The Evolution of the Cosmic SN Rate

Massimo Della Valle

Arcetri Astrophysical Observatory

Science cases (for OWL)

- Evolution of cosmological parameters
- Large scale structure
- Galaxy formation
- (Dark?) matter distribution
- SNe at high redshift
- Star formation history
- Proto Globular Clusters
- Primordial stars (Pop III)
- Re-ionization epoch

Roberto Gilmozzi, ESO

Nino Panagia, ESA/STScI

Jacqueline Bergeron, IAP

Piero Madau, UCSC

Jason Spyromilio, ESO

Philippe Dierickx, ESO

Science with ELTs

Science with OWL: a practical case

The cosmic SN rate up to z ~ 20

SNe as calibrated standard candles, (SNe-Ia, Phillips 1993) provide a direct measurement of q_0 at z>0.3 (Perlmutter et al. 1998, 1999; Riess et al. 1999)

Science with ELTs

Science with OWL: a practical case

The cosmic SN rate up to z ~ 20

SNe as calibrated standard candles (Phillips 1993, SNe-Ia) provide a direct measurement of q_o at z>0.3 (Perlmutter et al. 1998, 1999, Riess et al. 1999)

The evolution of the cosmic SN rate provides a <u>direct</u> measurement of the cosmic star formation rate.

SN Rate as Tracer of Star Formation Rate

a) The rate of CC SNe (II & Ib/c) is a direct measurement of the death rate of stars $M>8\ M_\odot$ (e.g. Iben & Renzini 1983)

$>40\ M_\odot$? Normal II SNe? Normal or Peculiar Ic/b? Collapsars?

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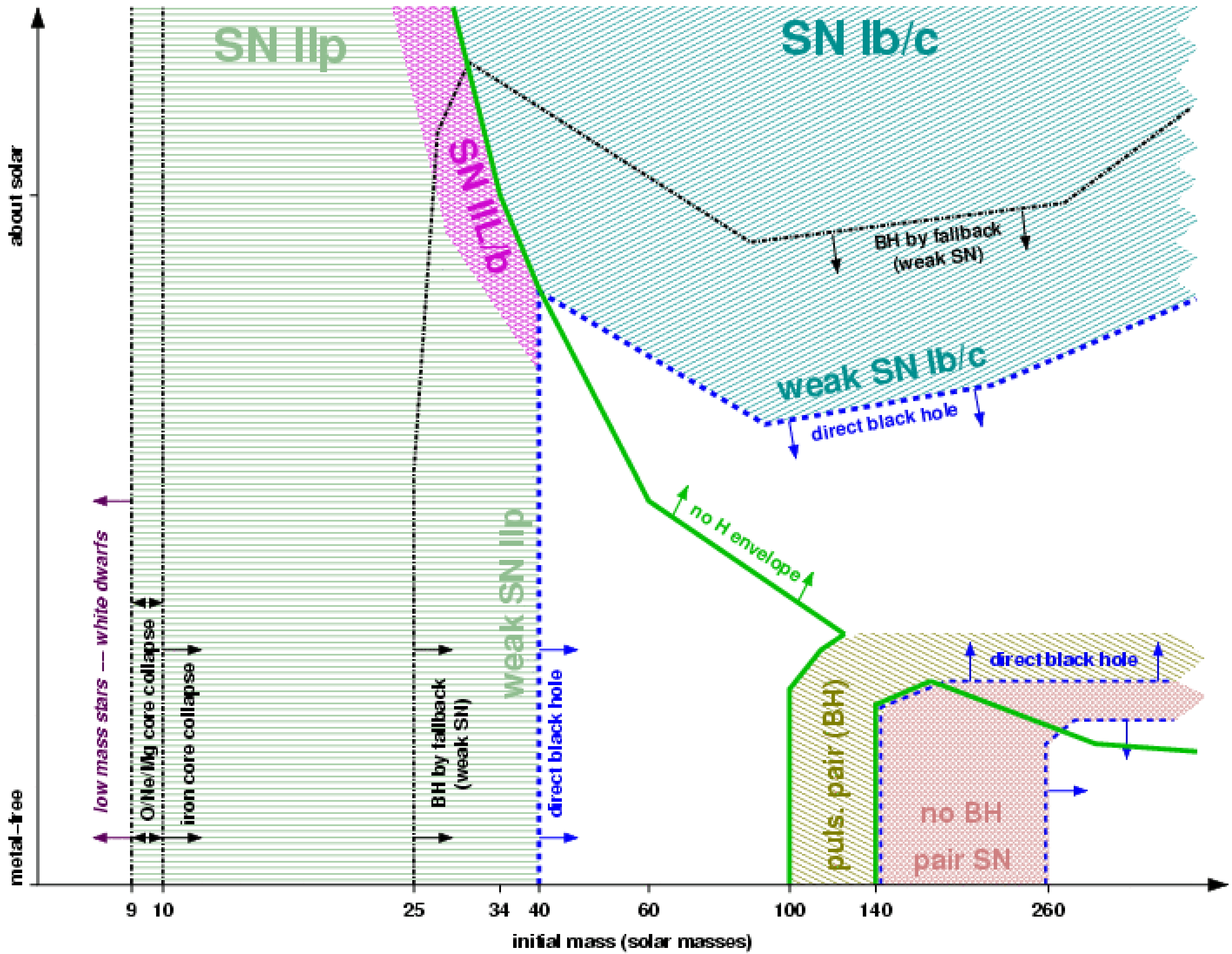

Heger et al. 2001

SN Rate as Tracer of Star Formation Rate

a) The rate of CC SNe (II & Ib/c) is a direct measurement of the death rate of stars $M > 8\ M_\odot$ ($>40\ M_\odot$? Normal II SNe? Normal or Peculiar Ic/b? Collapsars?)

b) The rate of type Ia SNe provide the history of star formation of moderate mass stars, 3-8 $M_\odot$

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SN Rate as Tracer of Star Formation Rate

a) The rate of CC SNe (II & Ib/c) is a direct measurement of the death rate of stars M>8 $M_\odot$ (>40 $M_\odot$? Normal II SNe? Normal or Peculiar Ic/b? Collapsars?)

b) The rate of type Ia SNe provide the history of star formation of moderate mass stars, 3-8 $M_\odot$

c) The evolution of the rate can clarify the nature of the progenitors of type Ia SNe (WD+WD or WD+MS) (Madau, Della Valle & Panagia 1998)

Science with OWL: a practical case

The cosmic SN rate up to z ~ 20

SNe as calibrated standard candles (Phillips 1993, SNe-Ia) provide a direct measurement of q_o at z>0.3 (Perlmutter et al. 1998, 1999, Riess et al. 1999)

The evolution of the cosmic SN rate provides a direct measurement of the cosmic star formation rate.

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SN rate→ Star formation Rate

For a Salpeter IMF with M_{up}=100 $M_\odot$
50% of all SN II are produced by stars with 8< $M_\odot$ <13.1 and 50% of the mass produced in SNe is in the interval 8-21.6 $M_\odot$

A 13 $M_\odot$ MS star has L=8000 $L_\odot$ and T_{eff}=22000 K

A 21.6 $M_\odot$ L=35000 $L_\odot$ and T_{eff}=27000

SN rate→ Star formation Rate

More than 50% of the stars producing SNe are poor sources of ionizing photons and of UV continuum photons

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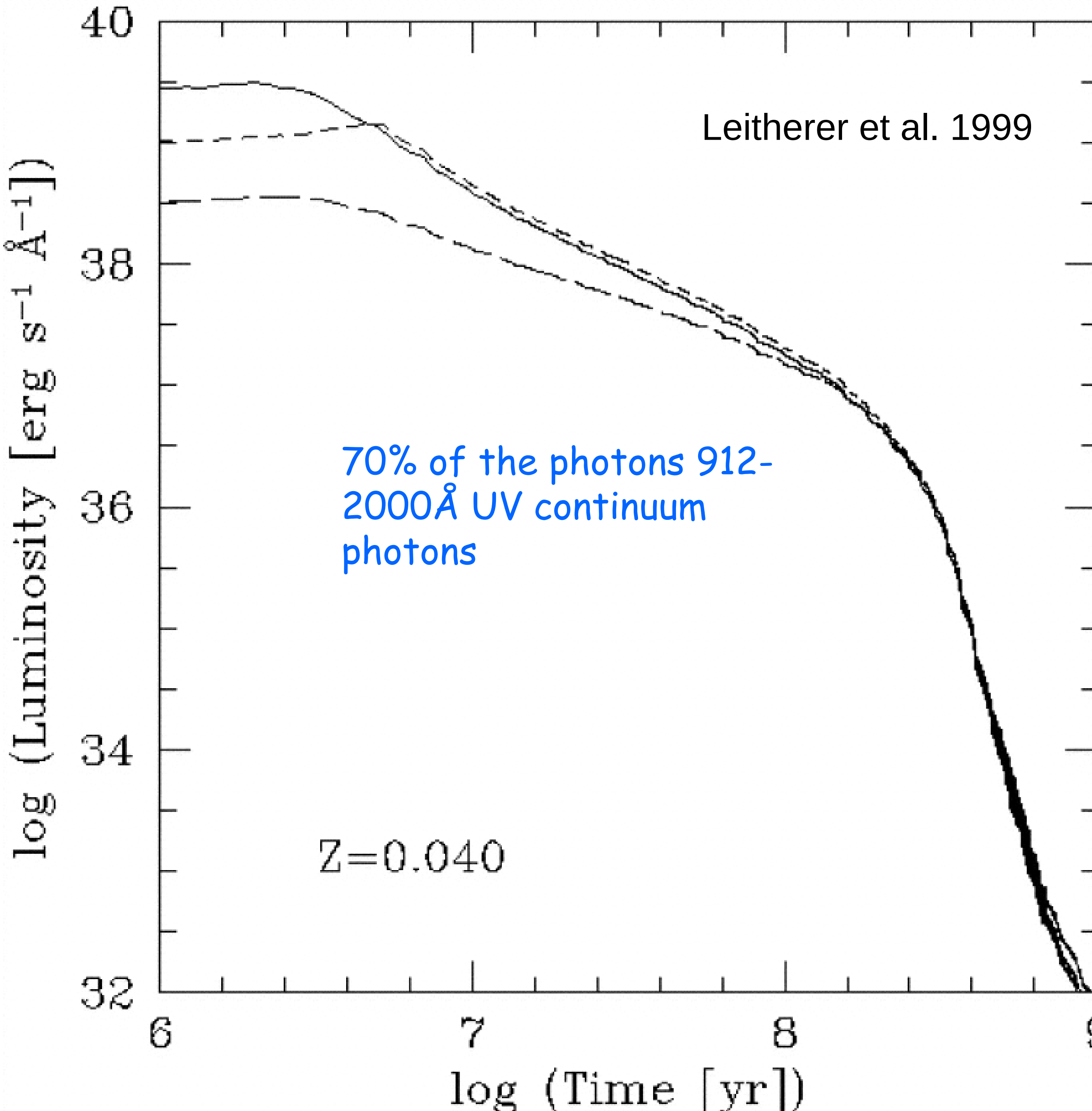

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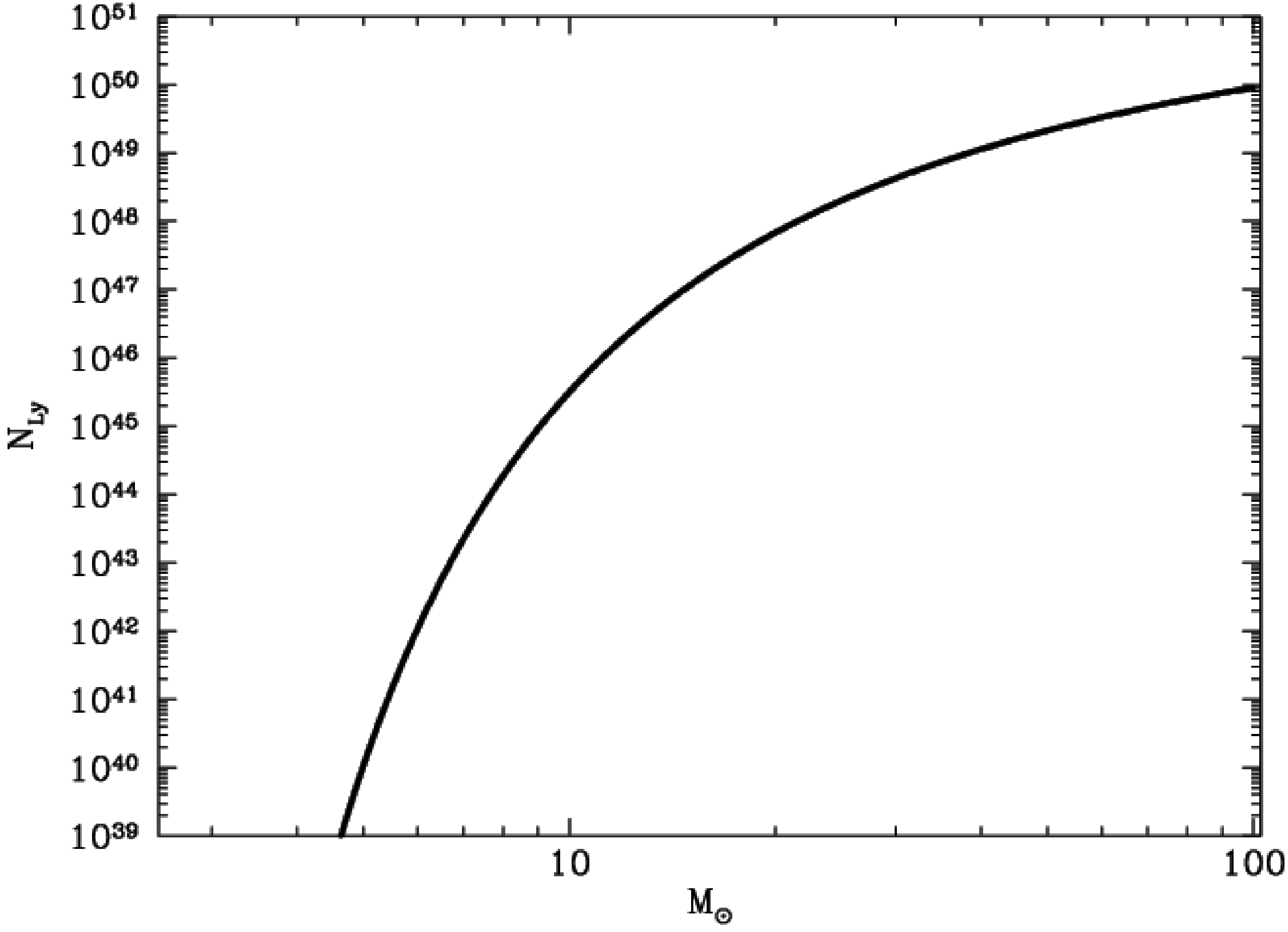
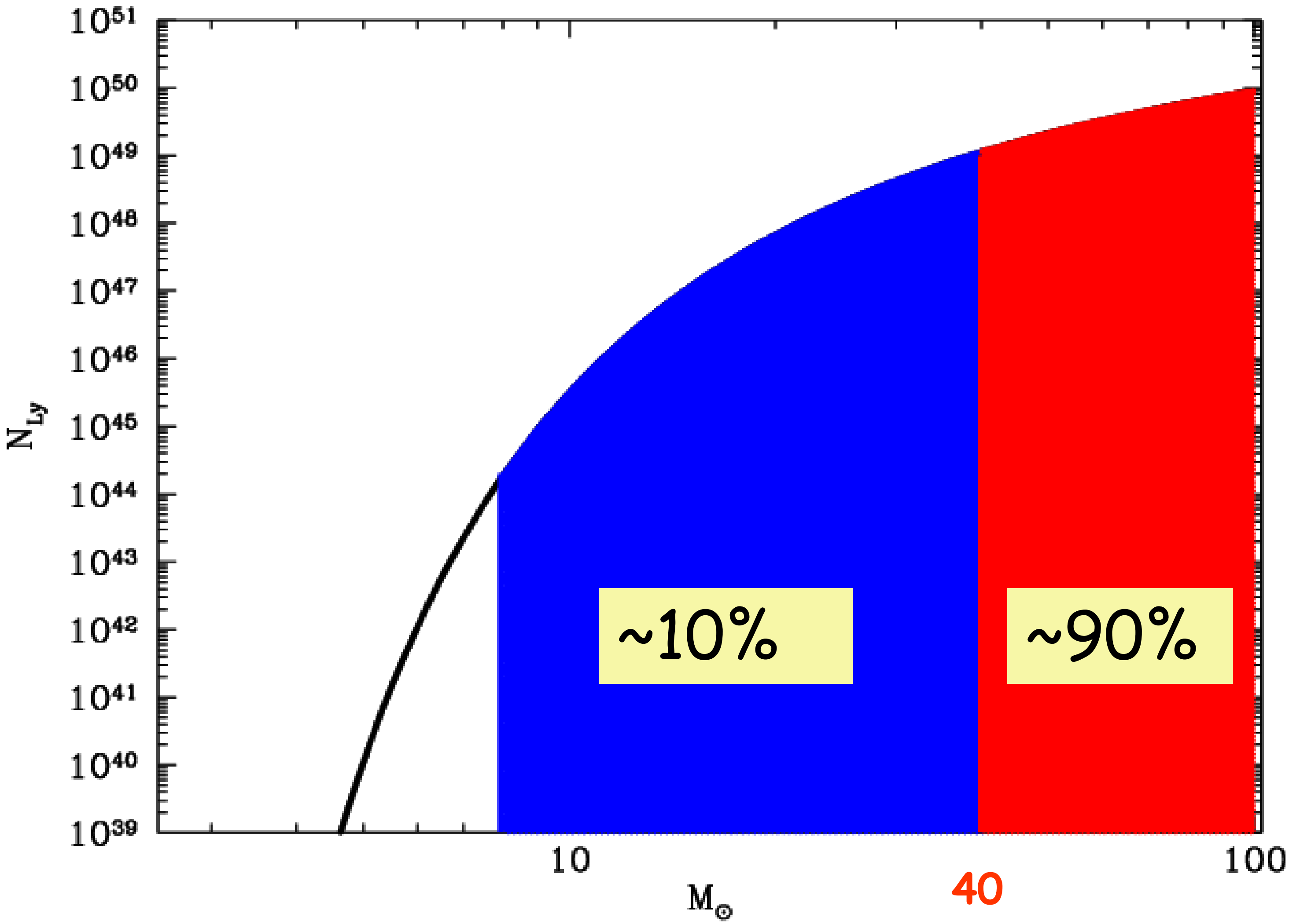




10^{51}
10^{50}
10^{49}
10^{48}
10^{47}
10^{46}
10^{45}
10^{44}
10^{43}
10^{42}
10^{41}
10^{40}
10^{39}
N_{Ly}
~10%
~90%
10
40
100
$M_\odot$

The bulk of the UV radiation both in the Balmer and in the Lyman continuum is produced by stars more massive than 40 $M_\odot$

SN rate→ Star formation Rate

Both the Halpha fluxes and the UV fluxes measure only the very upper part of the IMF [>40 $M_\odot$→ 8% ($M_\odot$>8)]

therefore:

SN rate→ Star formation Rate

Both the Halpha fluxes and the UV fluxes measure only the very upper part of the IMF [>40 $M_\odot$→ 8% ($M_\odot$>8)] therefore:

They are NOT good star formation rate indicators because

a) they require a huge extrapolation to lower masses

b) the extrapolation depends on the value of Mup which is not well known and may be not a constant quantity in different environment (Bressan et al. 2002) or at different z (Heger et al. 2002)

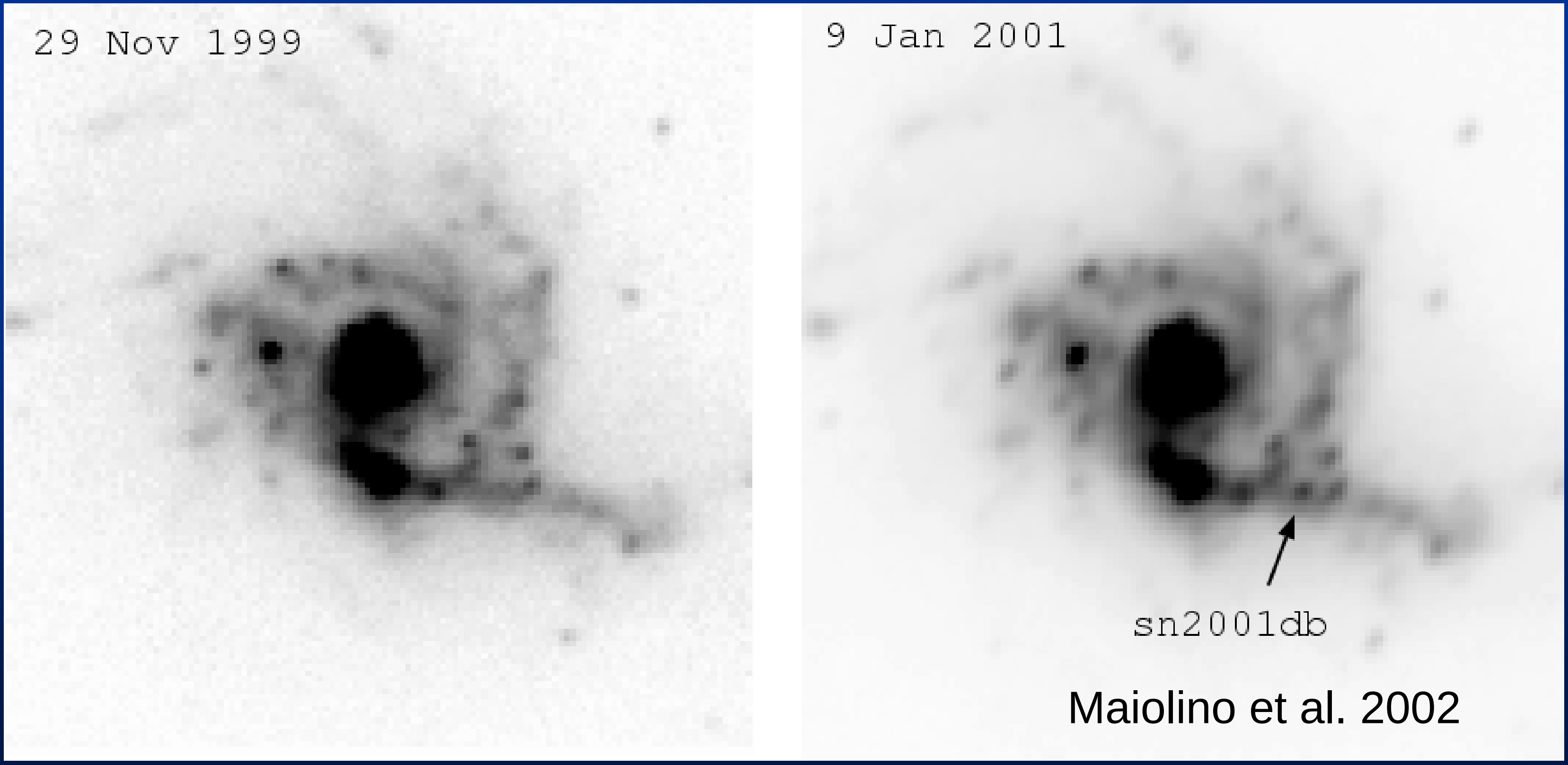
SN rate → Star formation rate

SNe provides a measurements of the star formation rate which is:

1. Independent of other possible determinations
2. More direct, because the IMF extrapolation is much smaller
3. More reliable because it is based on counting SN explosions rather than relying on identifying and measuring the sources of ionization (if using H-alpha flux) or the sources of UV continum

Drawbacks…..

Supernovae can be missed because of the extinction



Maiolino et al. 2002

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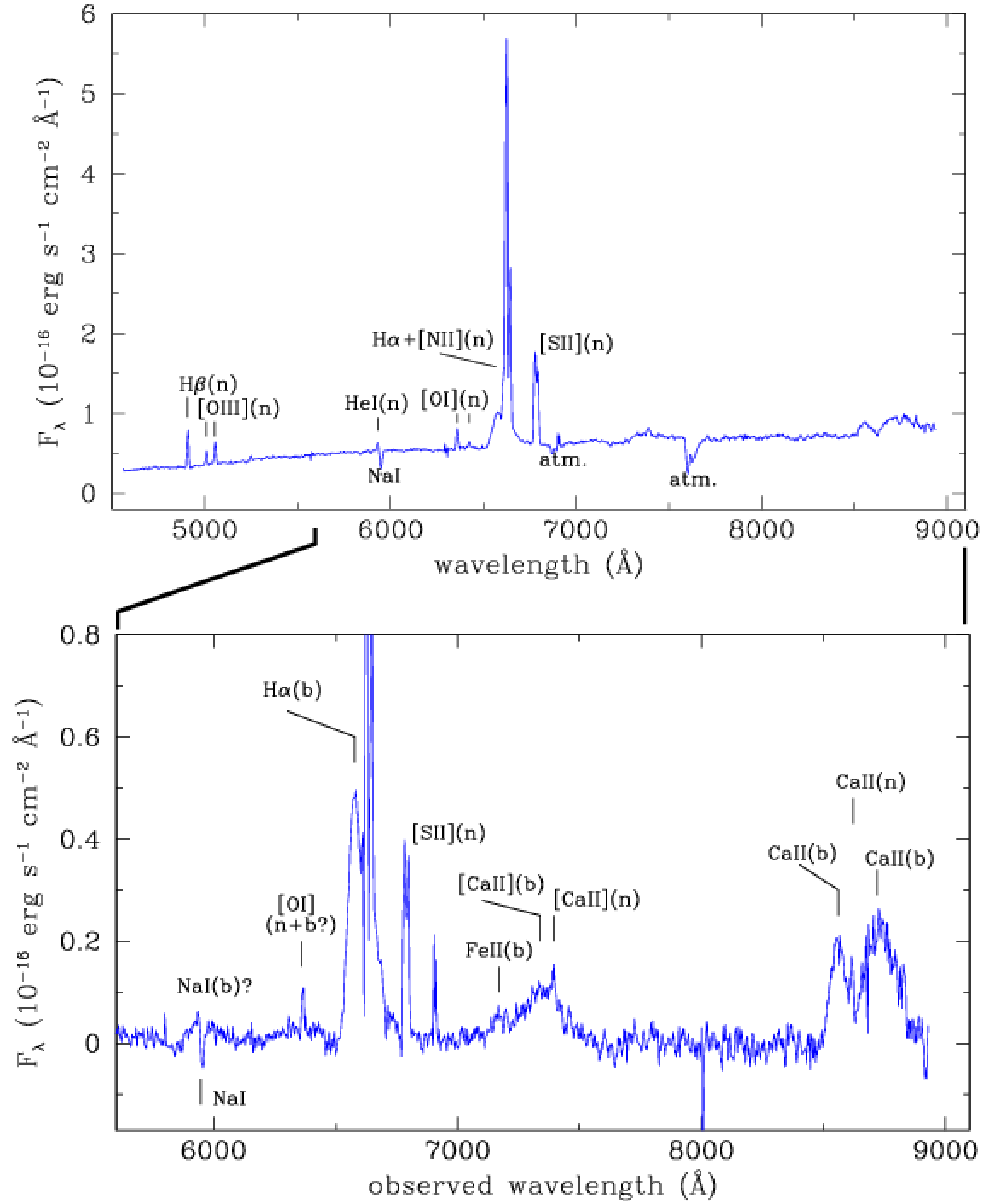

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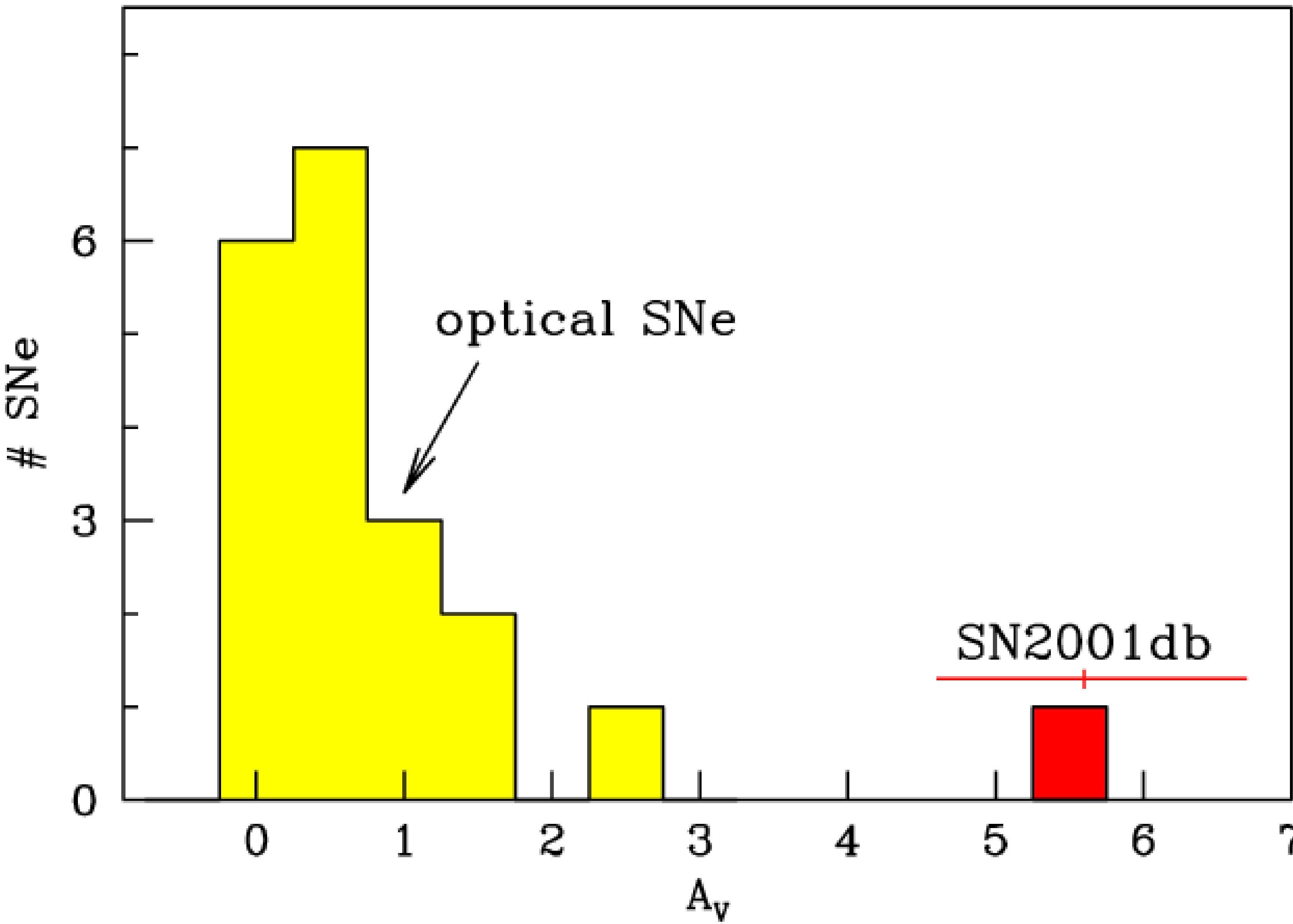

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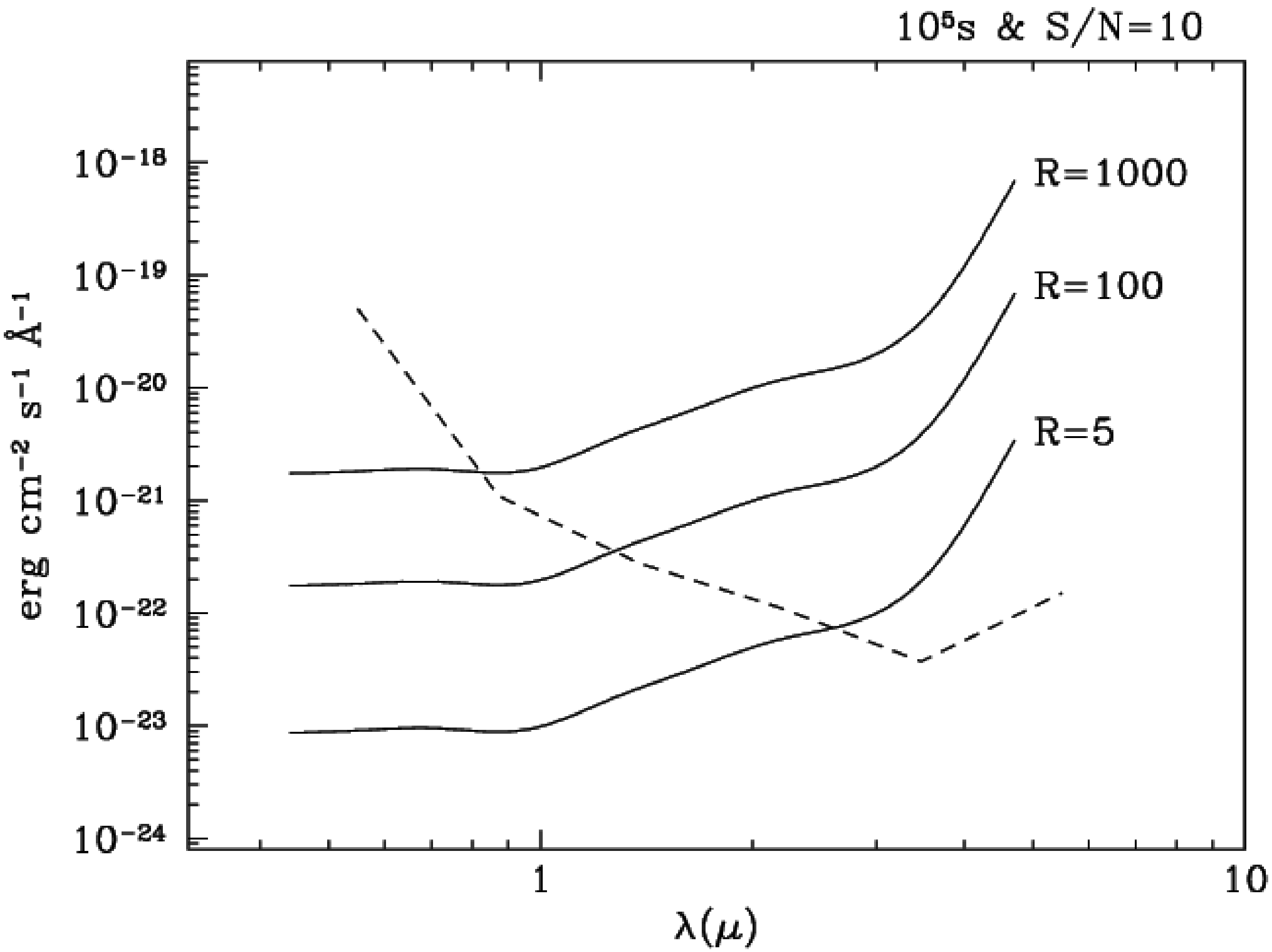
Ingredients of the simulation

1. OWL Performances: the imaging and spectroscopic limit of OWL as a function of S/N ratio and integration time
2. OWL Field: 2x2 arcmin (K)
3. Number of SNe expected in a single OWL Frame
4. SED of SNe (type II, Ia, Ib/c, Pop III)
5. Morphology of lightcurves and absolute magnitude at maximum of SNe (Control Time)
6. Distribution of SNe into the different spectroscopic types
7. Cosmology (Ω_m =0.3 Ω_Λ =0.7)

1.OWL Performances

http://www-astro.physisc.ox.ac.uk/imh/ELT

JWST: (Panagia 2003, Stiavelli 2003)



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2. & 3. Number of SNe in a single OWL Frame

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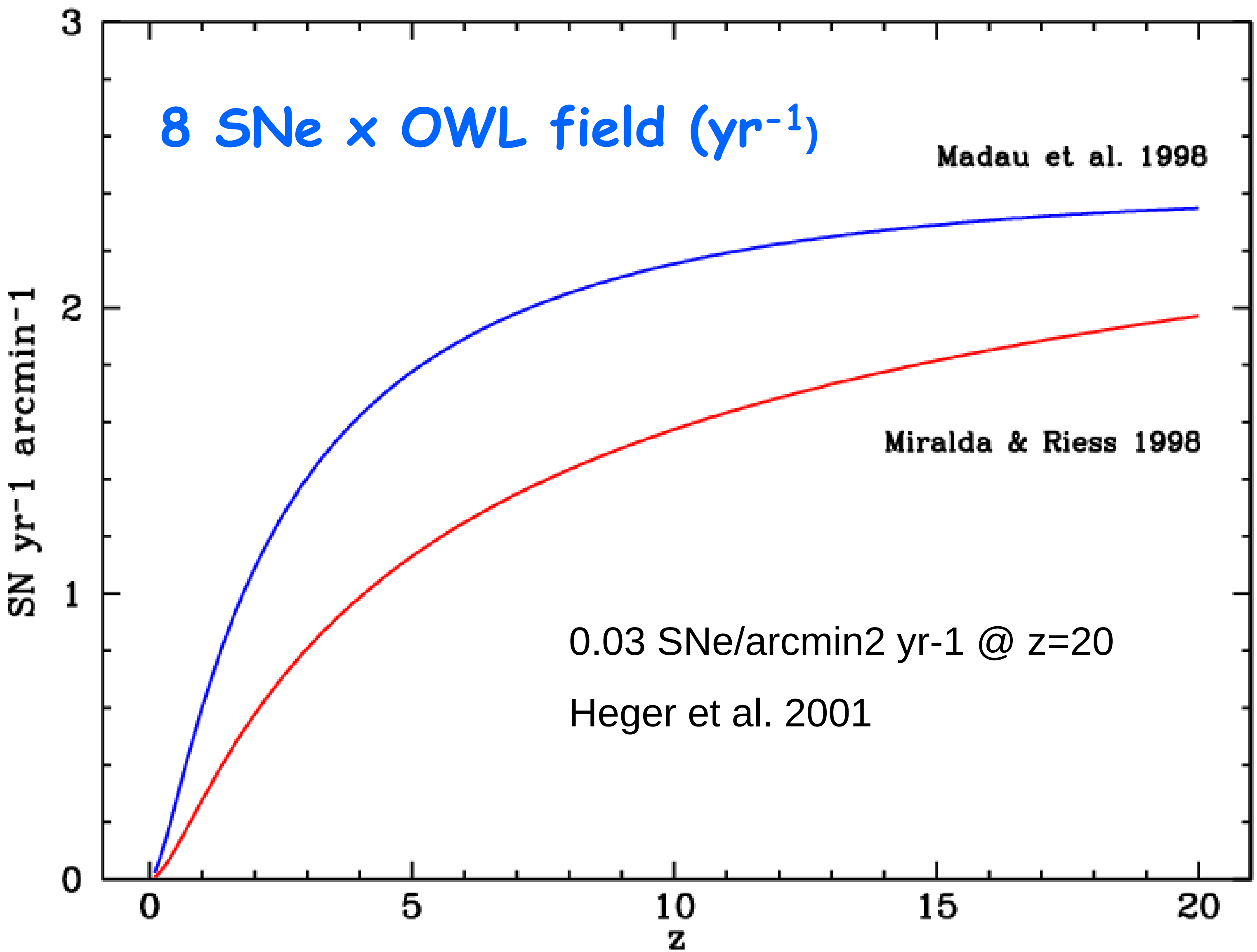
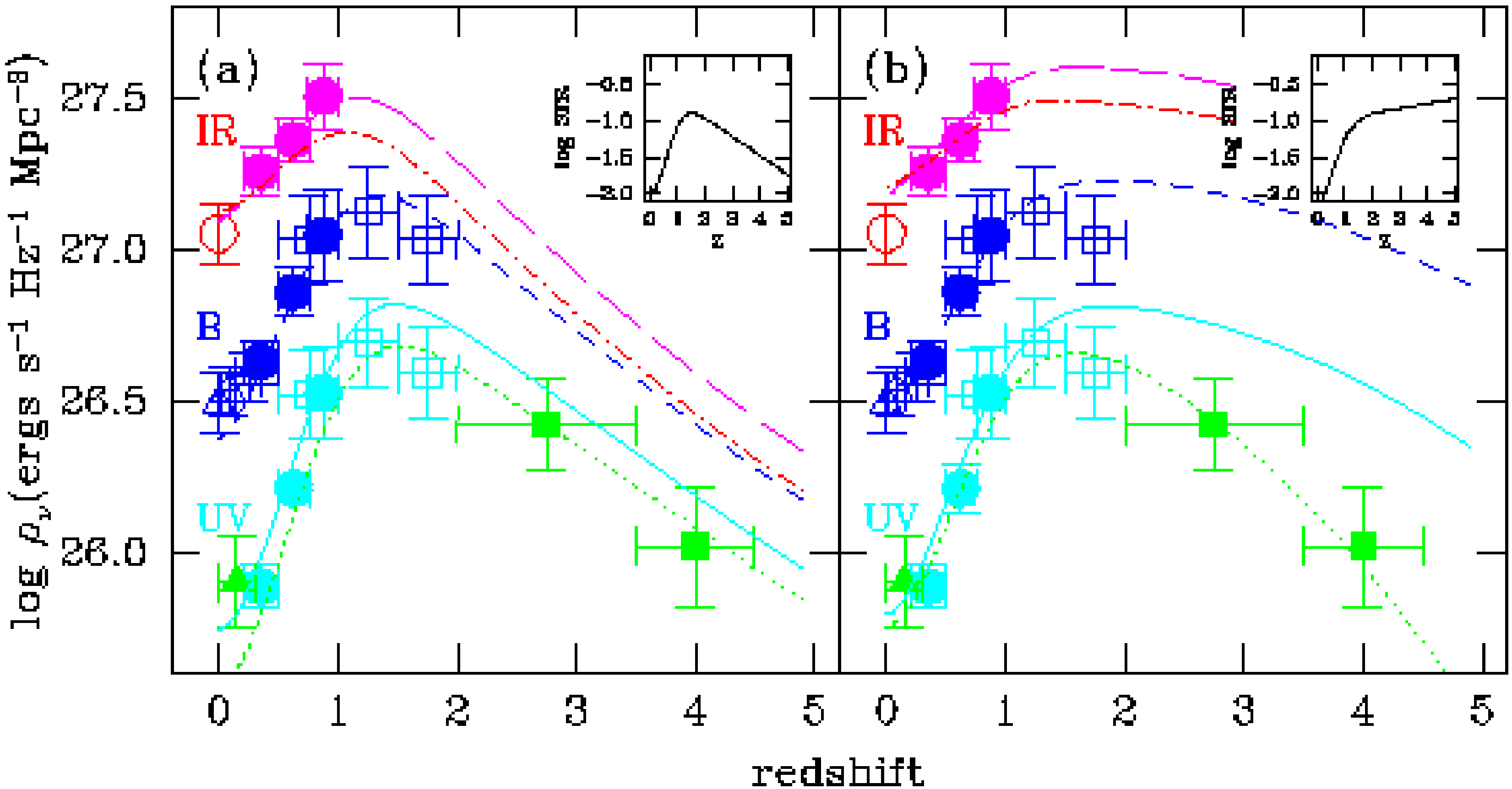
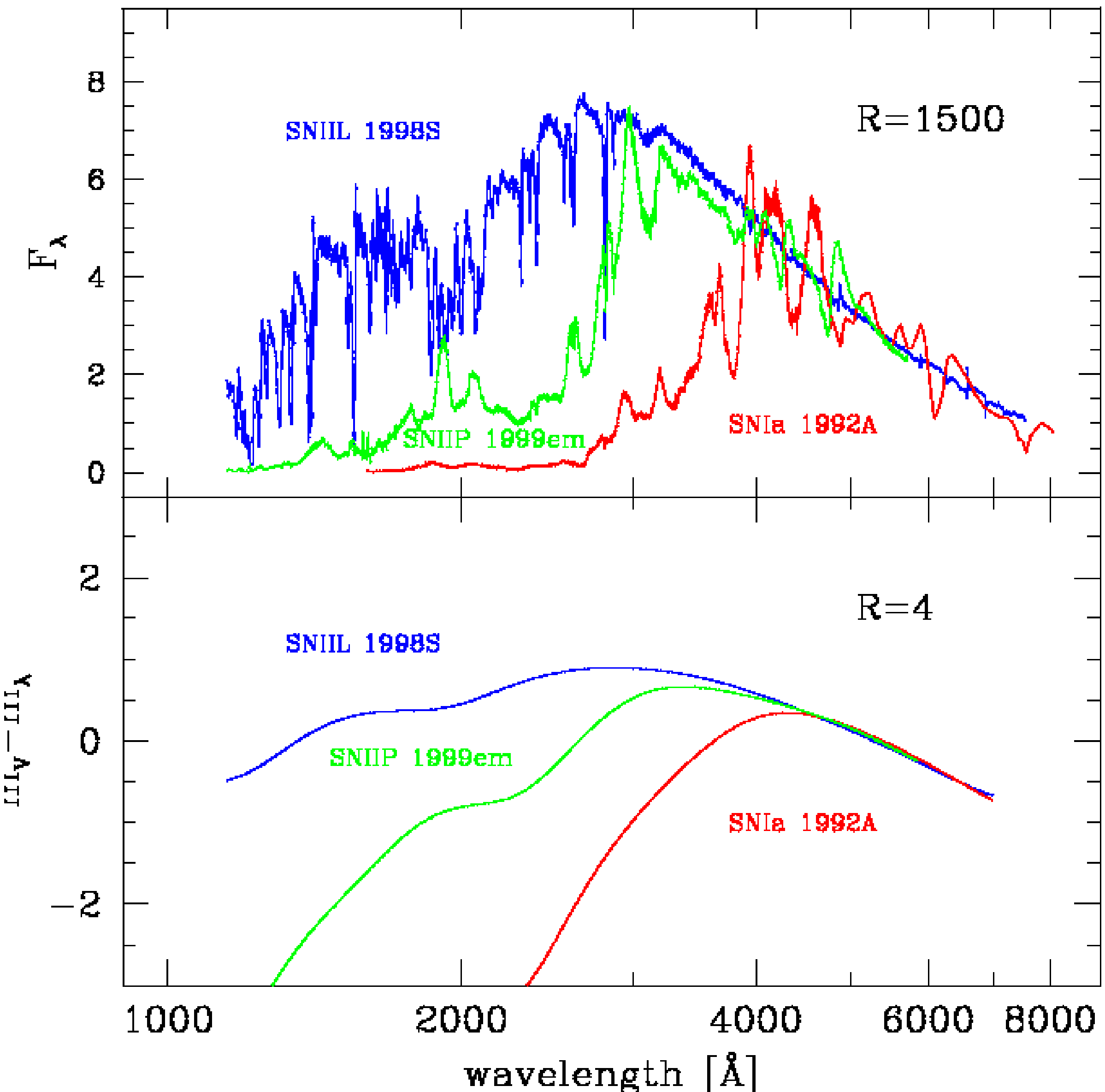




Madau, Della Valle & Panagia 1998

4. Spectral Distribution Energy

(Panagia 2003)



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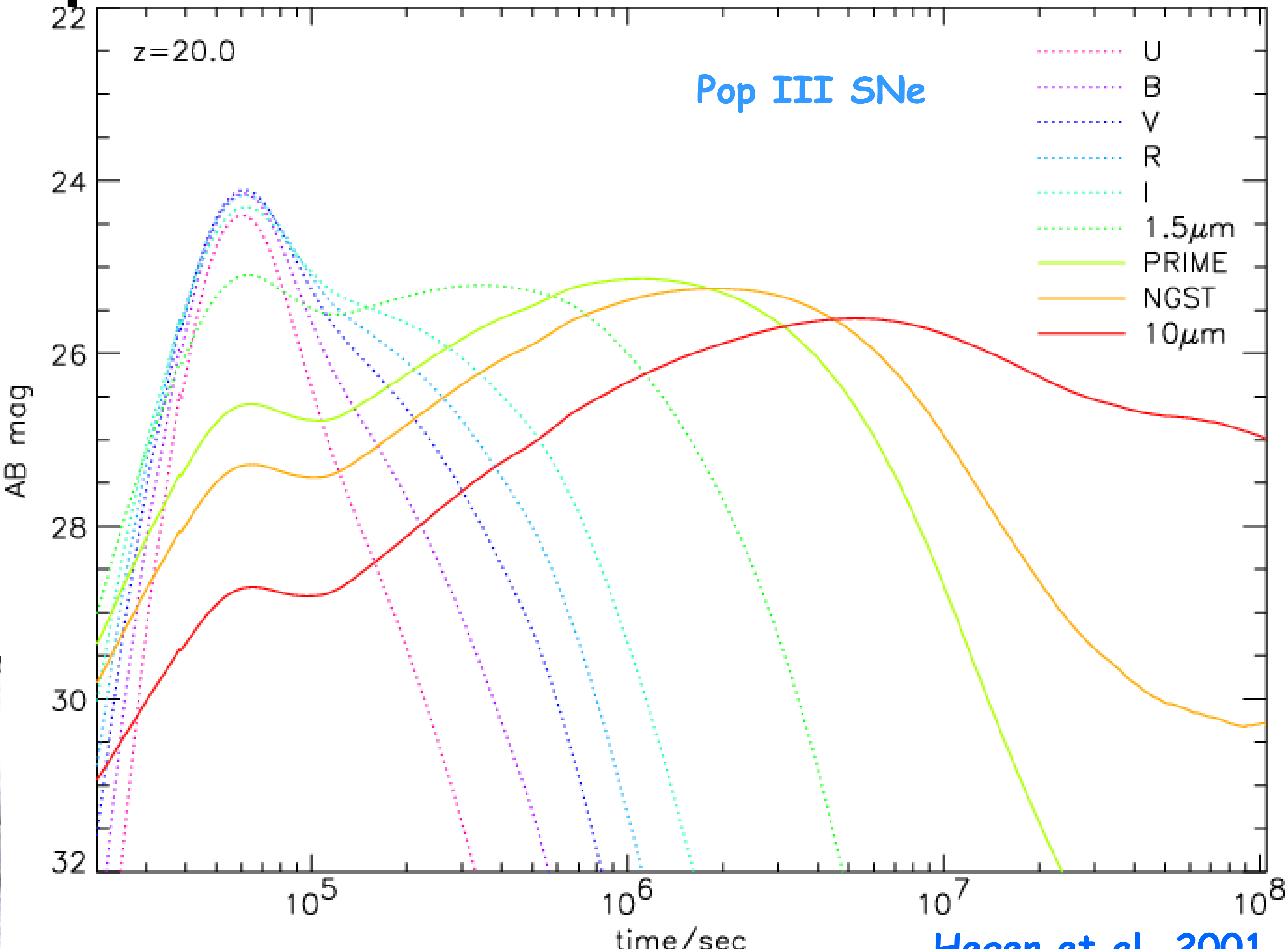
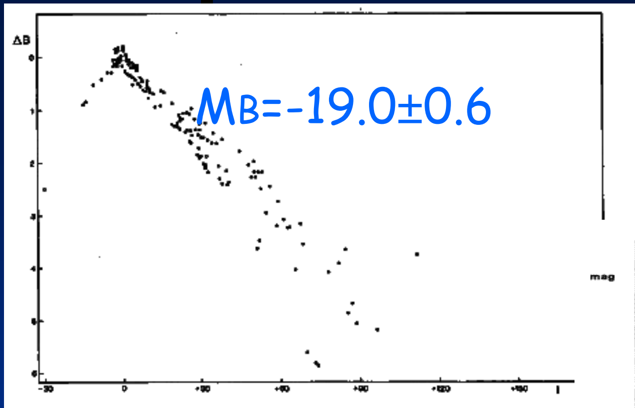


Heger et al. 2001



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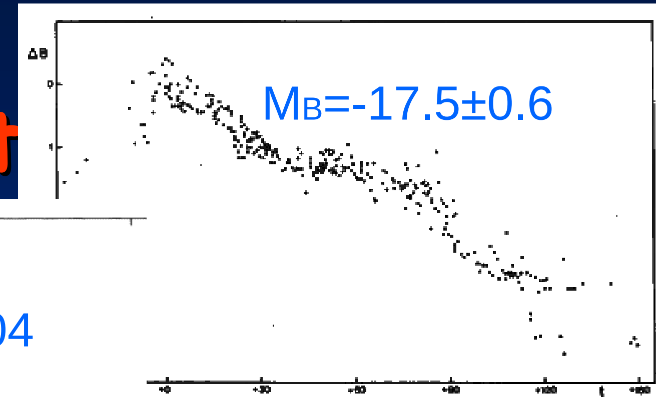

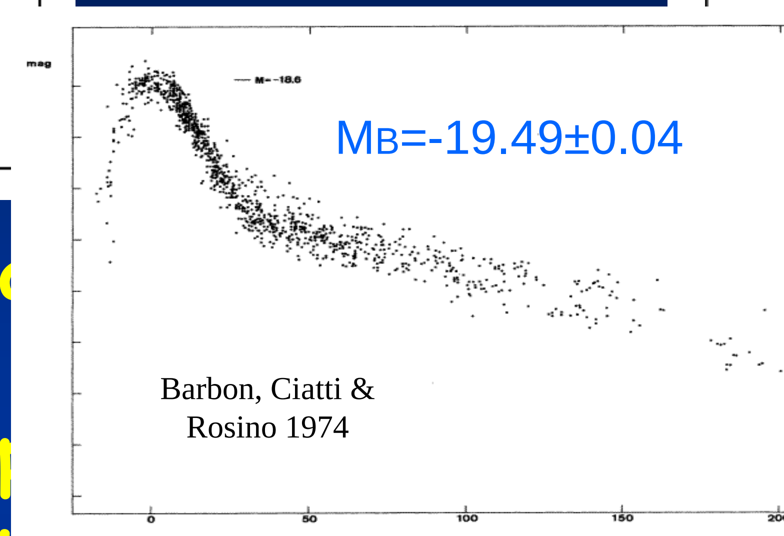

Science with

Observ... ...sed on the control ... ...vicky 1936). The typ... ... around maximum is 15-20 days and most SNe will occur at z<5 → light-curve width in the observer rest-frame about 100-120 days. 4 exposures at time intervals of 3 months will cover 1 year

(Saha et al. 2001; Patat et al. 1994)

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% of SNe in the different types

46% type II Normal (mostly plateau)

16% type II Bright (mostly linear)

20% type Ia

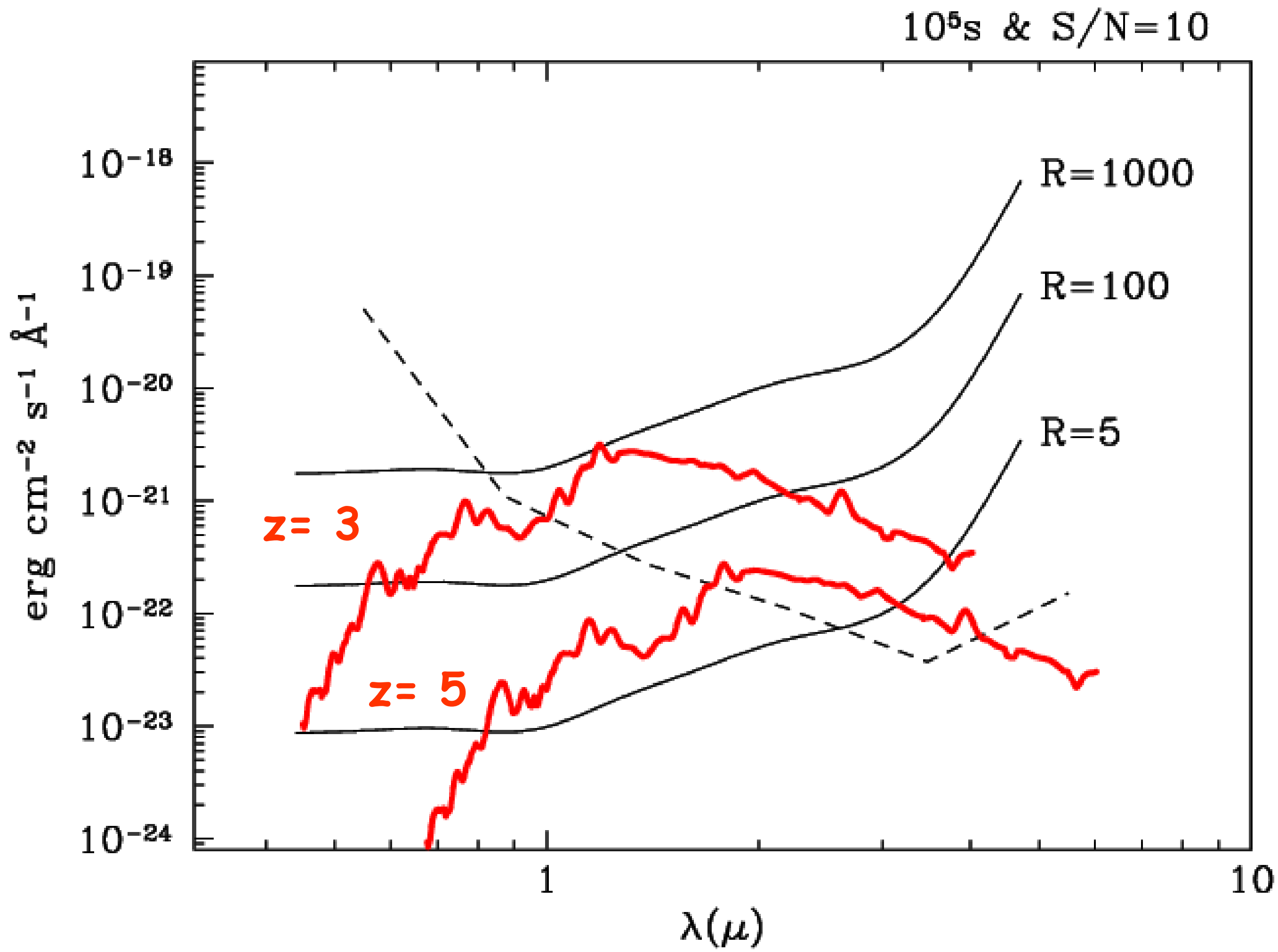
15% type Ib/c

0.5% hypernovae

1.5% SNe from Pop III

Mannucci et al. 2004; Della Valle et al. 2004; Heger et al. 2001; Podsiadlowski et al. 2004; Cappellaro et al. 1997

SNe-II Normal



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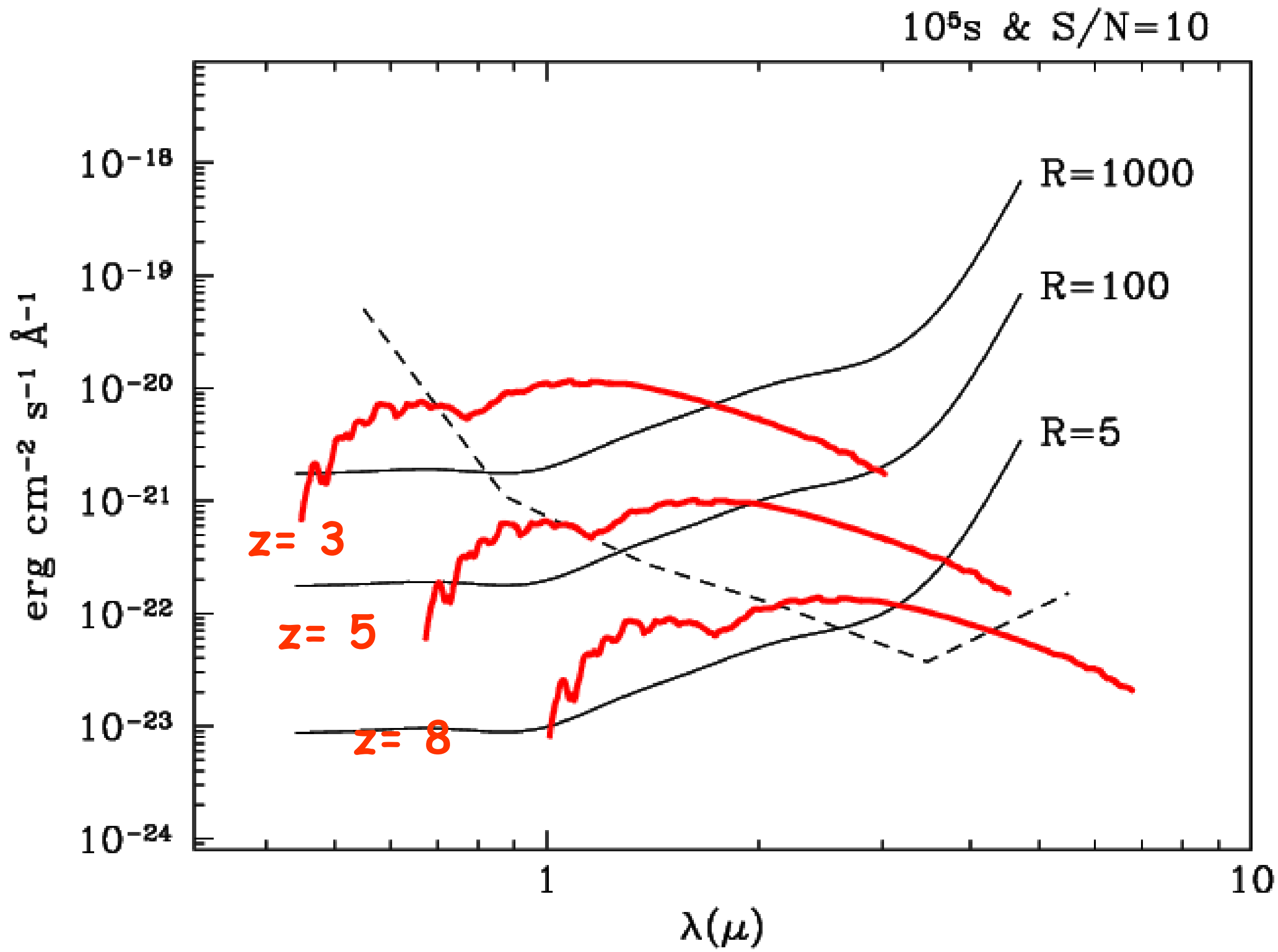
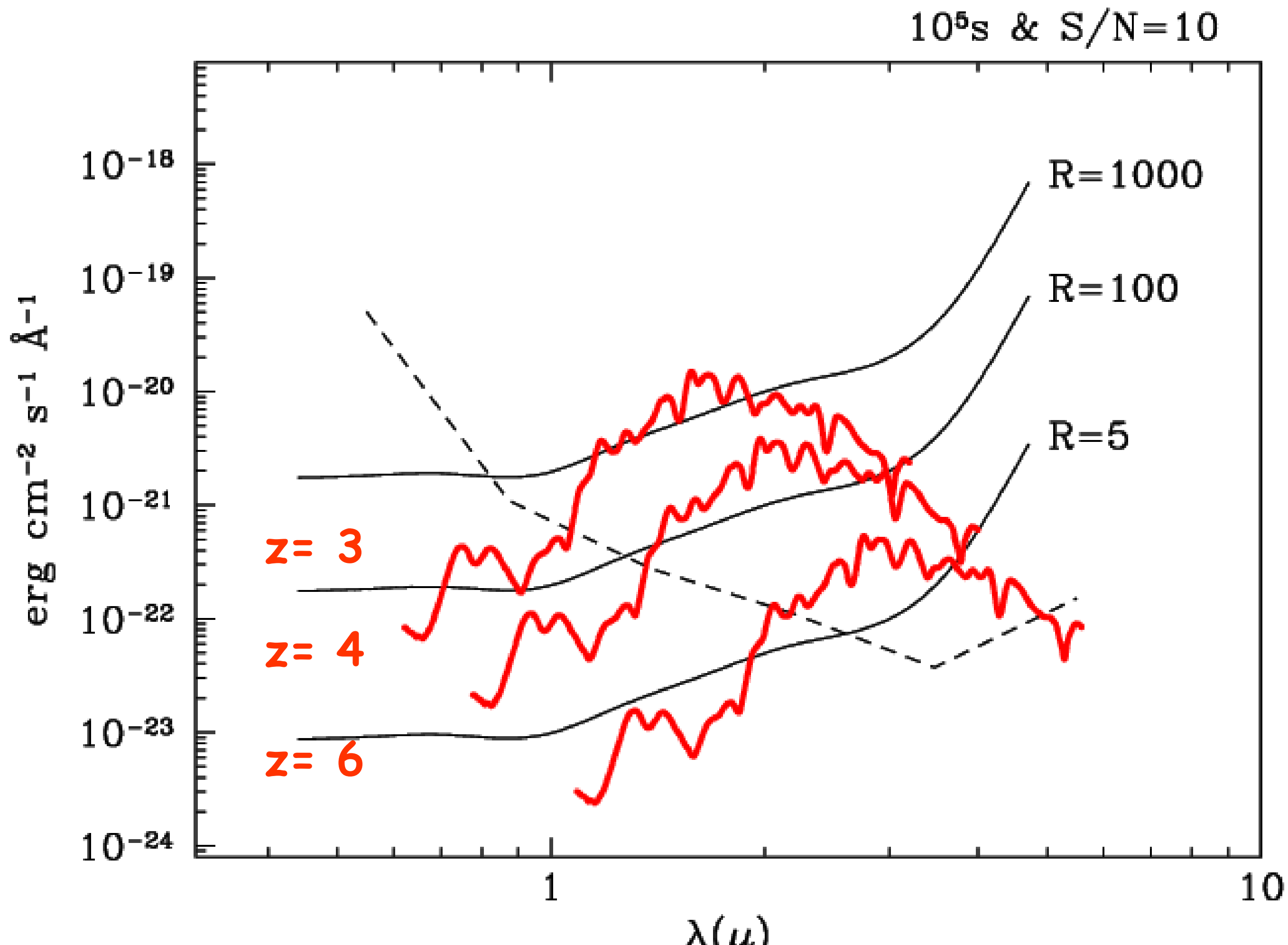
SNe II-Bright

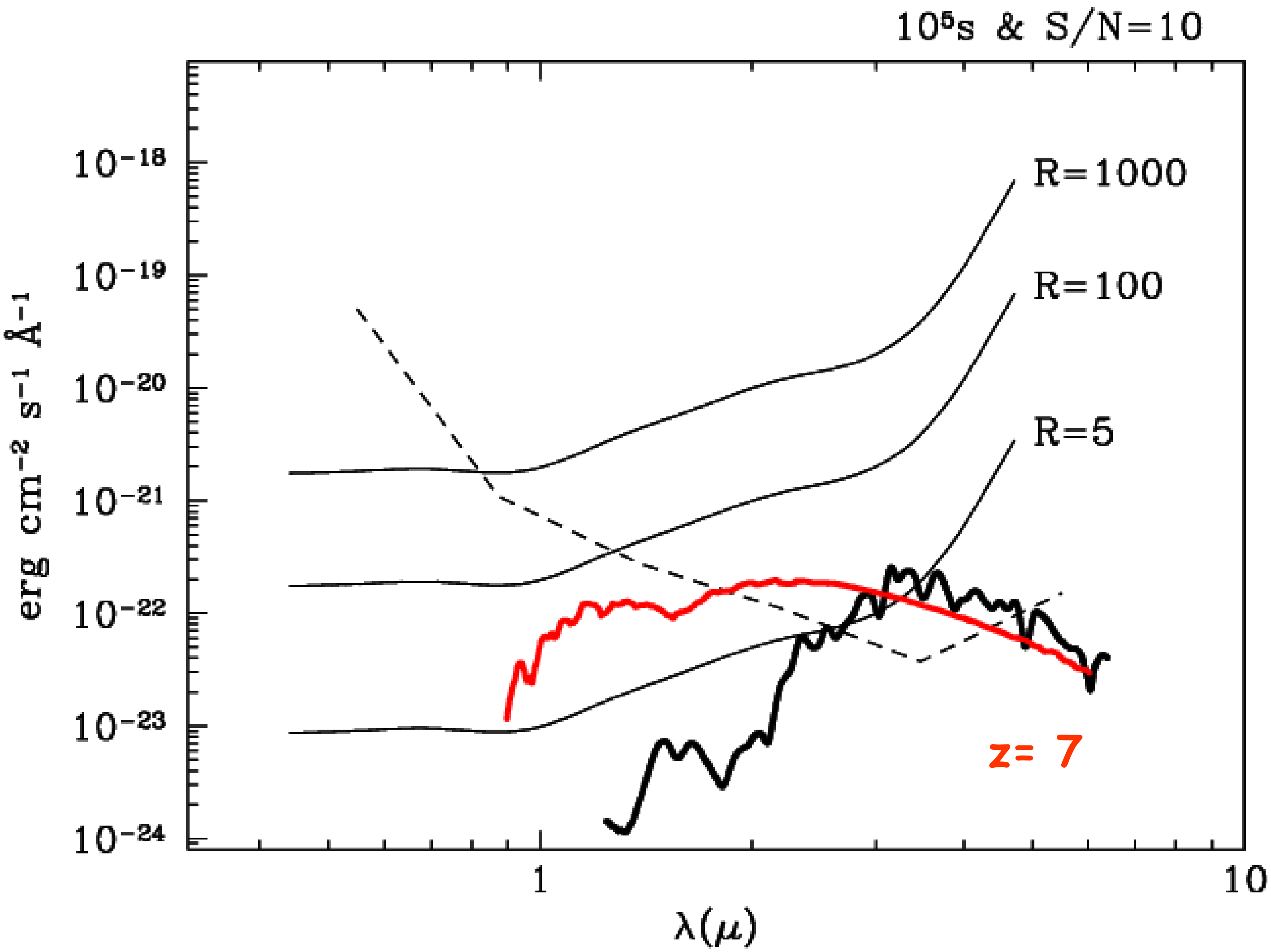


SNe Ia

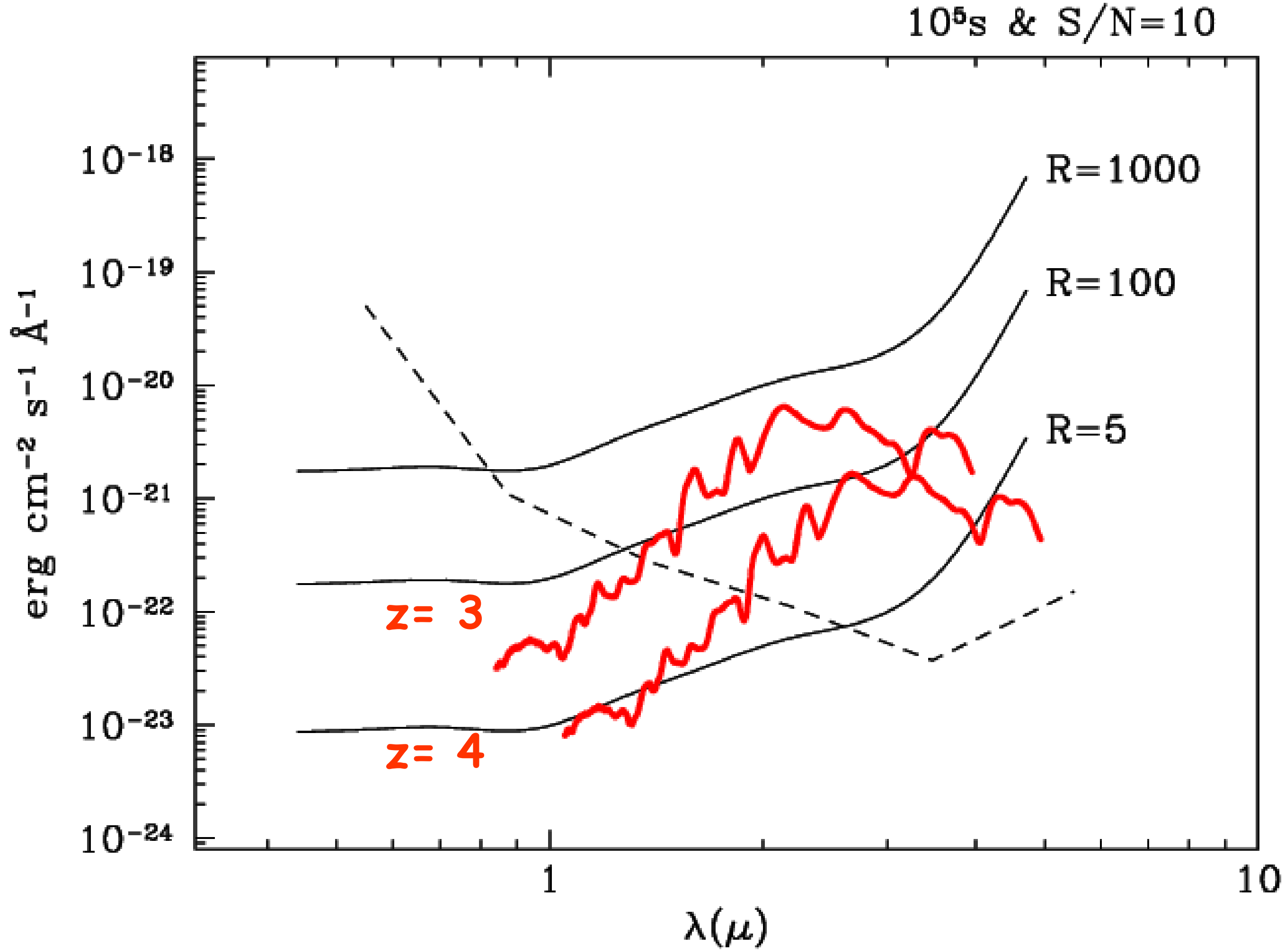


10^5s & S/N=10
10^{-18}
10^{-19}
10^{-20}
10^{-21}
10^{-22}
10^{-23}
10^{-24}
erg cm^{-2} s^{-1} Å^{-1}
R=1000
R=100
R=5
z= 3
z= 4
z= 6
1
10
$\lambda(\mu)$

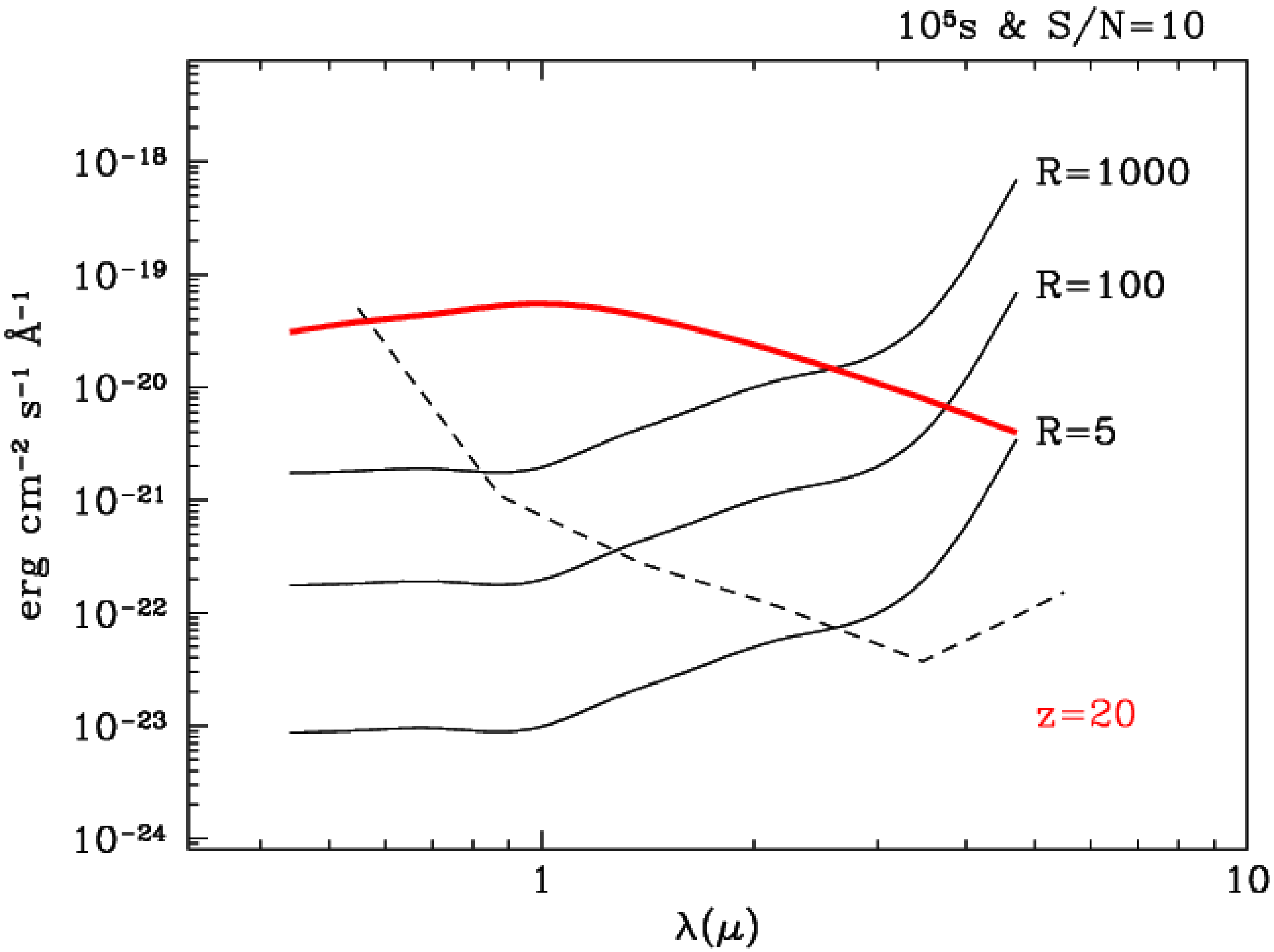
Science with ELTs



SNe Ib/c


10^5s & S/N=10
R=1000
R=100
R=5
z= 3
z= 4
10^{-18}
10^{-19}
10^{-20}
10^{-21}
10^{-22}
10^{-23}
10^{-24}
erg cm^{-2} s^{-1} Å^{-1}
1
10
$\lambda(\mu)$

Pop III SNe



10⁵s & S/N=10
R=1000
R=100
R=5
z=20
erg cm⁻² s⁻¹ Å⁻¹
10⁻¹⁸
10⁻¹⁹
10⁻²⁰
10⁻²¹
10⁻²²
10⁻²³
10⁻²⁴
1
10
λ(μ)

SN types vs. redshift

SNe Ia (3-8 $M_\odot$) visible up to z ~ 5 Blind below 2400A, K last useful band

SNe II (8-40 $M_\odot$) visible up to z ~ 8 Strong UV emitters (time-dilated UV flash)

Pop III SNe (100-300 $M_\odot$) visible to z ~20 (Heger & Woosley 2002)

The Simulation

- We plan to image 50 fields in the J, H and K bands (1h each) at 4 different epochs (="SN search") +
- 3 epochs in the K band for the photometric follow-up (i.e. seven K photometric points for each SN) +
- 4h for each SN (z< 4.5-5) to get the spectroscopic classification +
- 4h for each SN (z~20; 4 epochs 120h)
- Grand Total=600h (search)+150h (K follow-up)+200h(spectr. II & Ia)+120h (spectr. Pop III SNe)=1070h+10%

Results of the Simulation

1200h or 150 nights to study 400 SNe up to z =20

This is about twice as much the size of a current Treasury programme (450 orbits) and it is comparable with SNAP (now Join Dark Energy Mission), about 2000 SNe (Ia) in 2 yr (z<1.7).

1200h+200h (II epoch spectroscopy)

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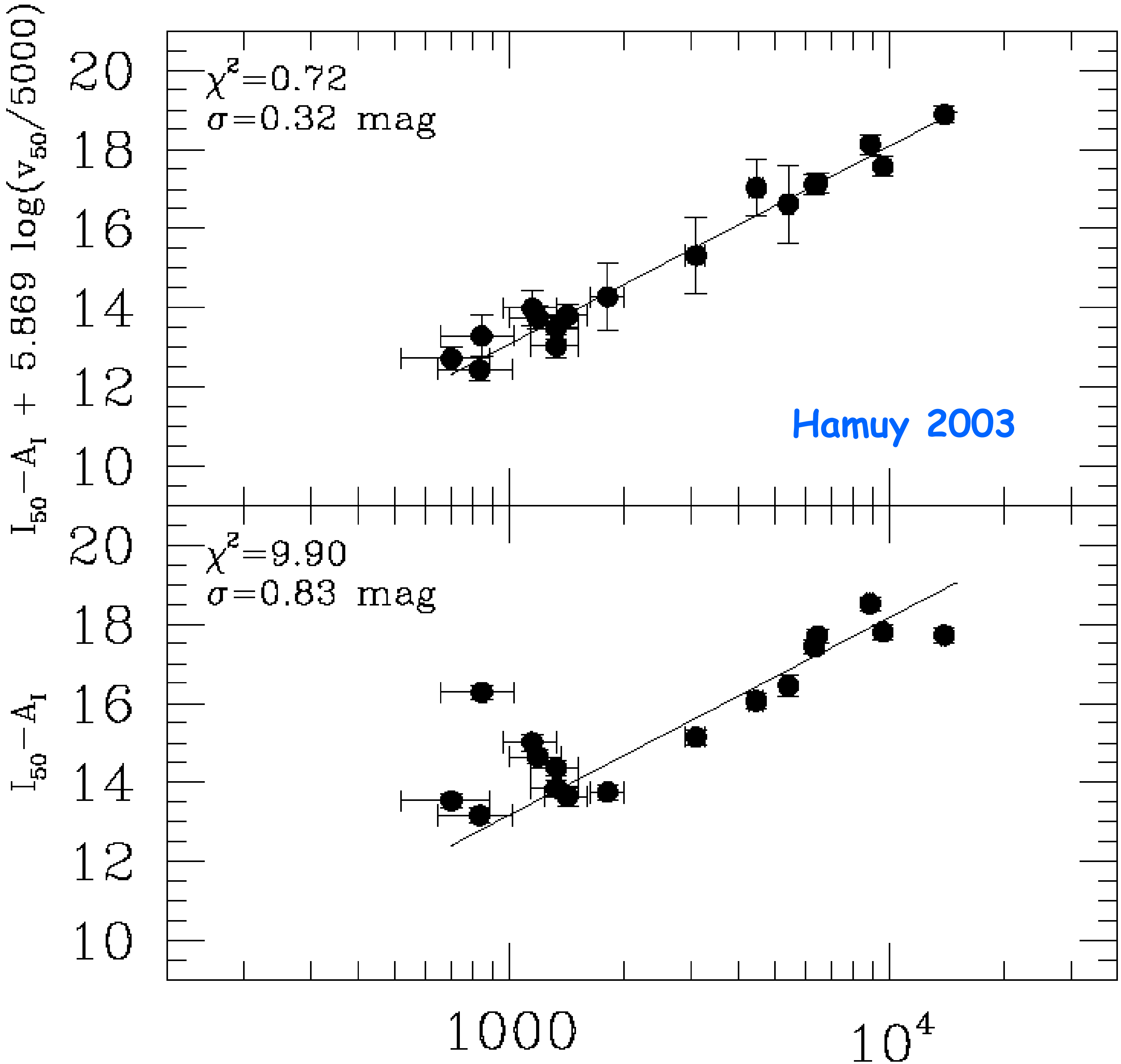

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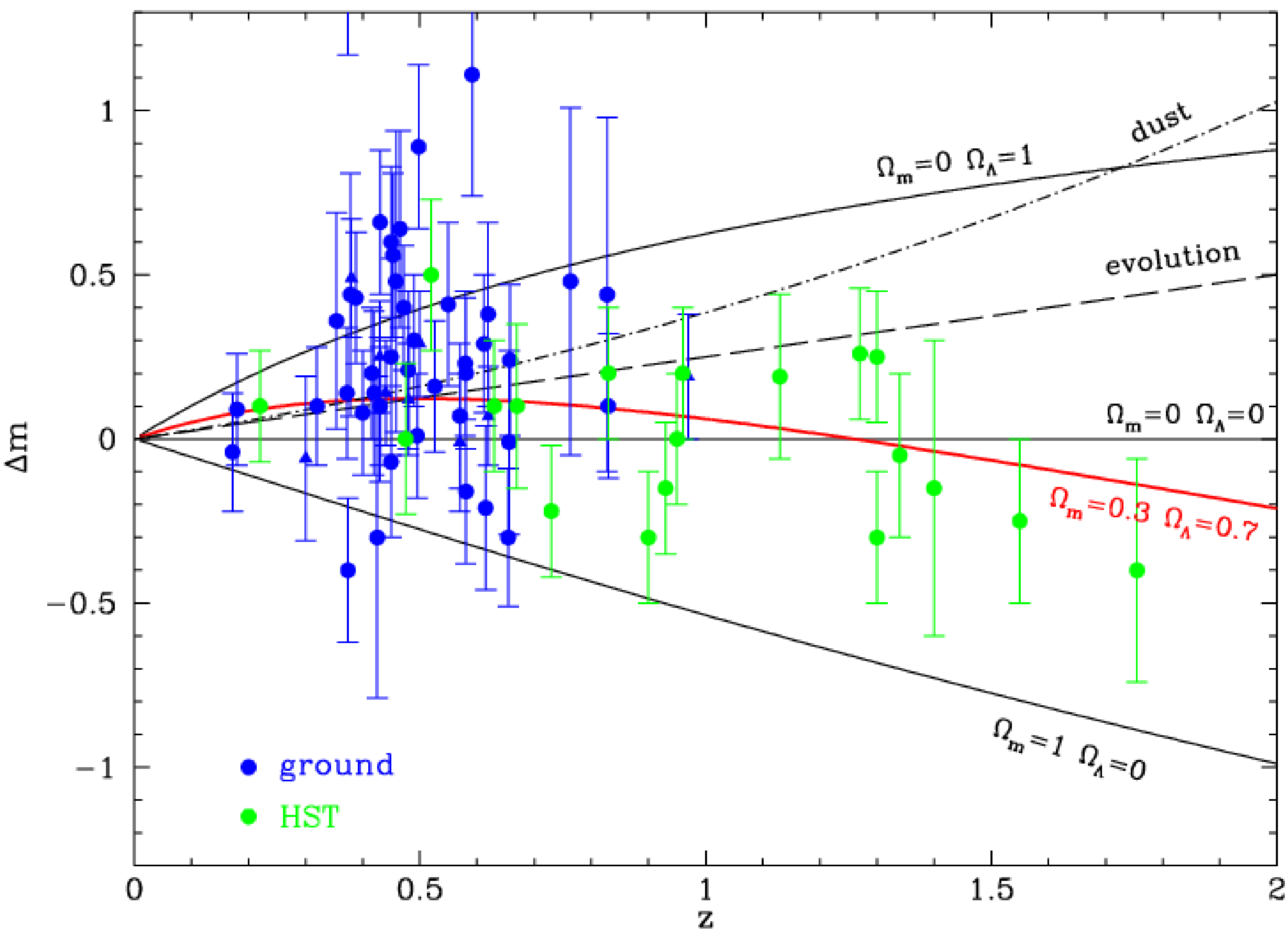

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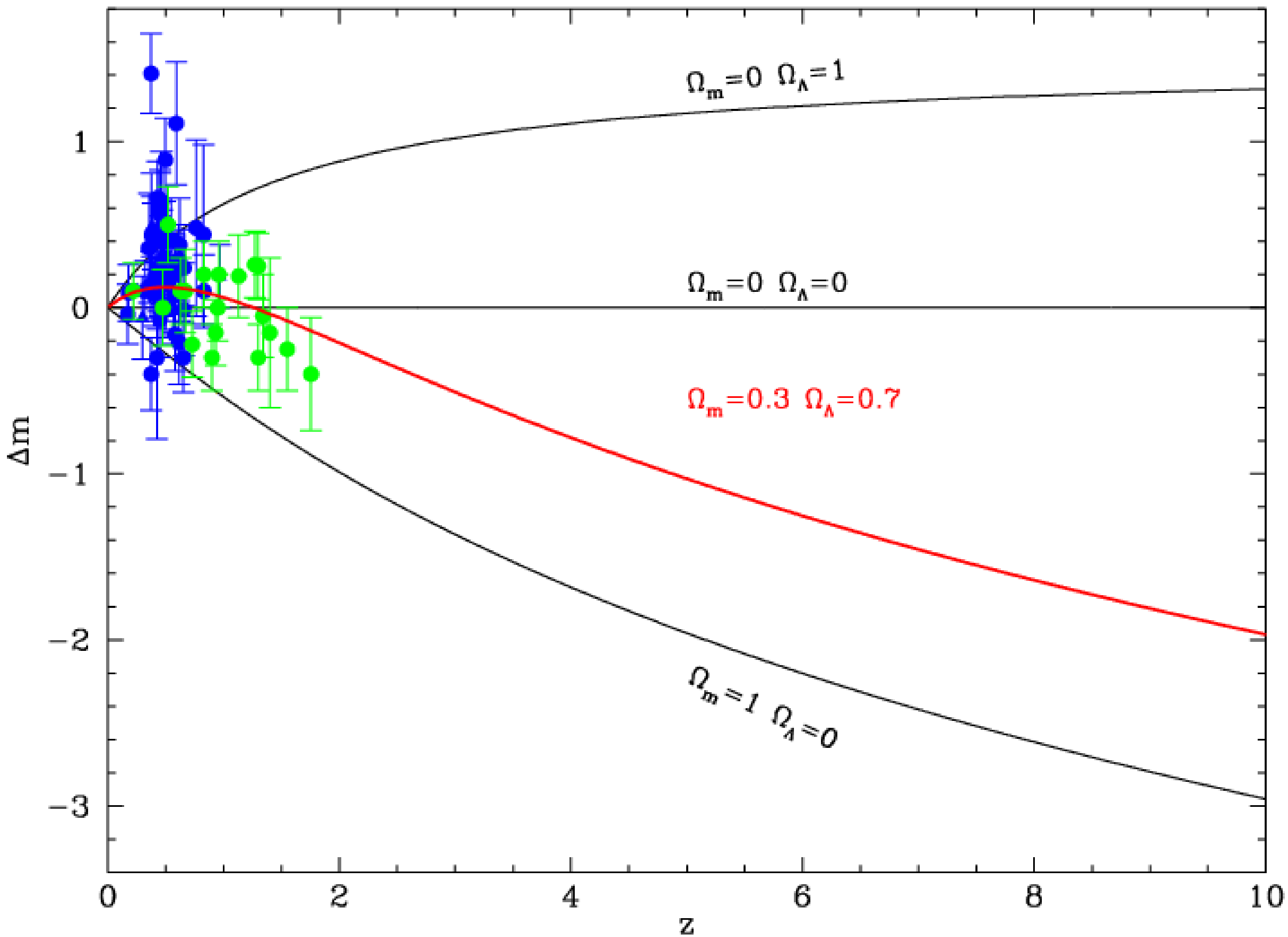

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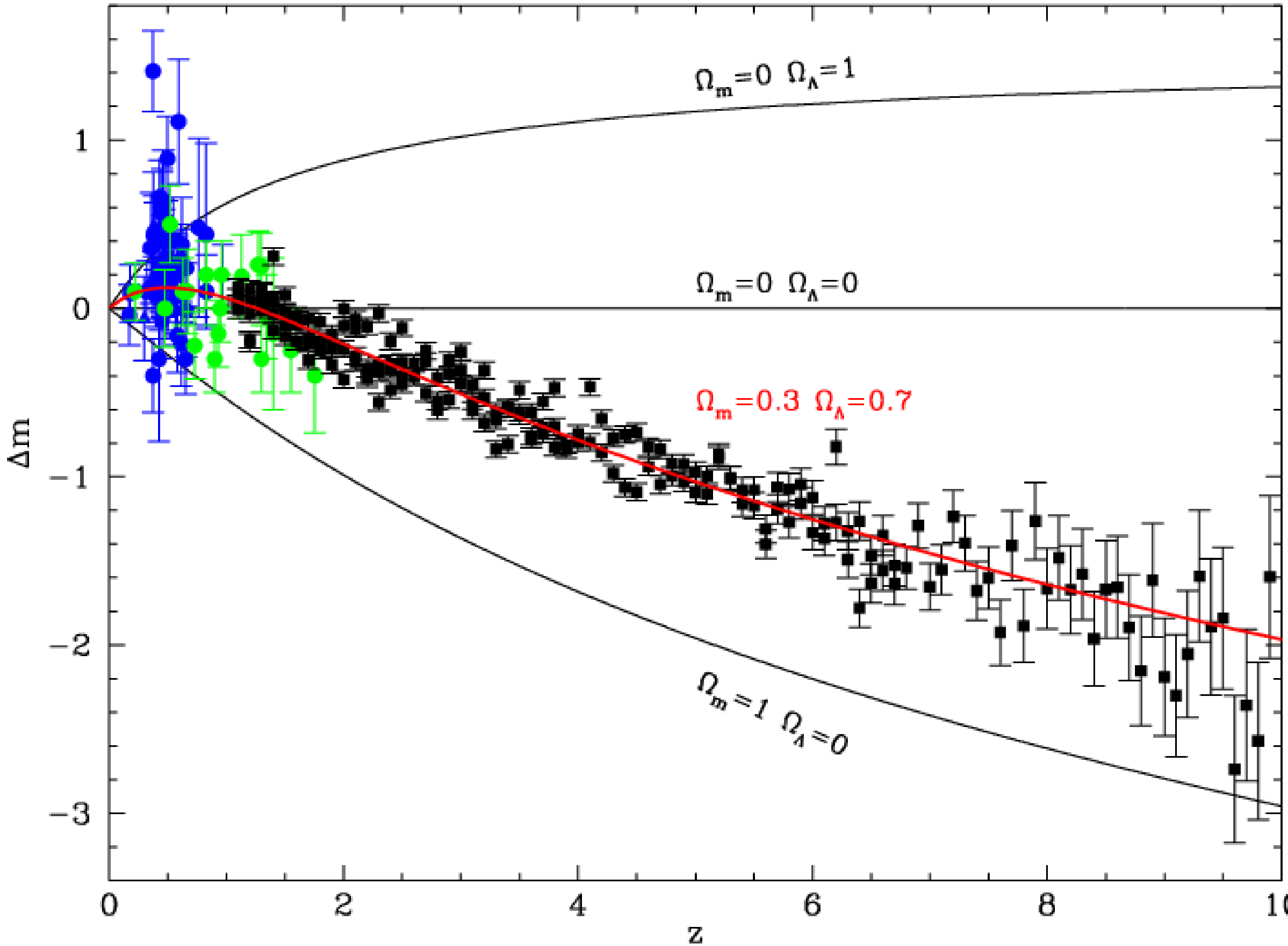

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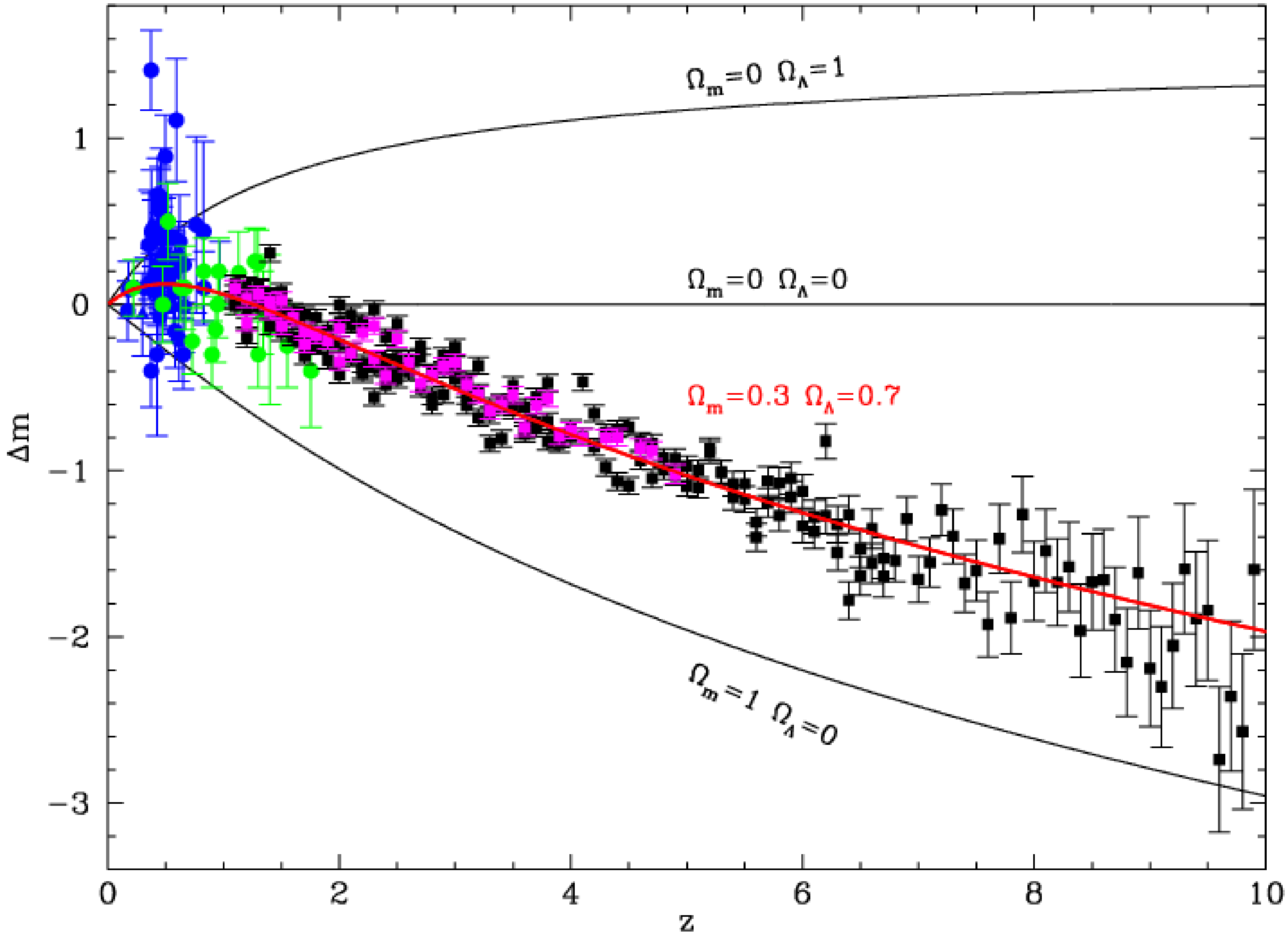

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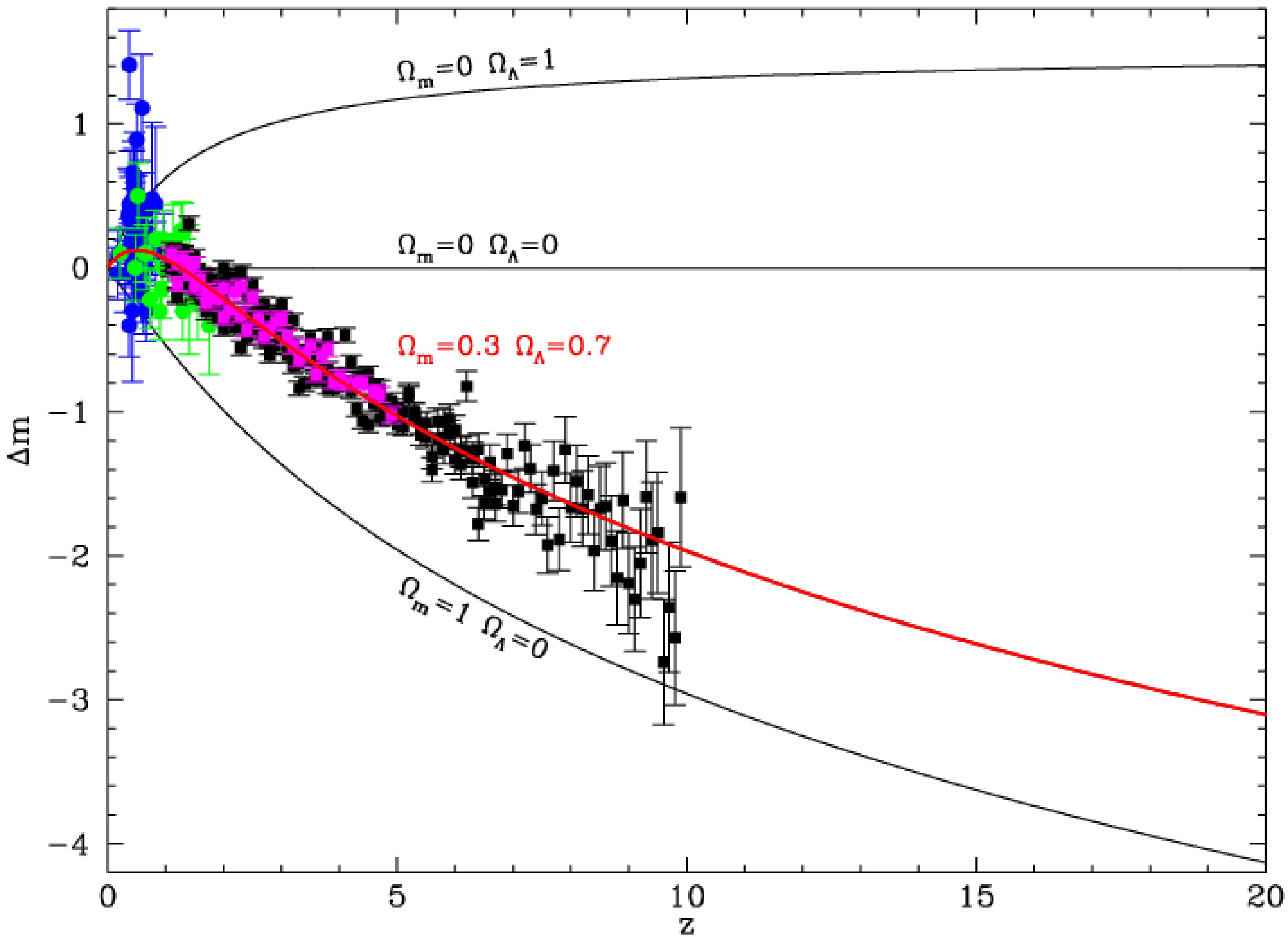

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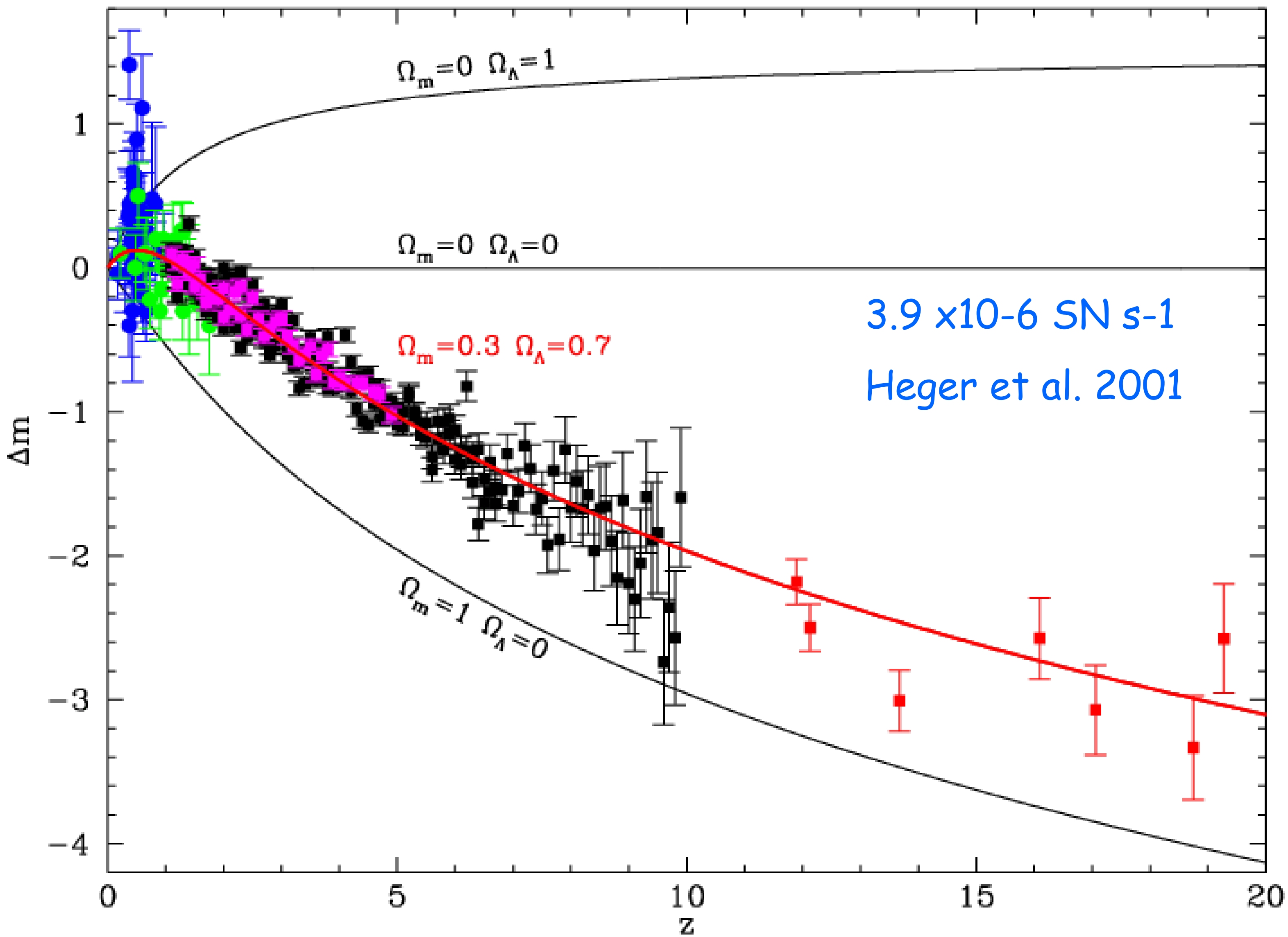

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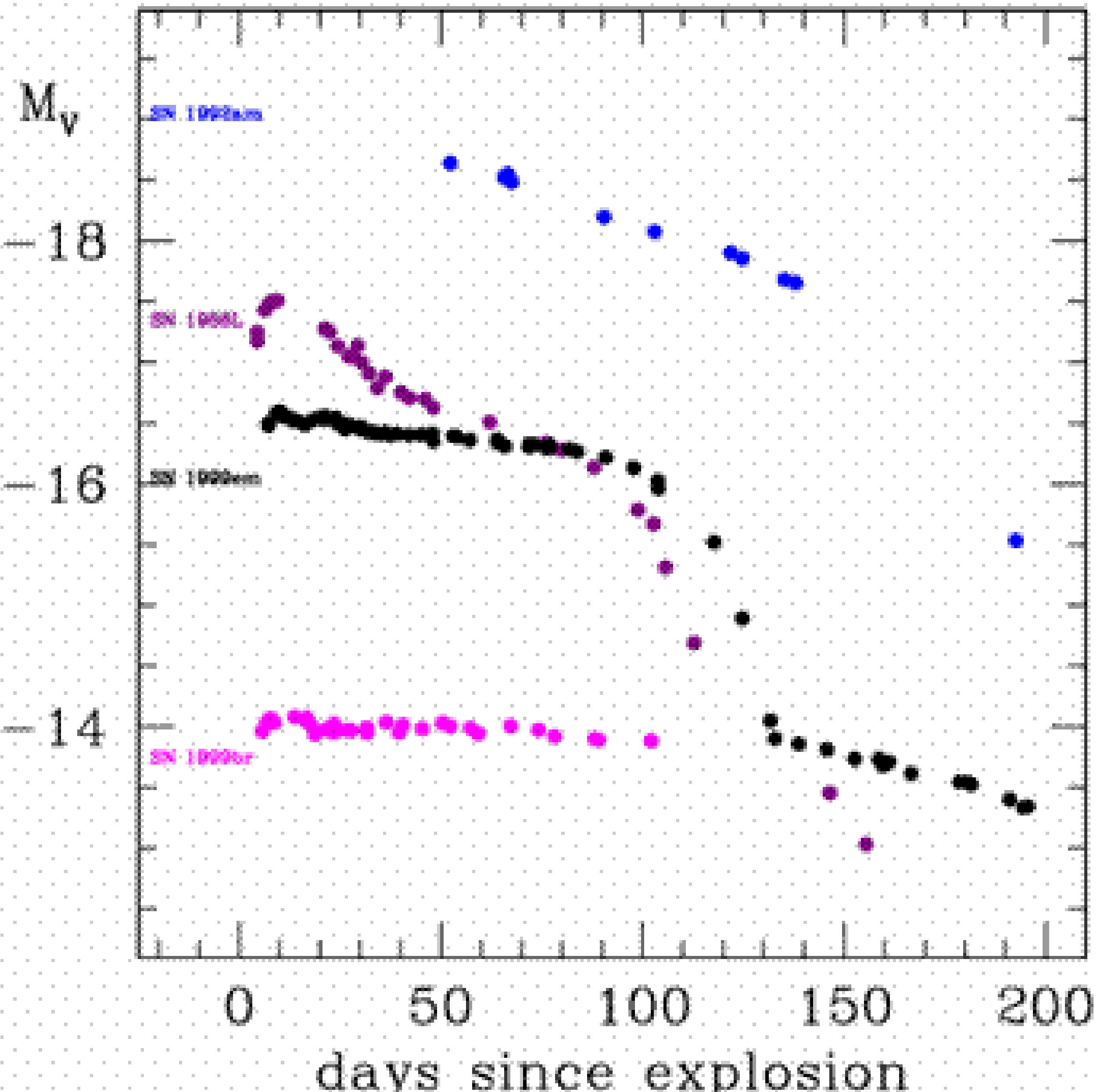

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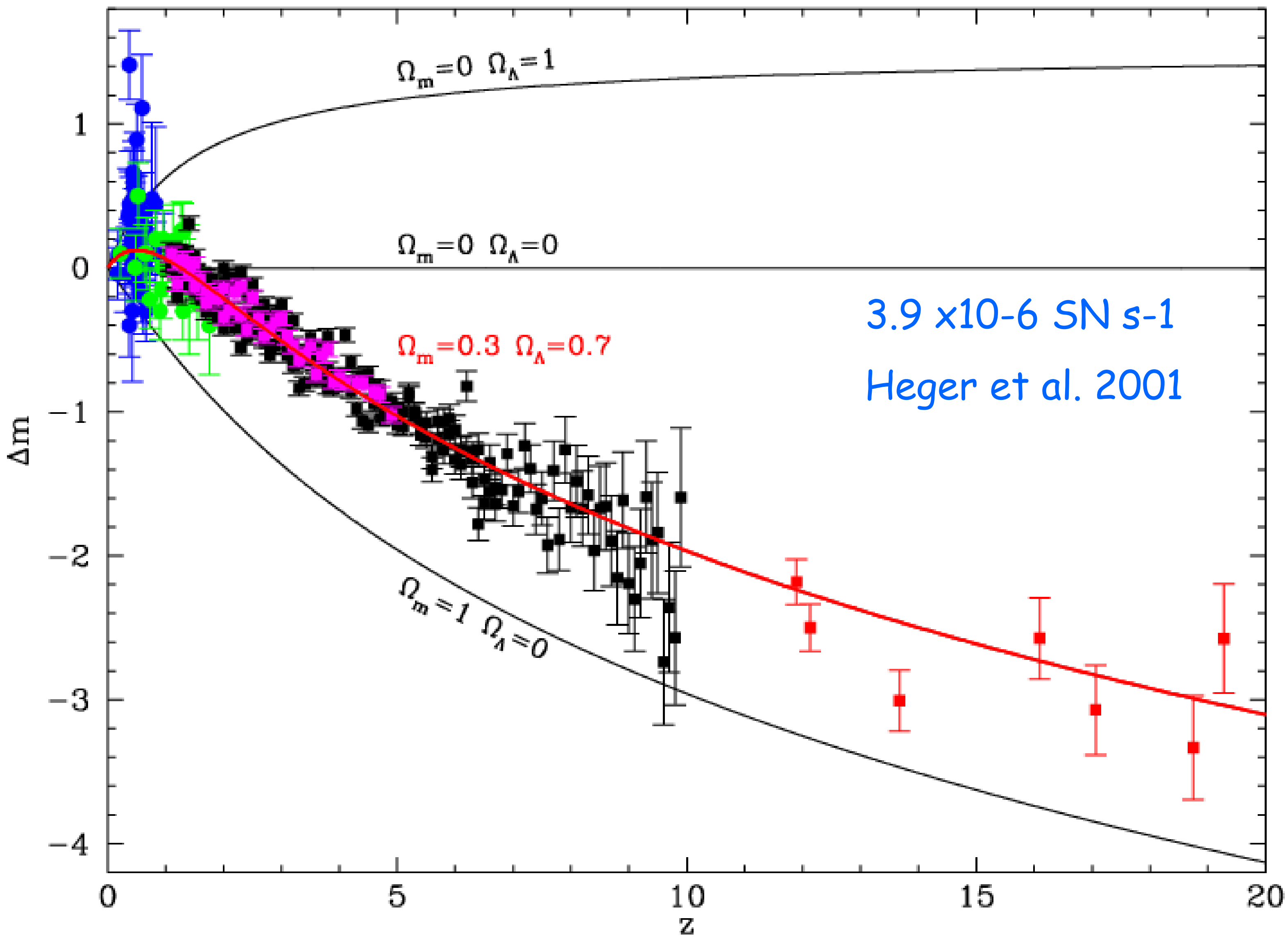

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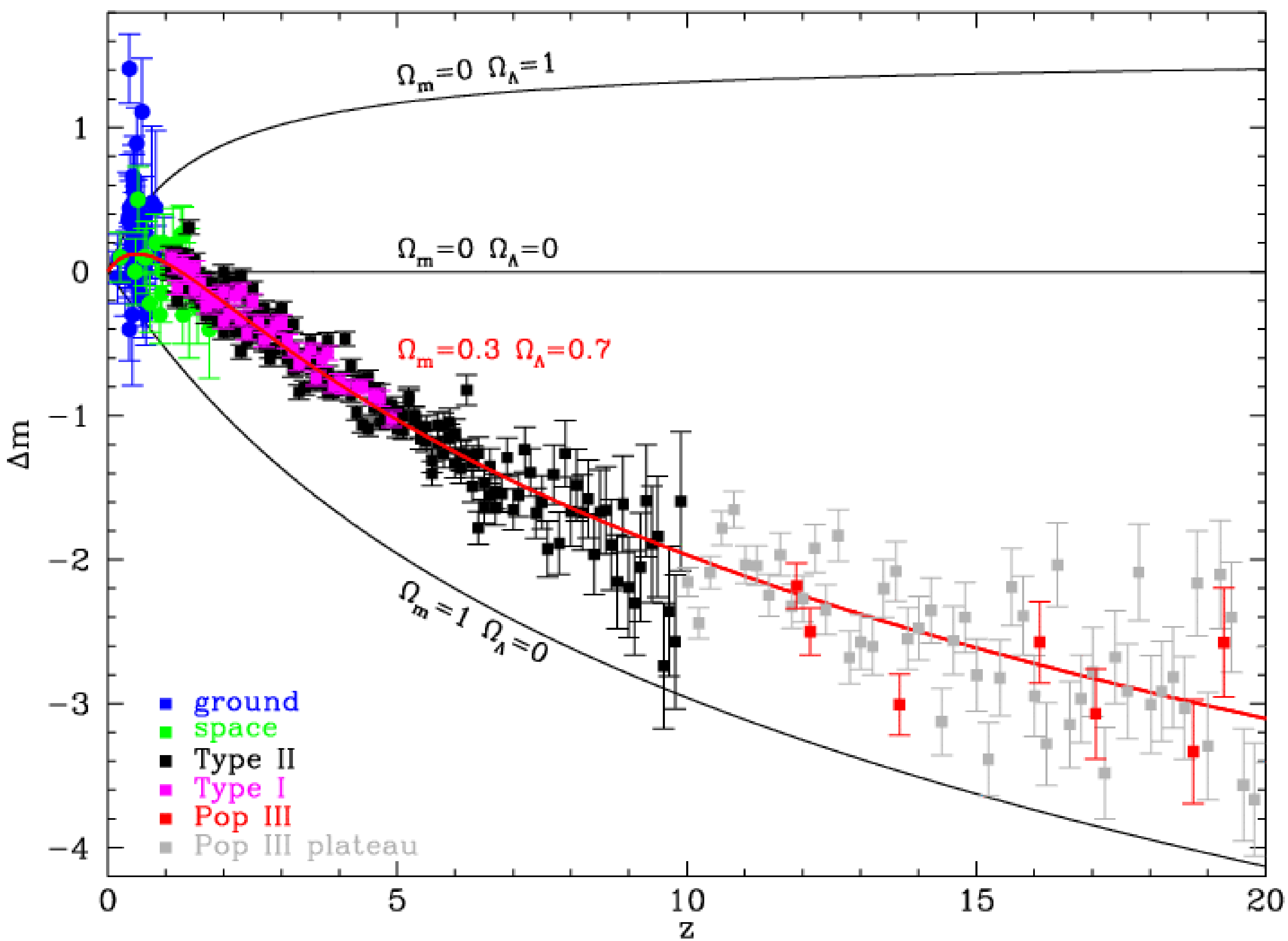


Science with ELTs

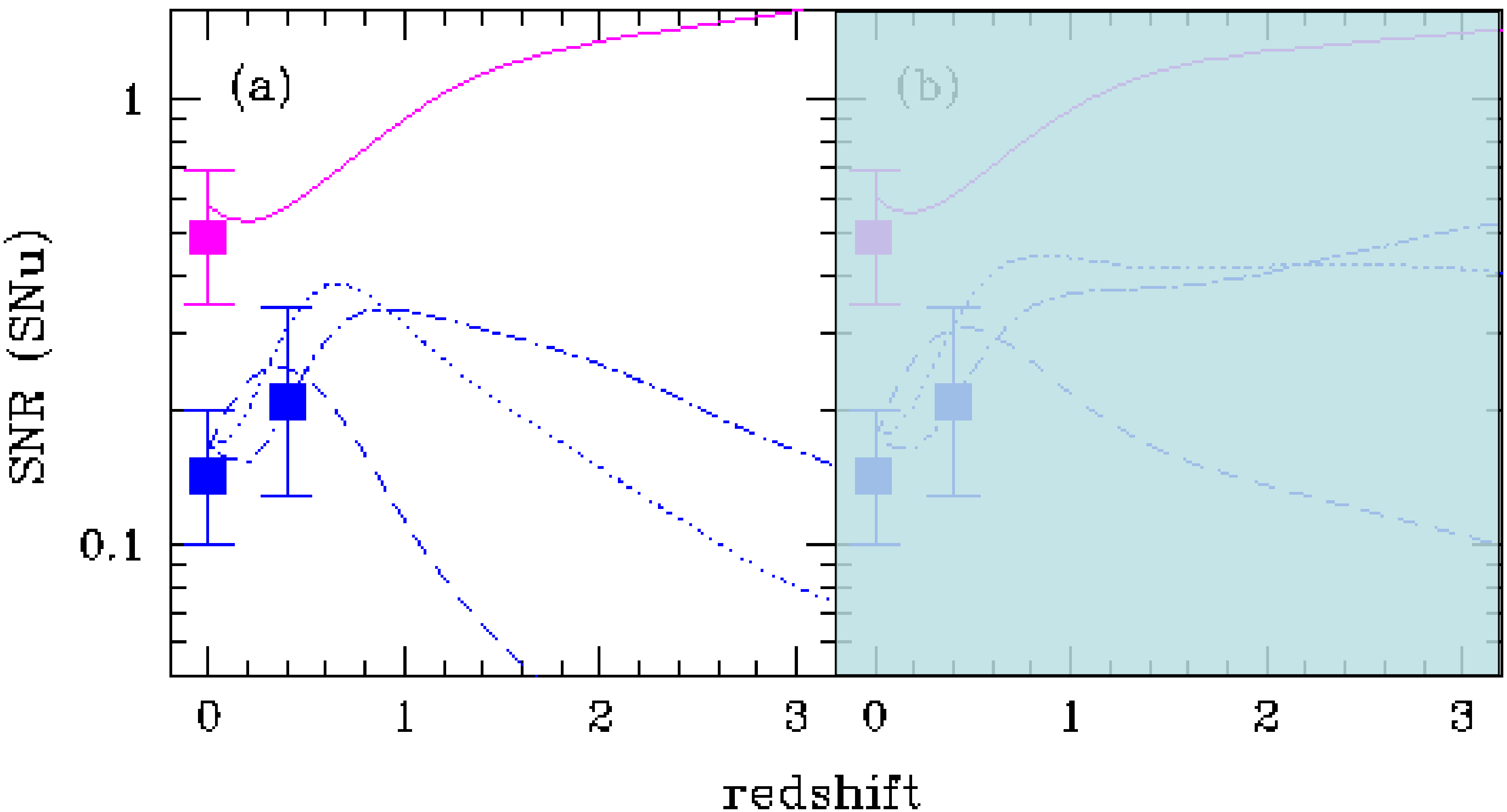
Conclusions

An ELT sample of SNe provides a measurements of the star formation rate up to z=20 which is:

1. Independent of other possible determinations
2. More direct –because the IMF extrapolation is much smaller
3. More reliable because it is based on counting SN explosions rather than relying on identifying and measuring the source of ionization (if using H-alpha flux) or the source of UV continuum
4. We can learn to what extent the IMF was more skewed toward massive stars (relatively to a normal Salpeter) at low metallicities

Conclusions II

5. Disentangling models alternative to Λ
   - Supernovae at z up to ~10: quintessence (?)
6. Also ideal to probe the progenitors of type Ia SNe (sd *vs* dd)
7. To probe the physical properties of the ISM and IGM at z>10 through high spectral resolution ($R \sim 10^4$) of Pop III SNe (feasible with 50-100m)
8. To explore the metal enrichment of the IGM at early epochs (up to z~4) via observations of bright type II and Ia at resolution of ~ 1000 (with 50-100m)



Madau, Della Valle & Panagia 1998

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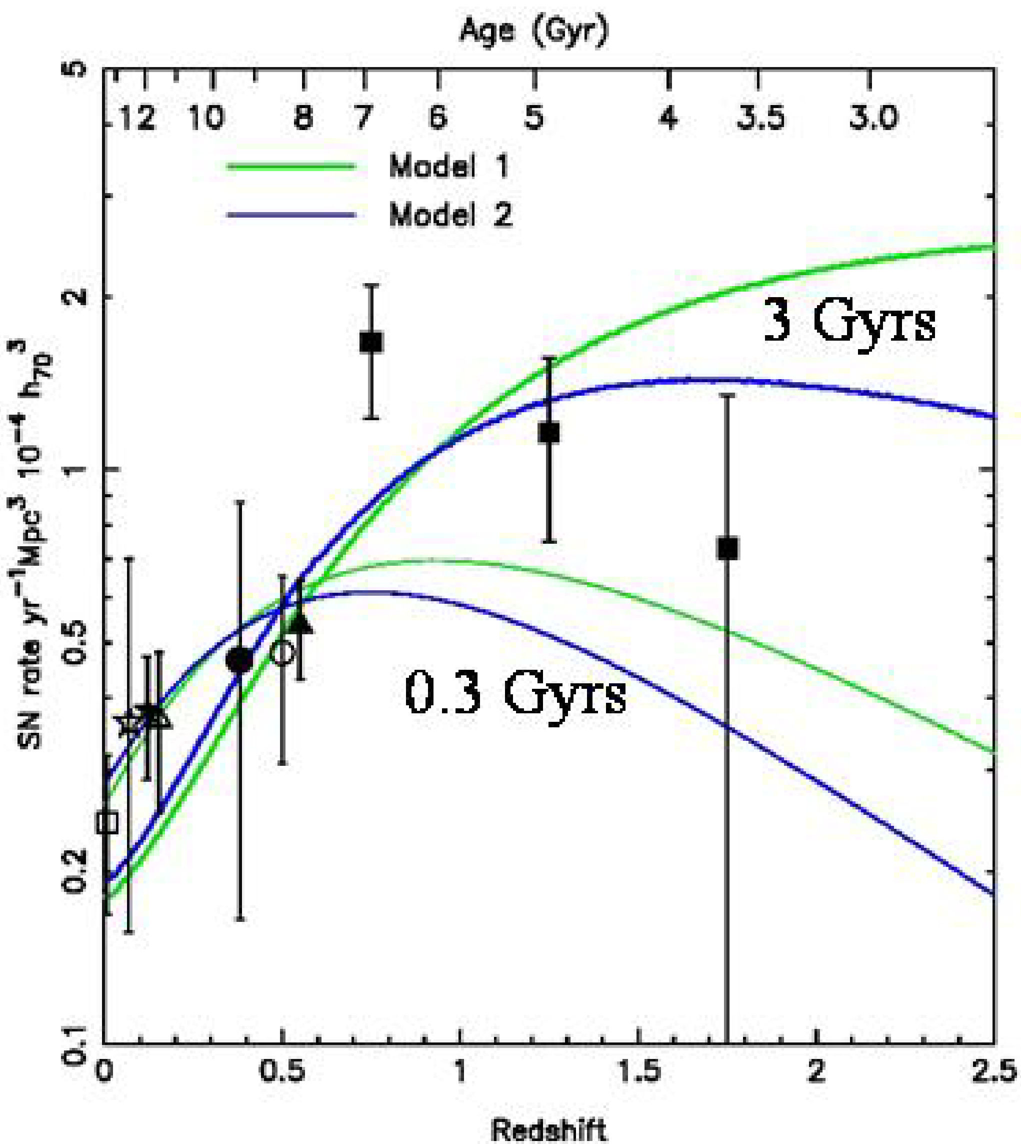

Dahlen et al. 2004

Conclusions III

GRB/SN connection:

- Hypernovae (peculiar type Ic) **1998bw/GRB 980425** (Galama et al. 1998), **2003dh/030329** (Stanek et al. 2003; Hjorth et al. 2003) **2003lw/031203** (Malesani et al. 2004, in press)

  (z < 0.2 and faint gamma sources: 10^4-10)
- Normal type Ic (?) **2002lt/GRB021211** (Della Valle et al. 2003) (z=1)

What are the progenitors of GRBs at very high z ?

LIGO Gravitational Wave Observatory (see also Virgo, Geo, Tama 300 and Aciga)

$M = 100\text{-}300\ M_\odot$ $E_{GW} \sim 10^{-3} M_\odot c^2$

$\rightarrow h \sim 10^{-22}$

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