

DFG-Forschergruppe: Magnetisation of interstellar and intergalactic media

Commissioning and pilot observations
(A9, Beck)

RM Synthesis and signal extraction
(A8, Enßlin)

Research Infrastructure & Service



Milky Way
(A1, Reich)

Groups and interacting
winds
(A5, Bomans)

Dynamo
models with
winds
(A4, Elstner)

Radio halos around dwarf galaxies
(A3, Klein)

Magnetic fields in clusters
(A6, Brüggen)

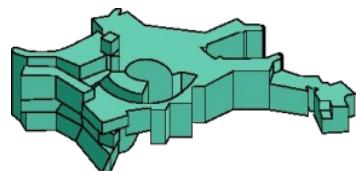
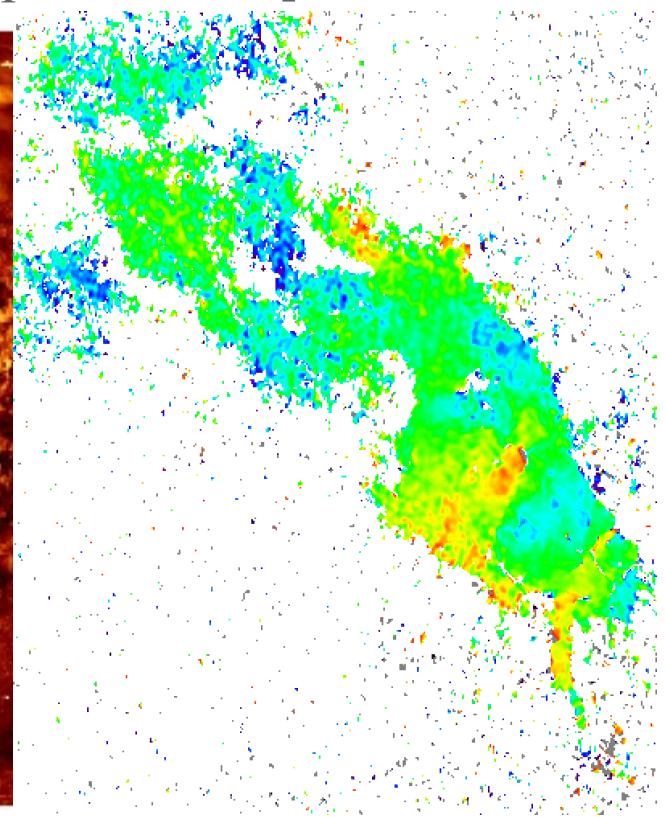
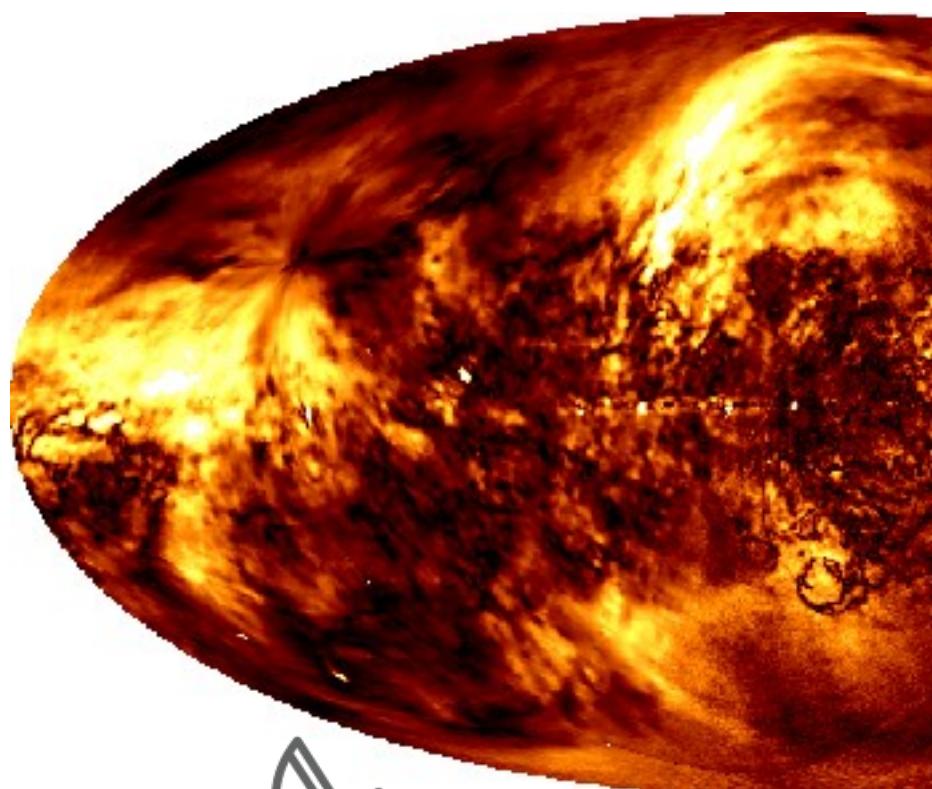
Nearby spiral and elliptical galaxies
(A2, Dettmar)

Radio relics and halos
(A7, Klein)

Scientific Research

A8 RM Synthesis & magnetic field statistics

Torsten Enßlin, Rainer Beck, Marcus Brüggen
Michael Bell, Thomas Riller, Henrik Junklewitz
[Andre Waelkens, Petr Kuchar, Niels Oppermann, ...]

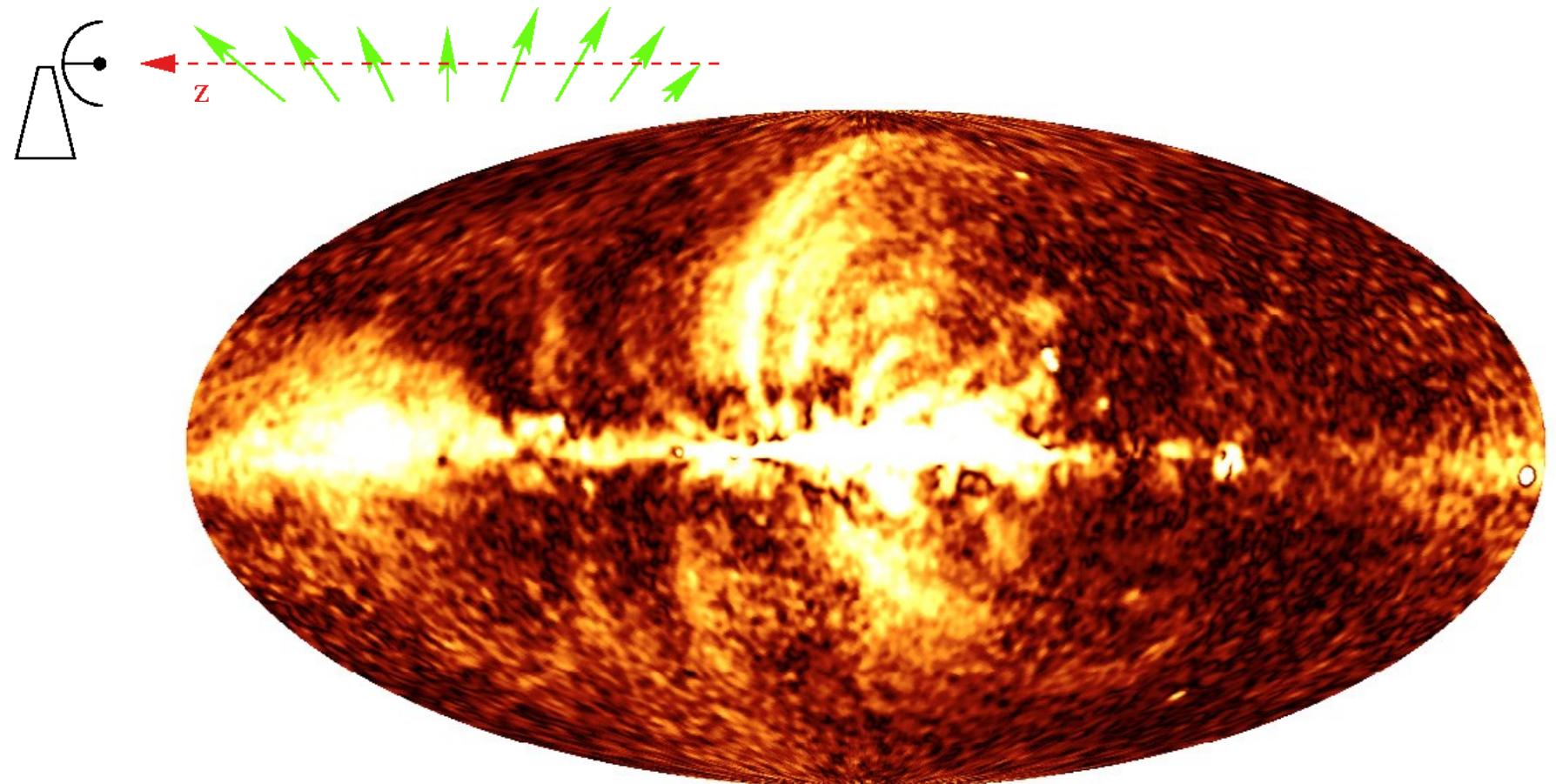


Max-Planck-Institut
für Astrophysik



Max-Planck-Institut
für
Radioastronomie

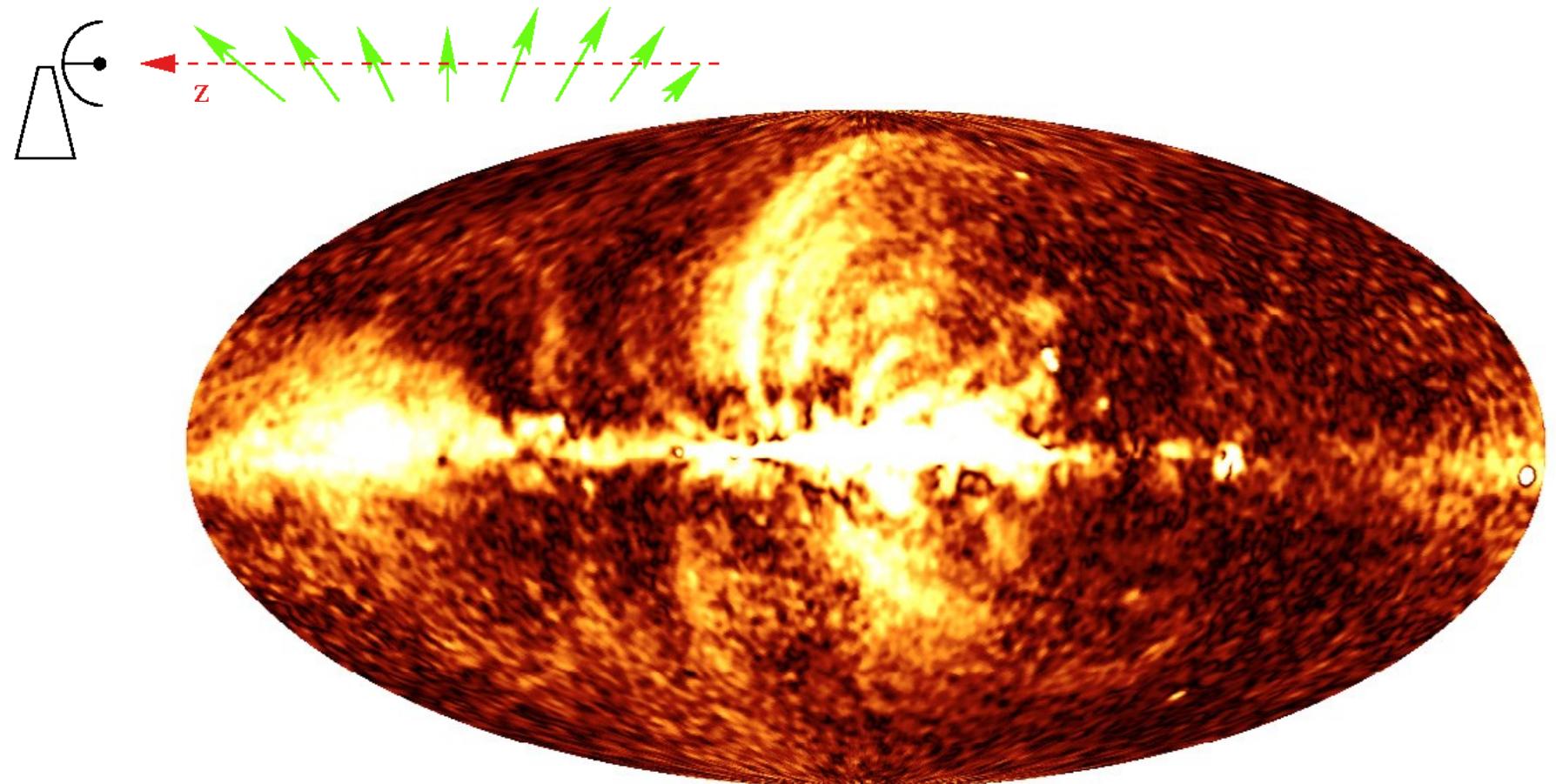
Galactic Magnetic Fields



1.4 GHz: Reich & Wolleben

22 GHz: WMAP team

Galactic Magnetic Fields

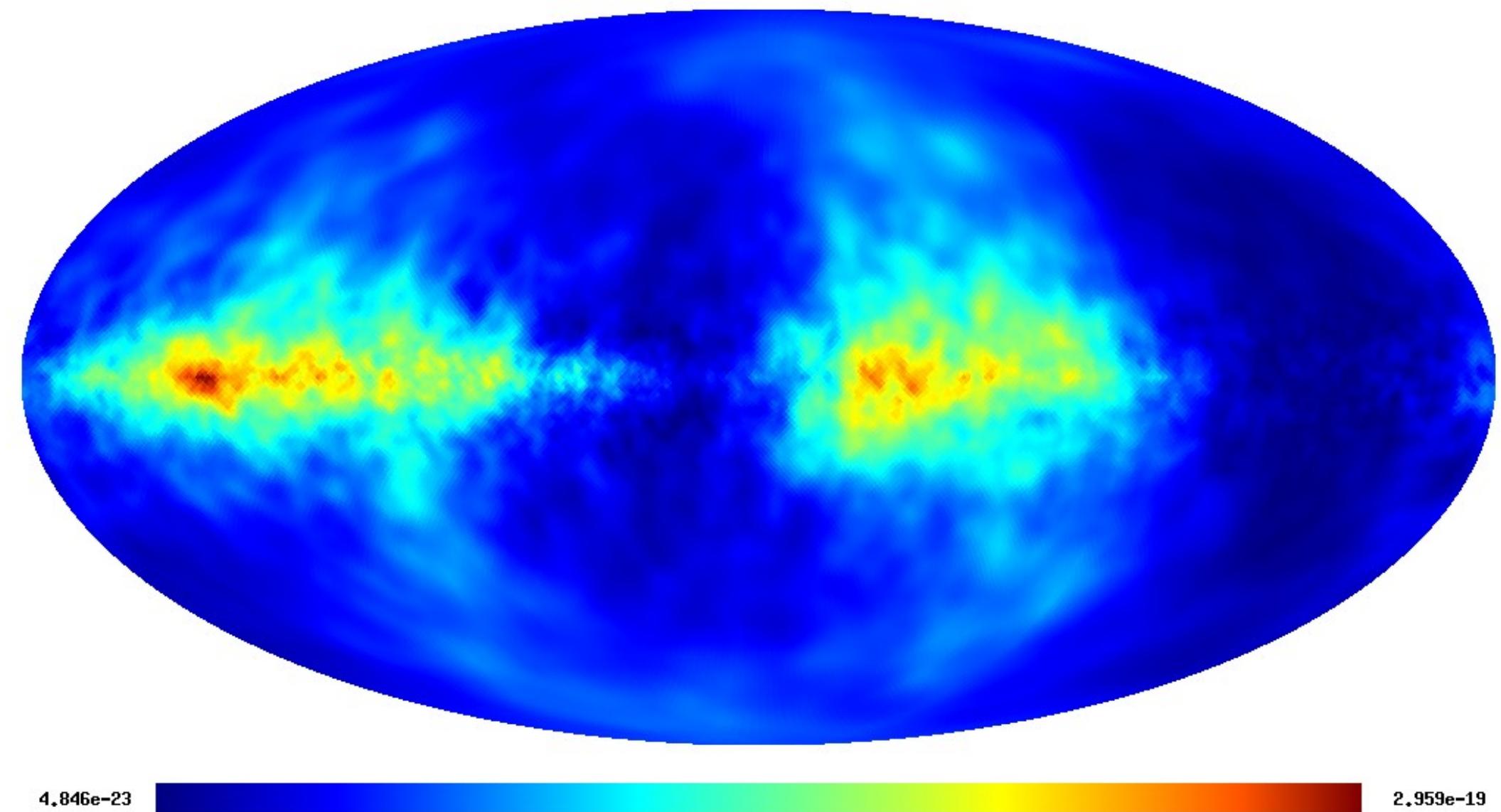


1.4 GHz: Reich & Wolleben

22 GHz: WMAP team

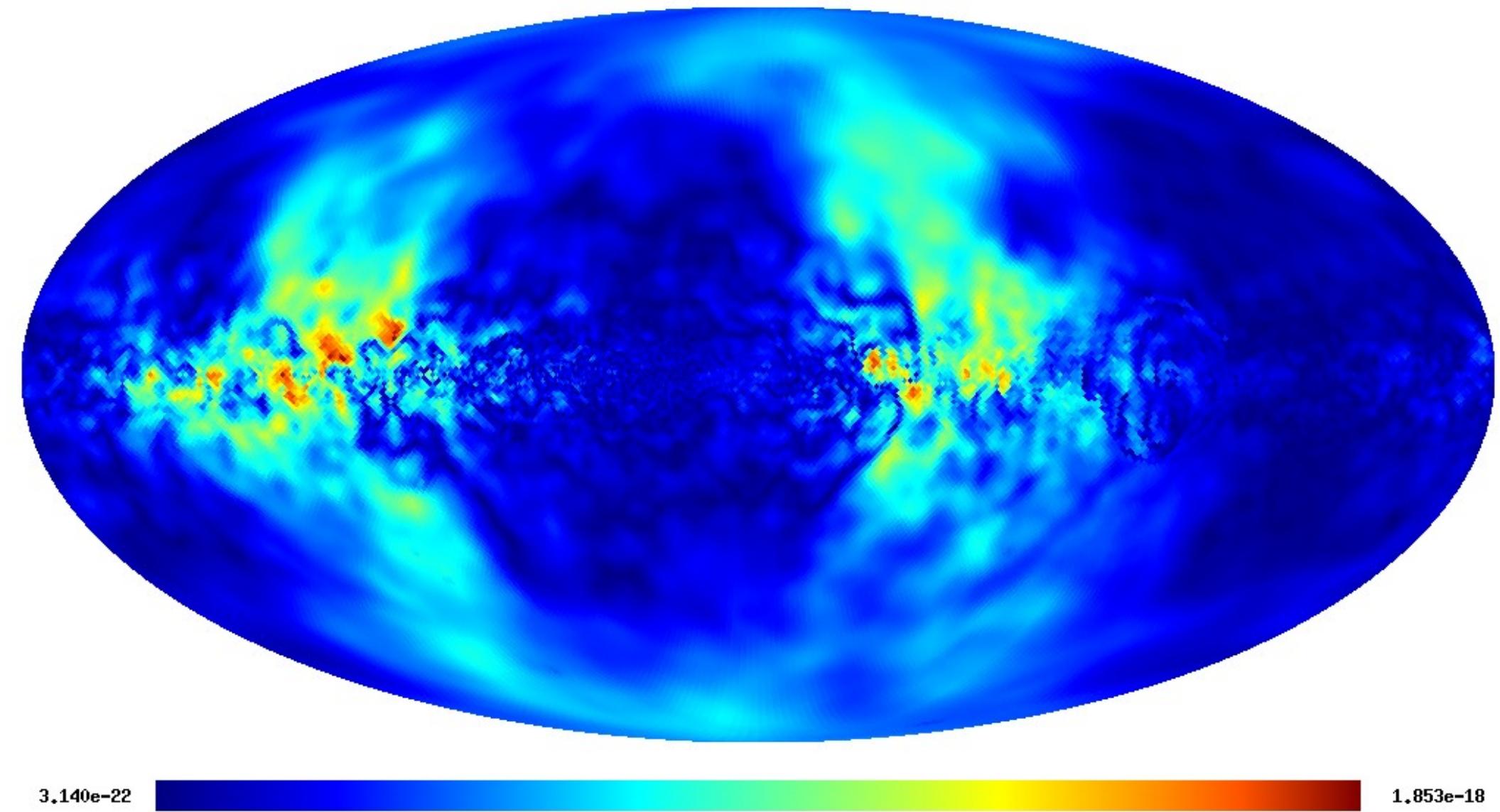
22 GHz

Hammurabi-Simulation



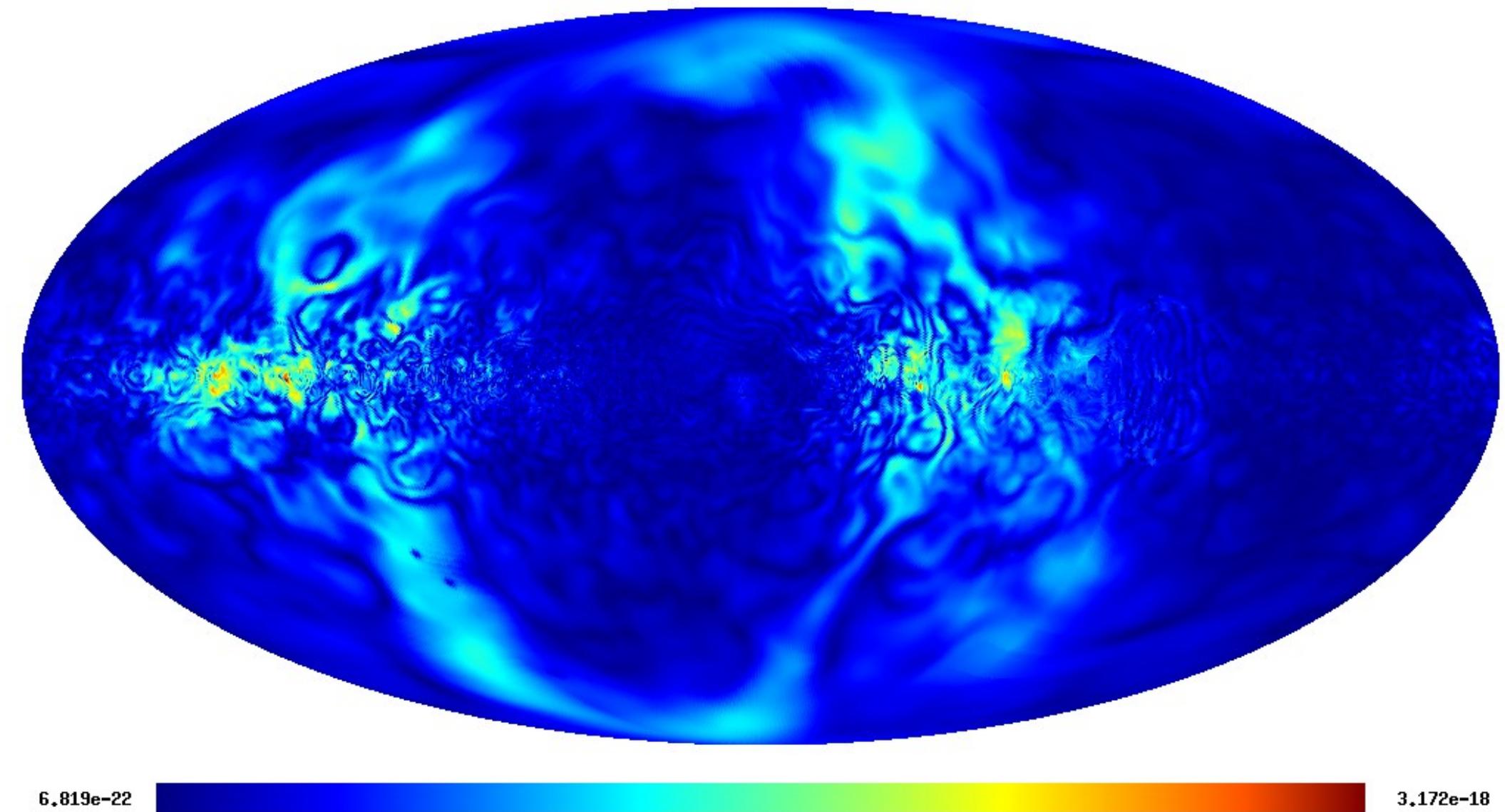
2.4 GHz

Hammurabi-Simulation



1.4 GHz

Hammurabi-Simulation

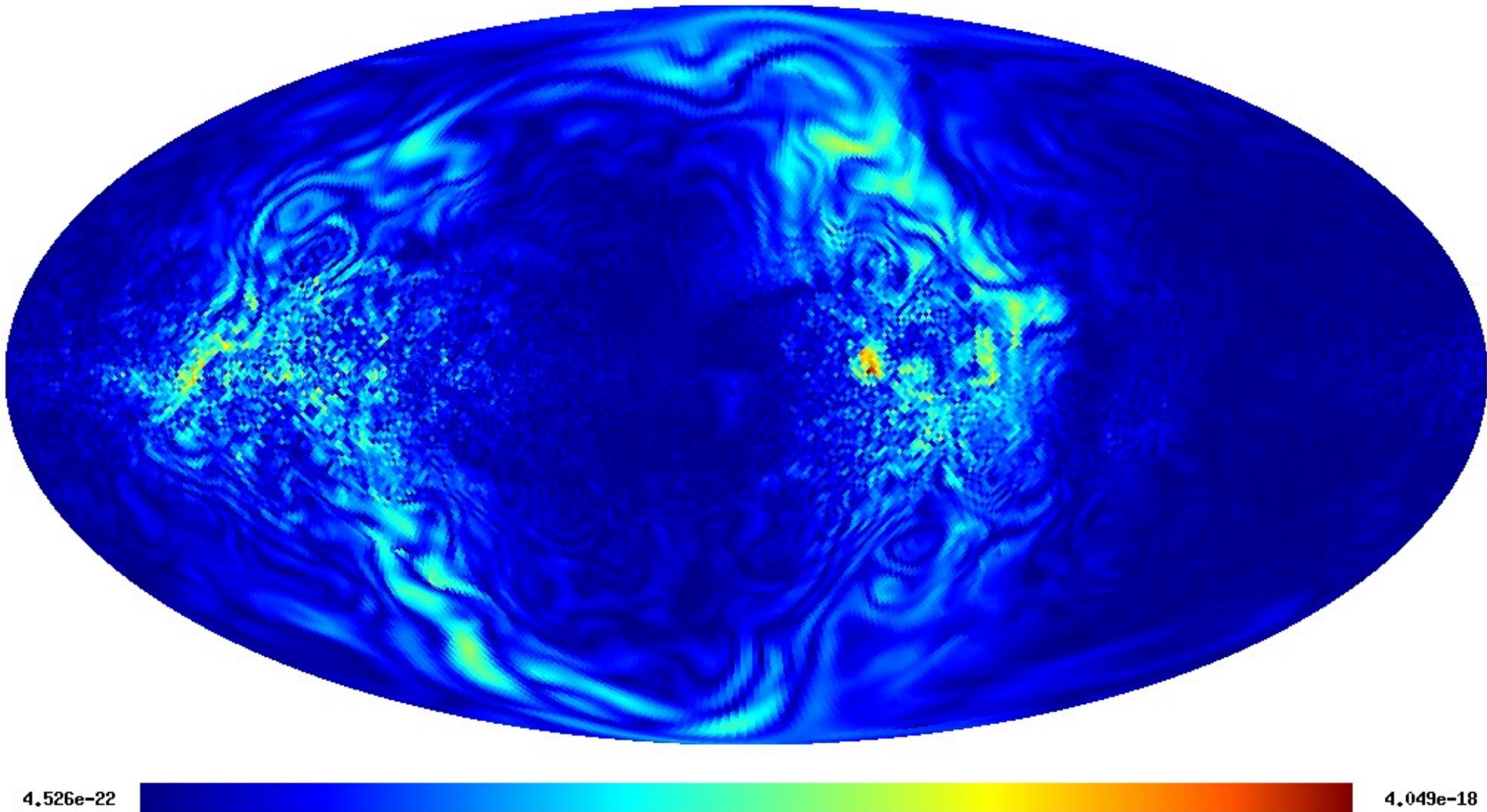


$6.819\text{e-}22$

$3.172\text{e-}18$

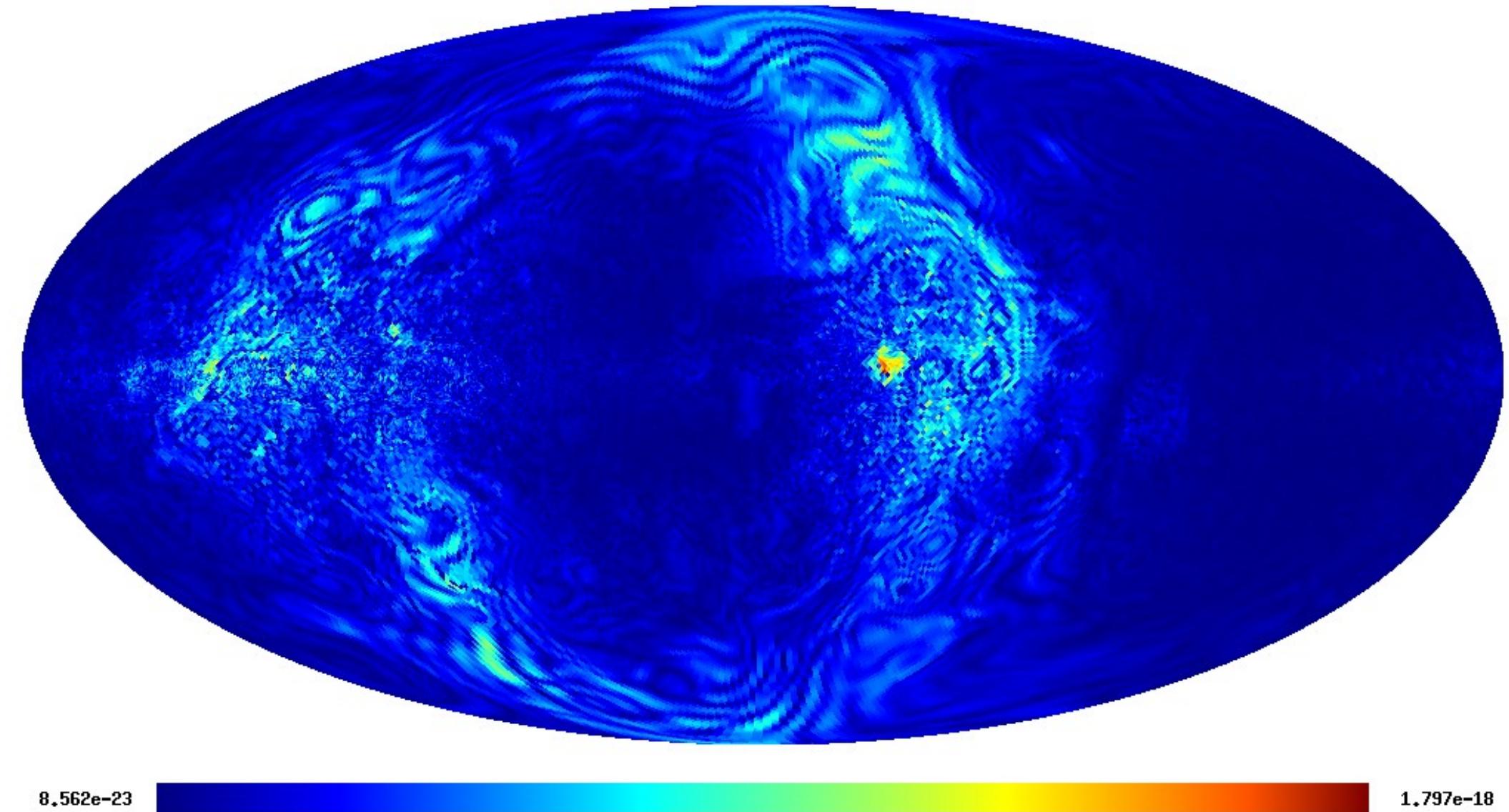
0.8 GHz

Hammurabi-Simulation



0.4 GHz

Hammurabi-Simulation

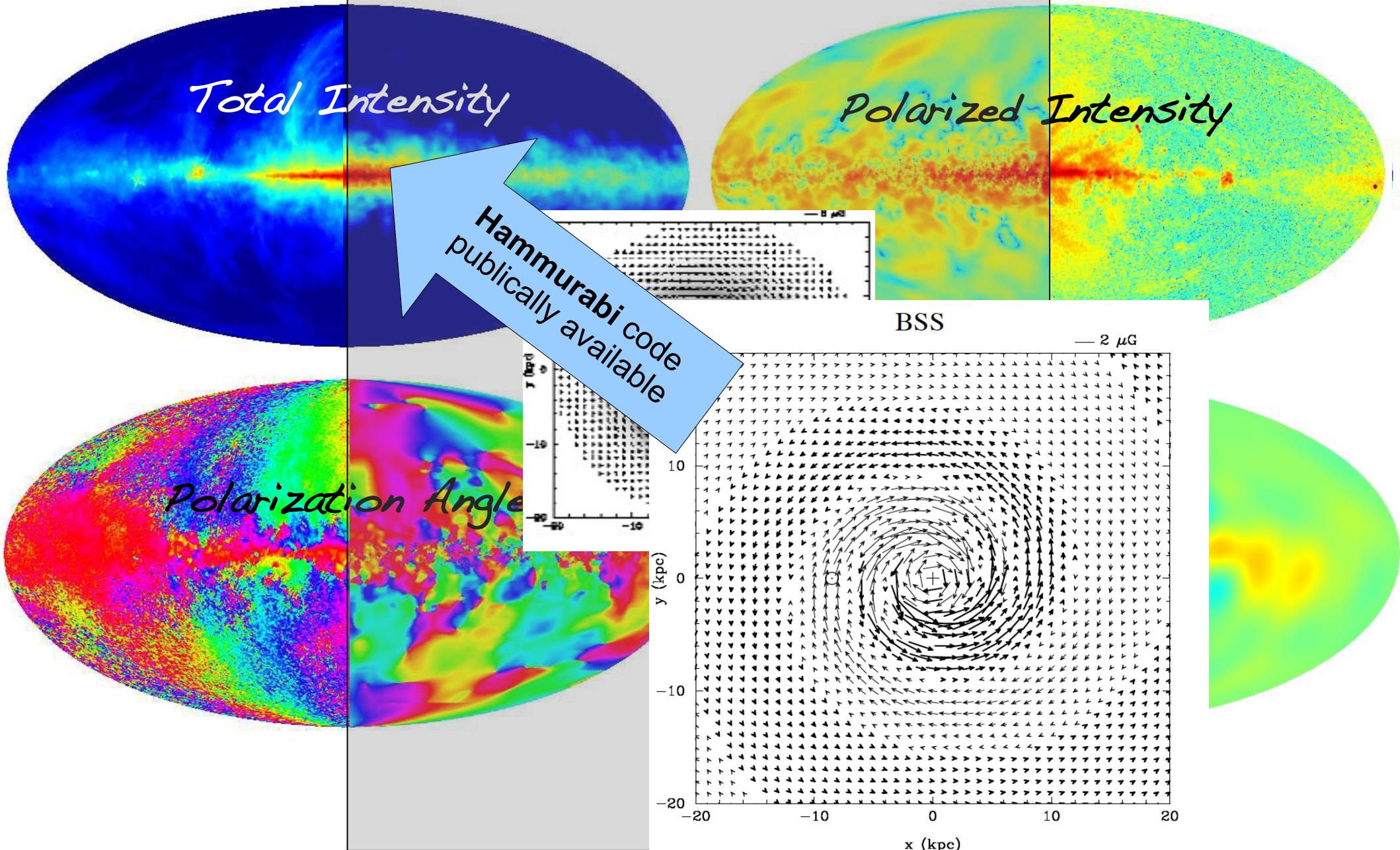


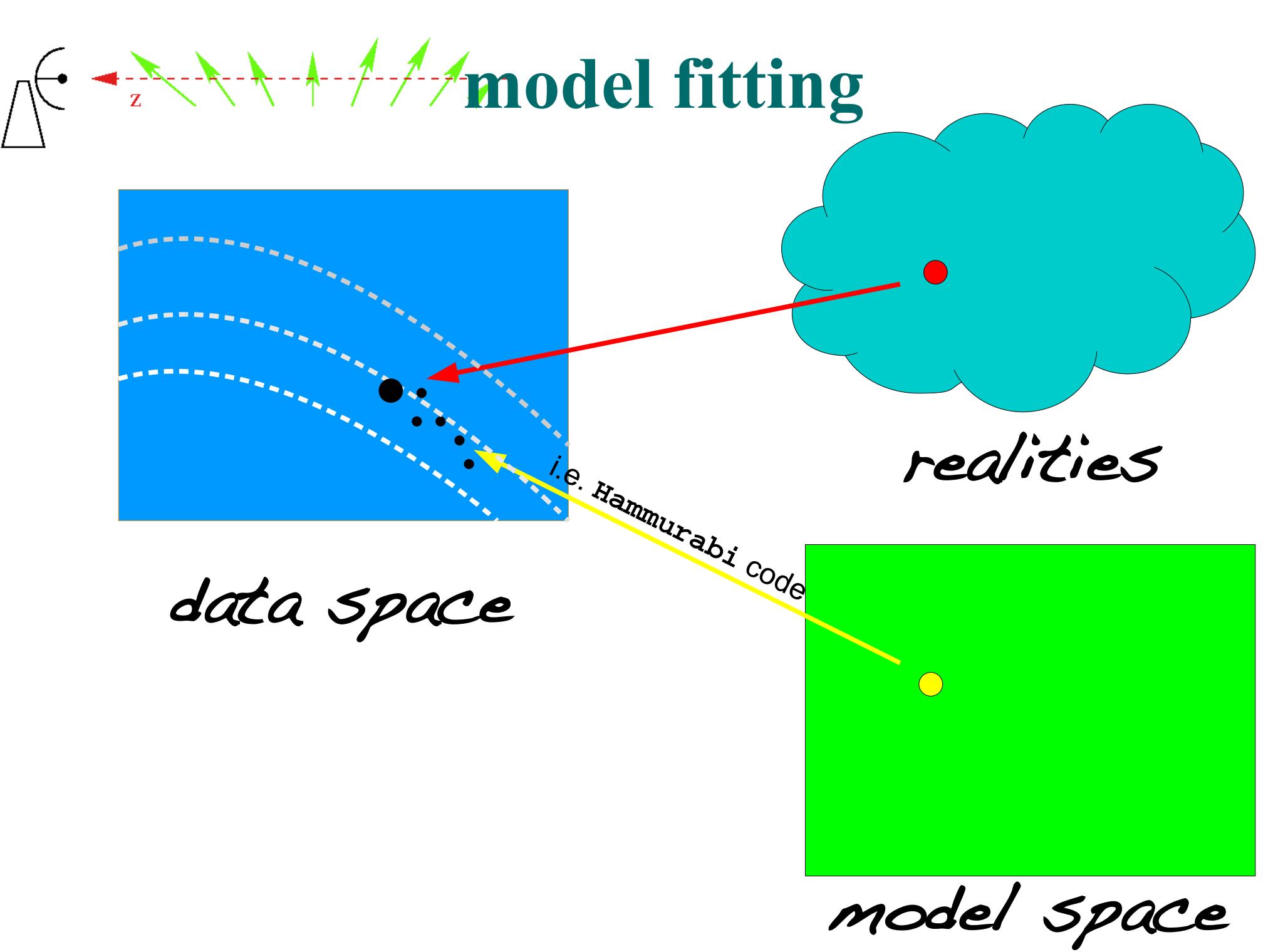
8.562e-23

1.797e-18

Hammurabi simulation of galactic magnetic field observables

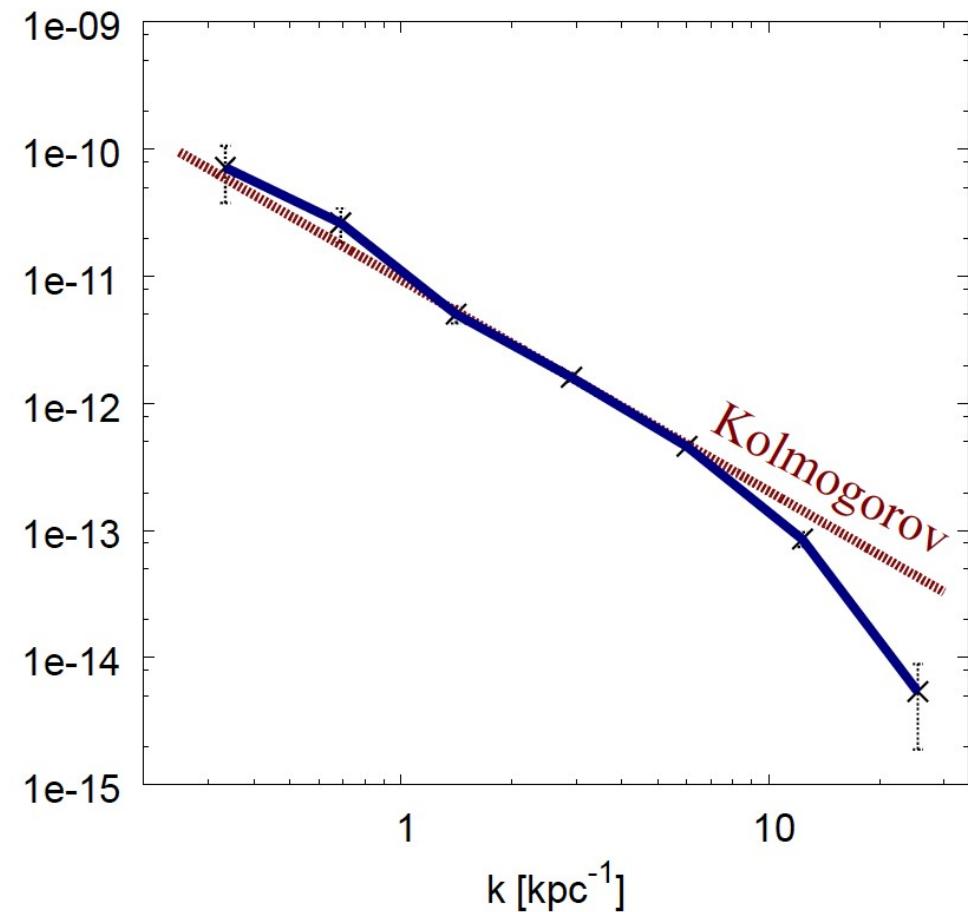
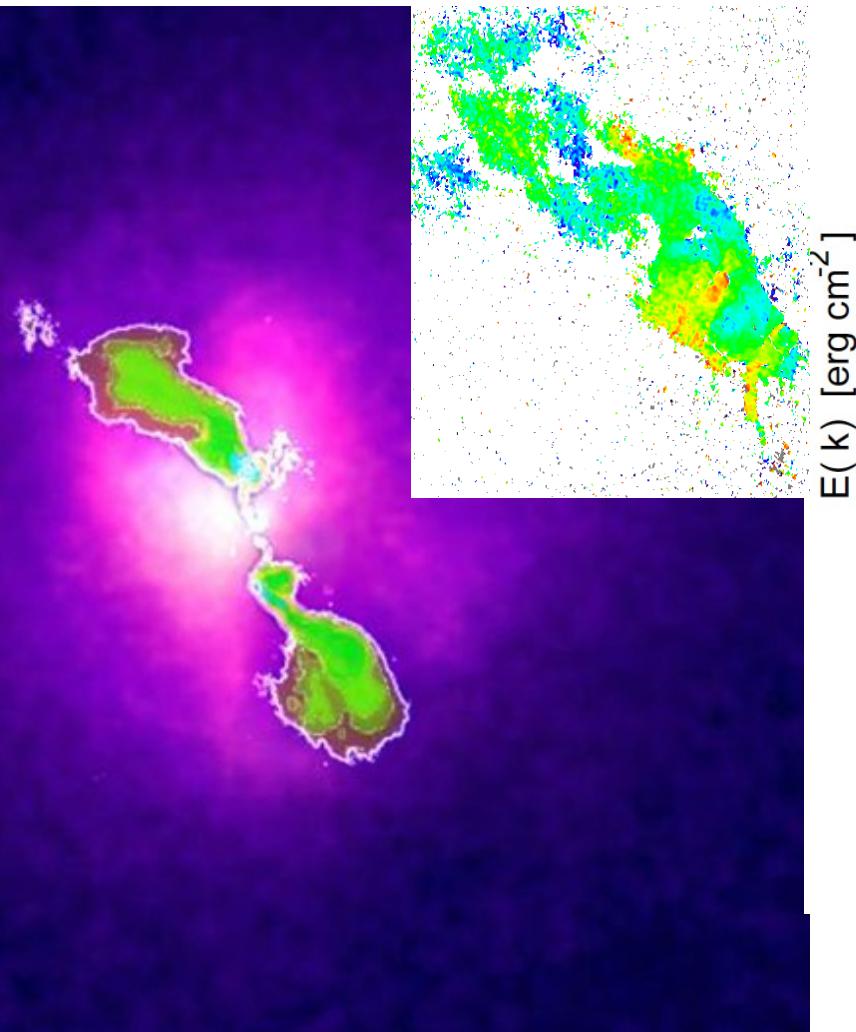
Waerken et al. (2009)
Sun et al. (2008)
Janson et al. (2008)



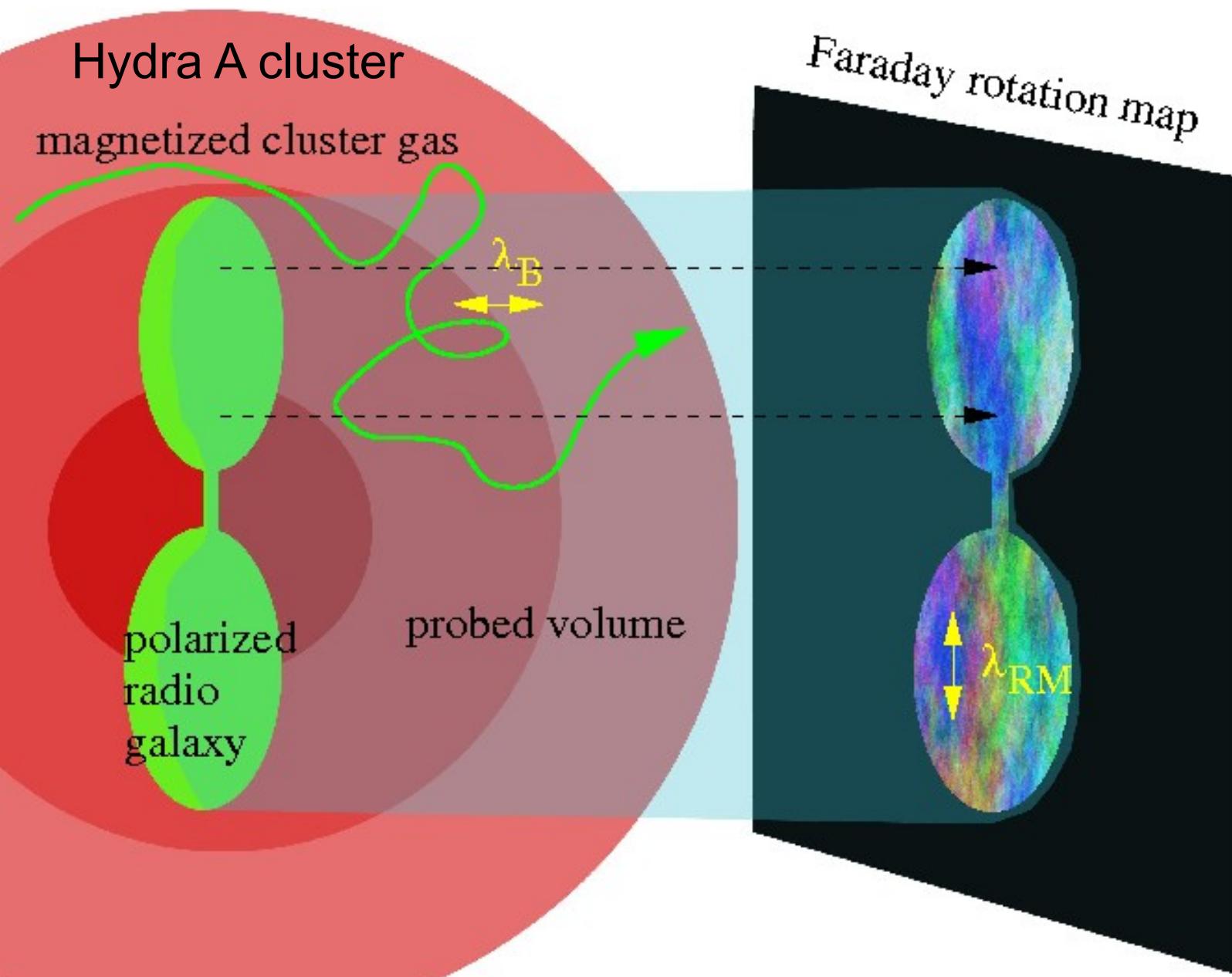


Magnetic power spectrum in Hydra A cool core cluster

Kuchar & EnBlin (2009)



Observational Setup



Observational Setup

Hydra A cluster

$$w(r) = \langle B(x) \cdot B(x+r) \rangle$$

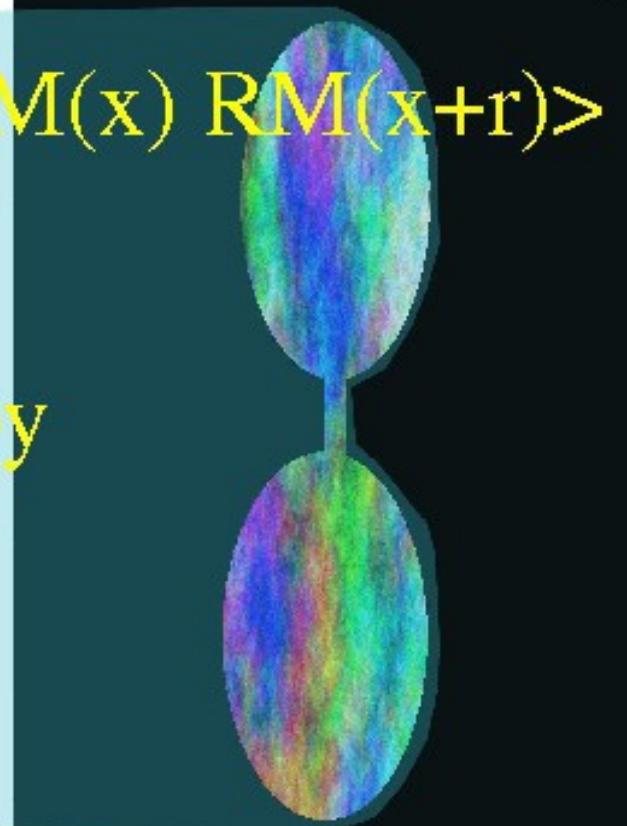
$$C_{RM}(r) = \langle RM(x) \cdot RM(x+r) \rangle$$

Assumptions:

- statistical isotropy
- $\text{div } B = 0$

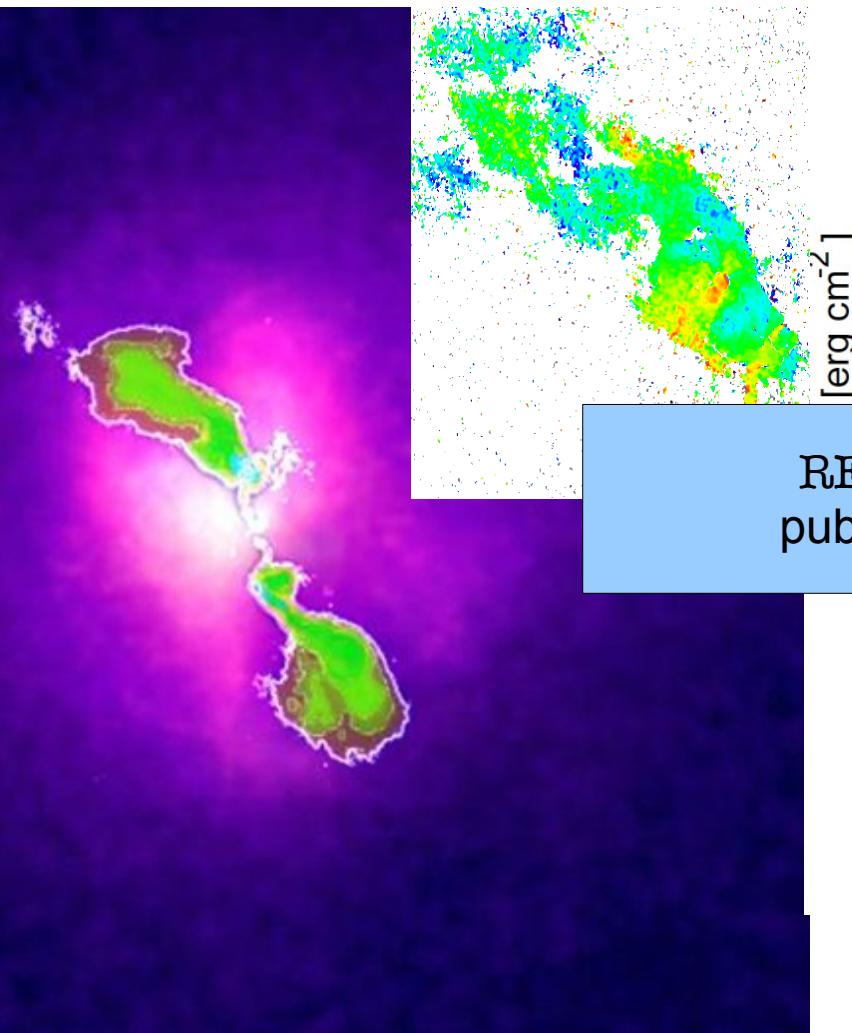
$$C_{RM}(k) \sim w(k)/2$$

Faraday rotation map

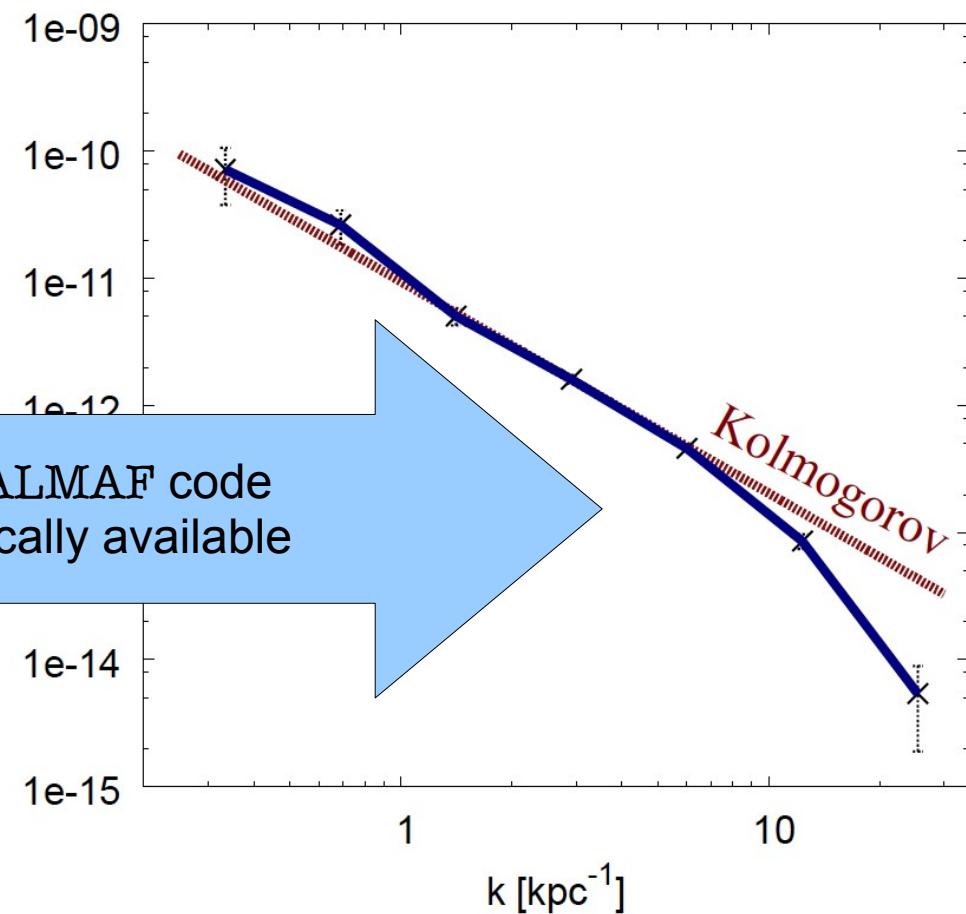


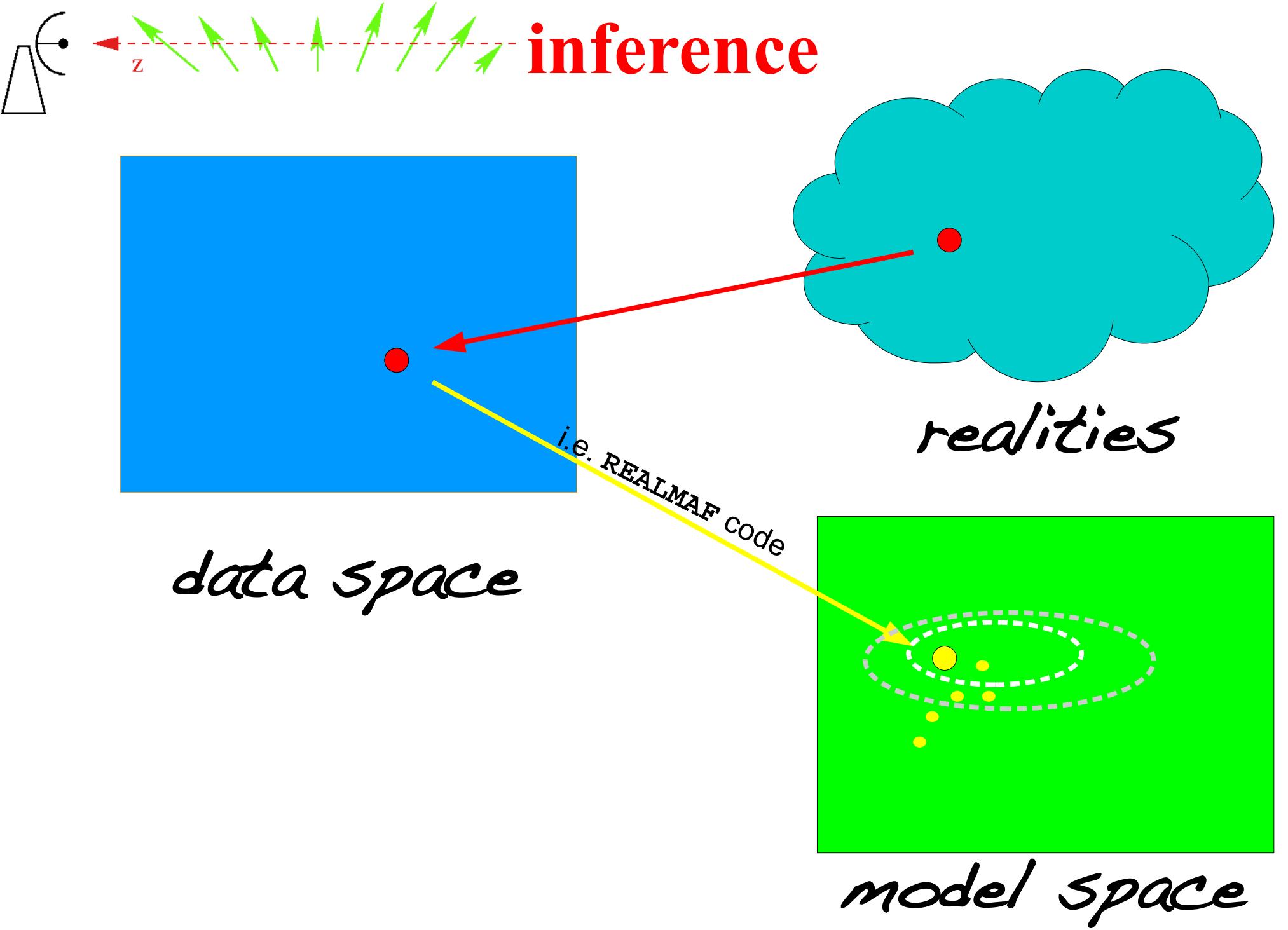
Magnetic power spectrum in Hydra A cool core cluster

Kuchar & EnBlin (2009)



REALMAF code
publically available



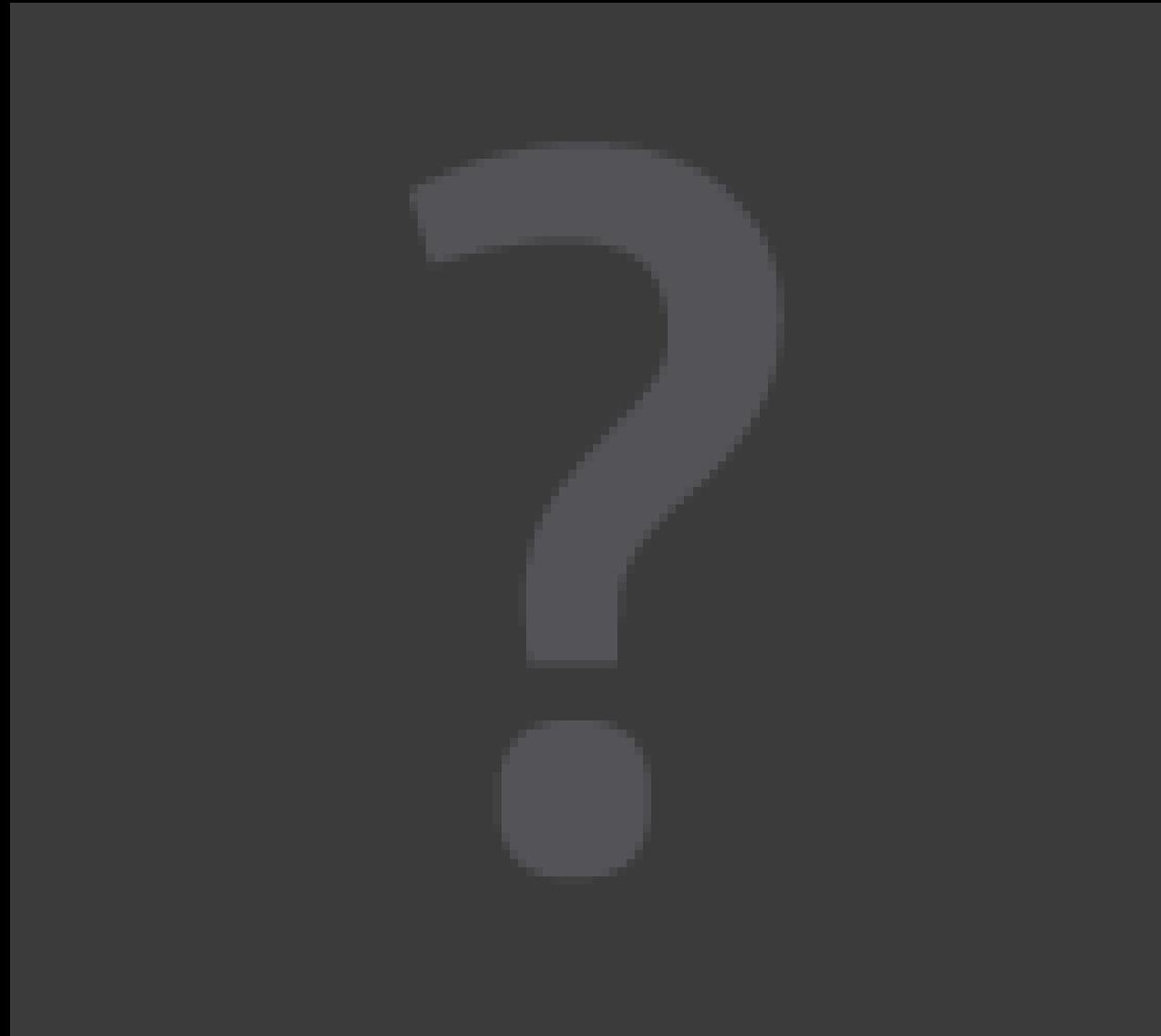


RM Synthesis

de Bruyn & Brentjens (2005)

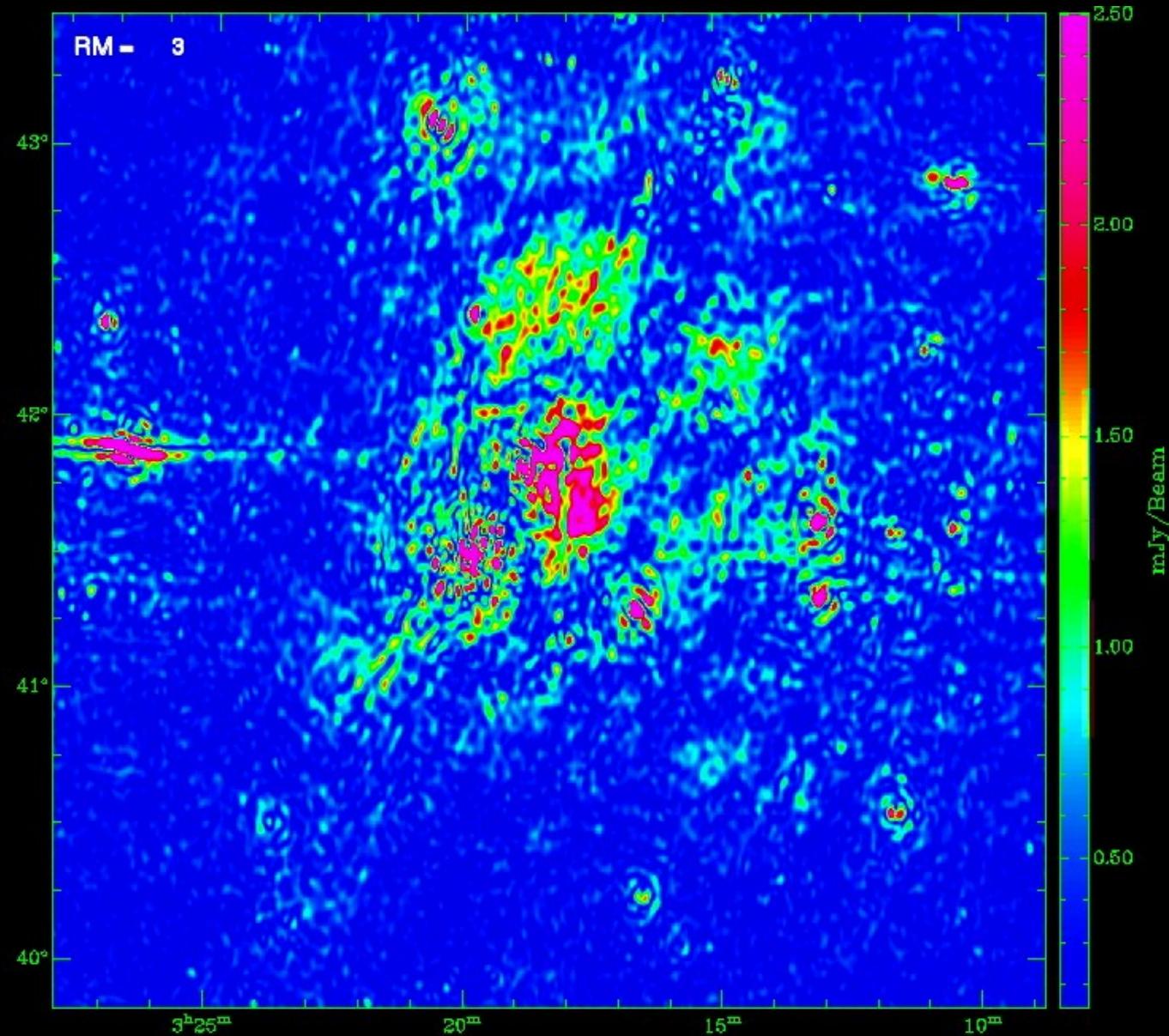
RM Synthesis

de Bruyn & Brentjens (2005)



RM Synthesis

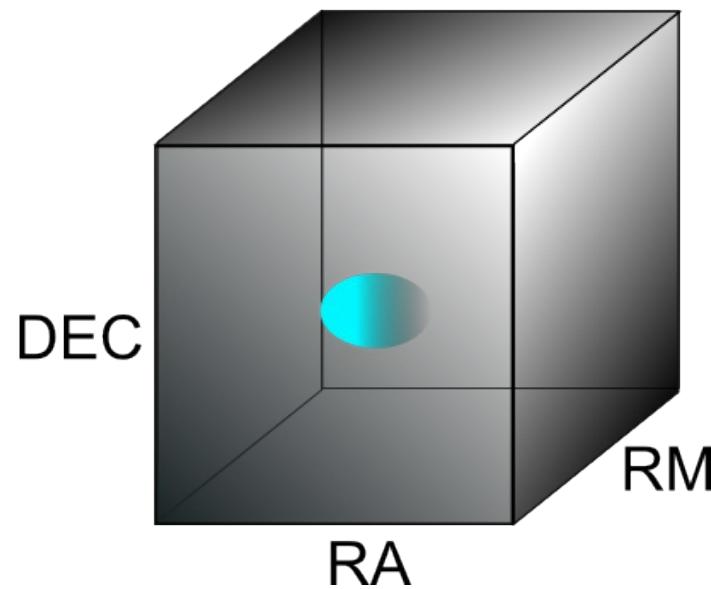
de Bruyn & Brentjens (2005)



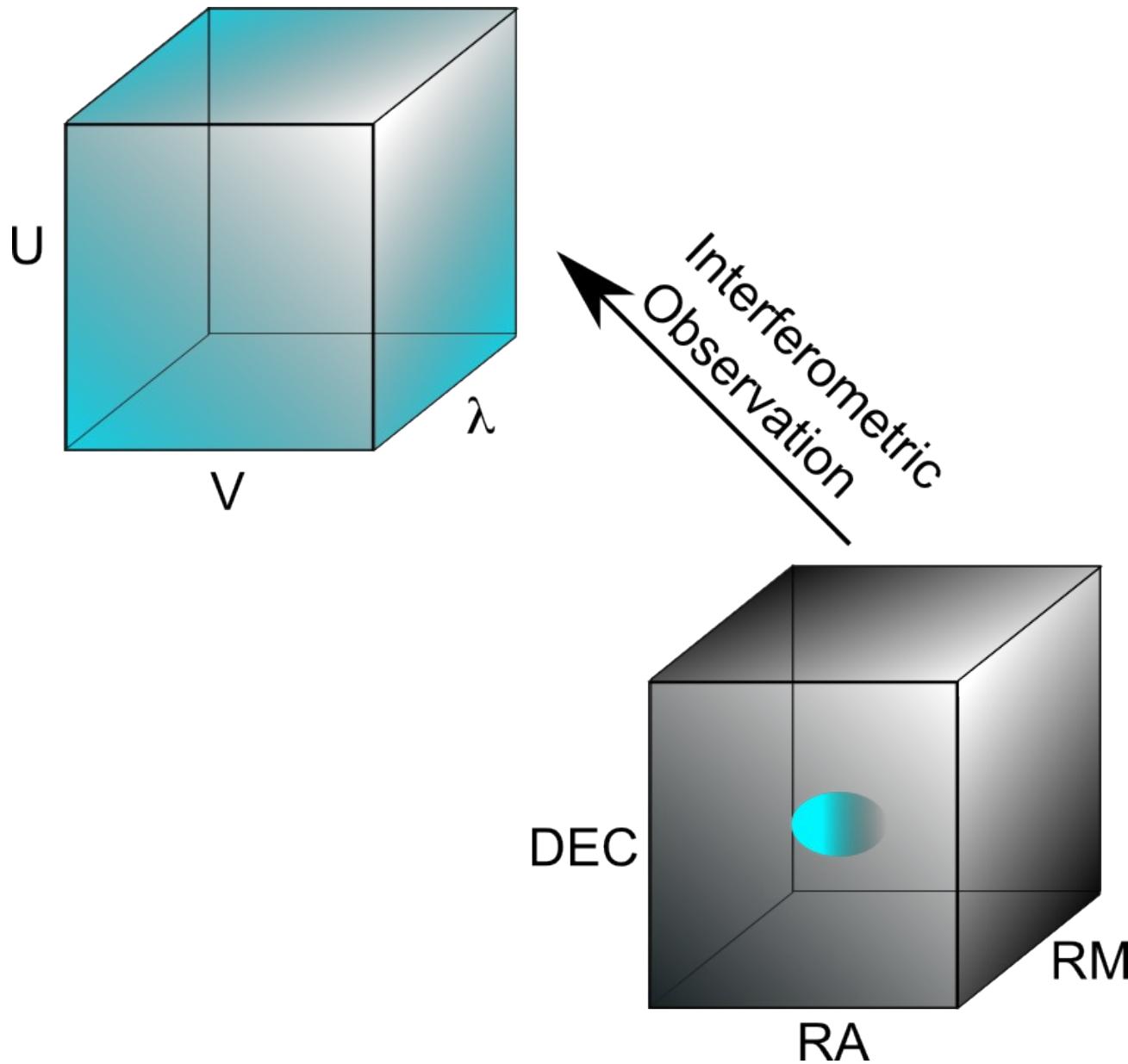
RM Synthesis

3d RM Synthesis

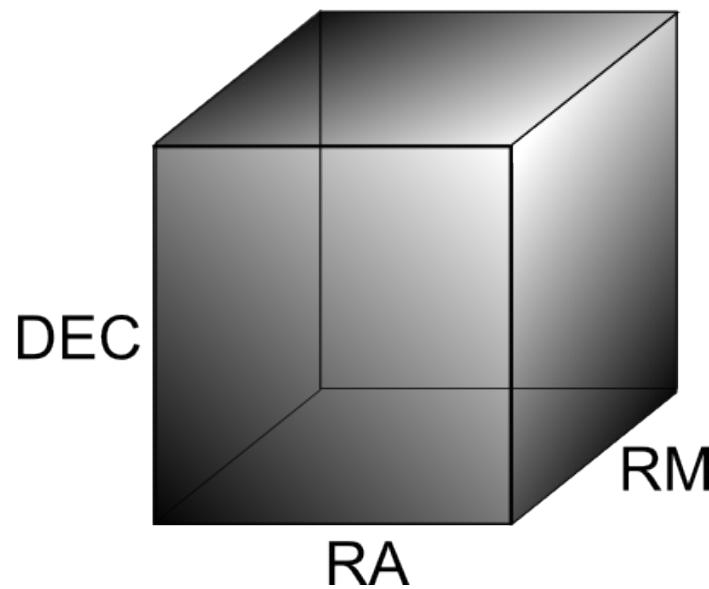
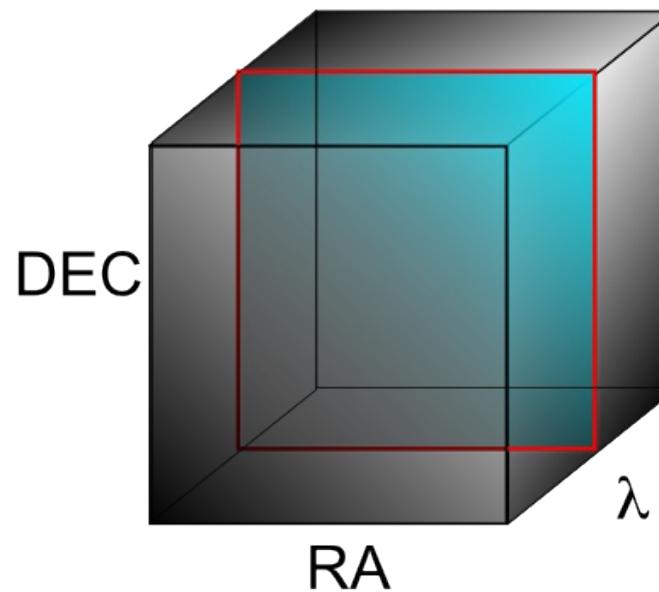
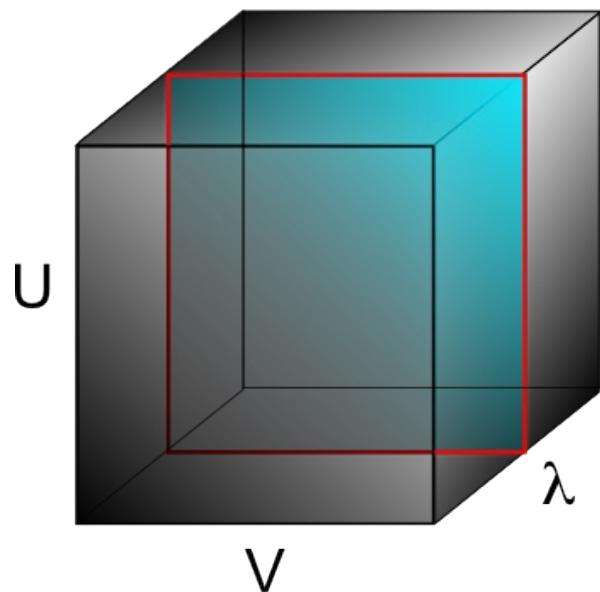
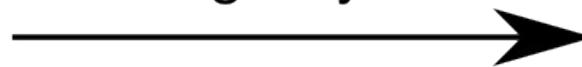
3d RM Synthesis



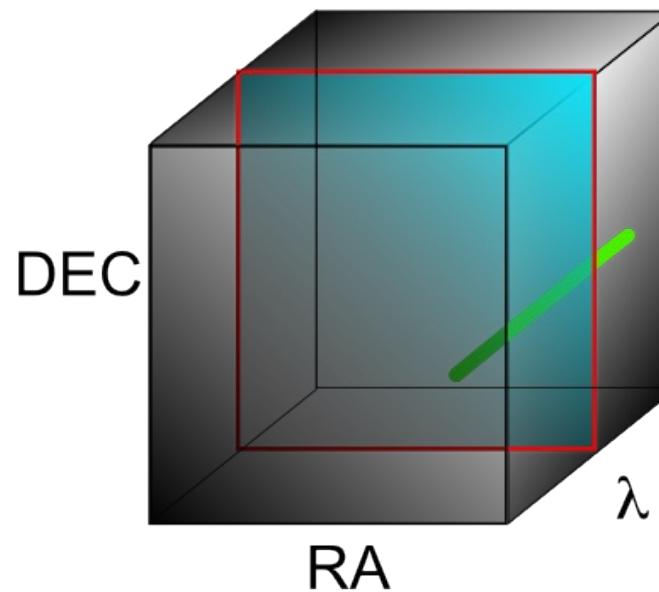
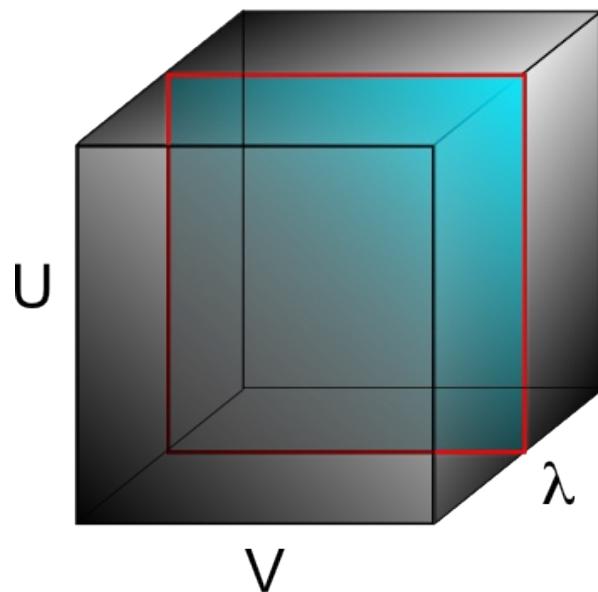
3d RM Synthesis



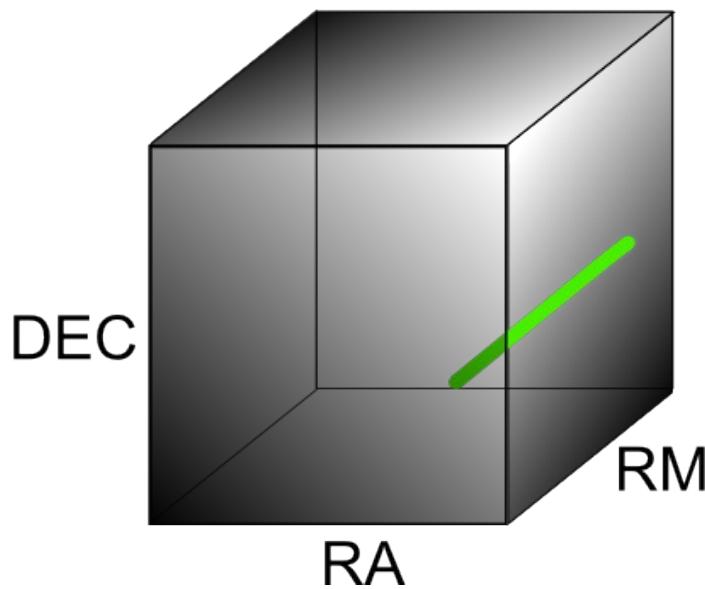
2-d Image Synthesis



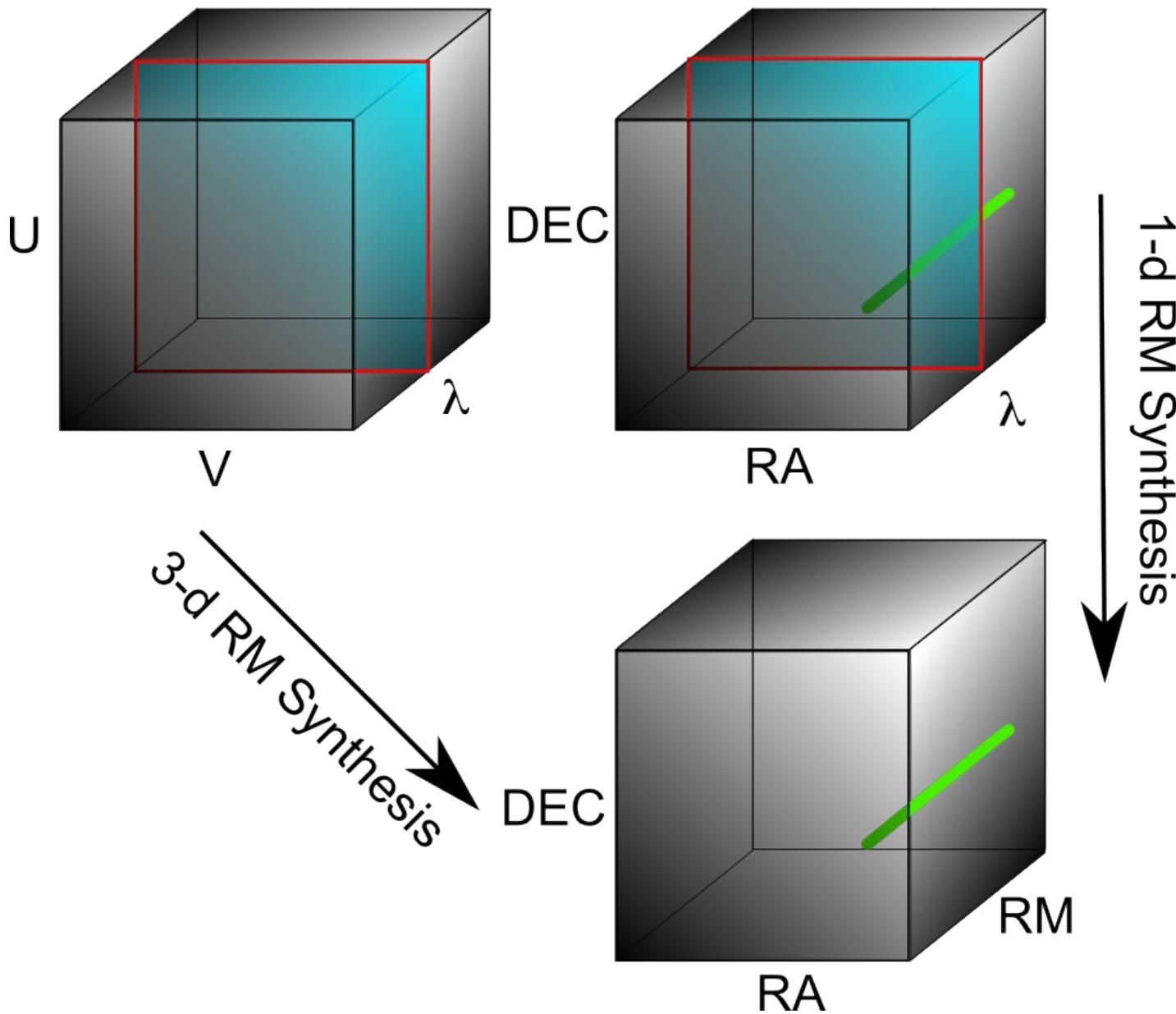
2-d Image Synthesis

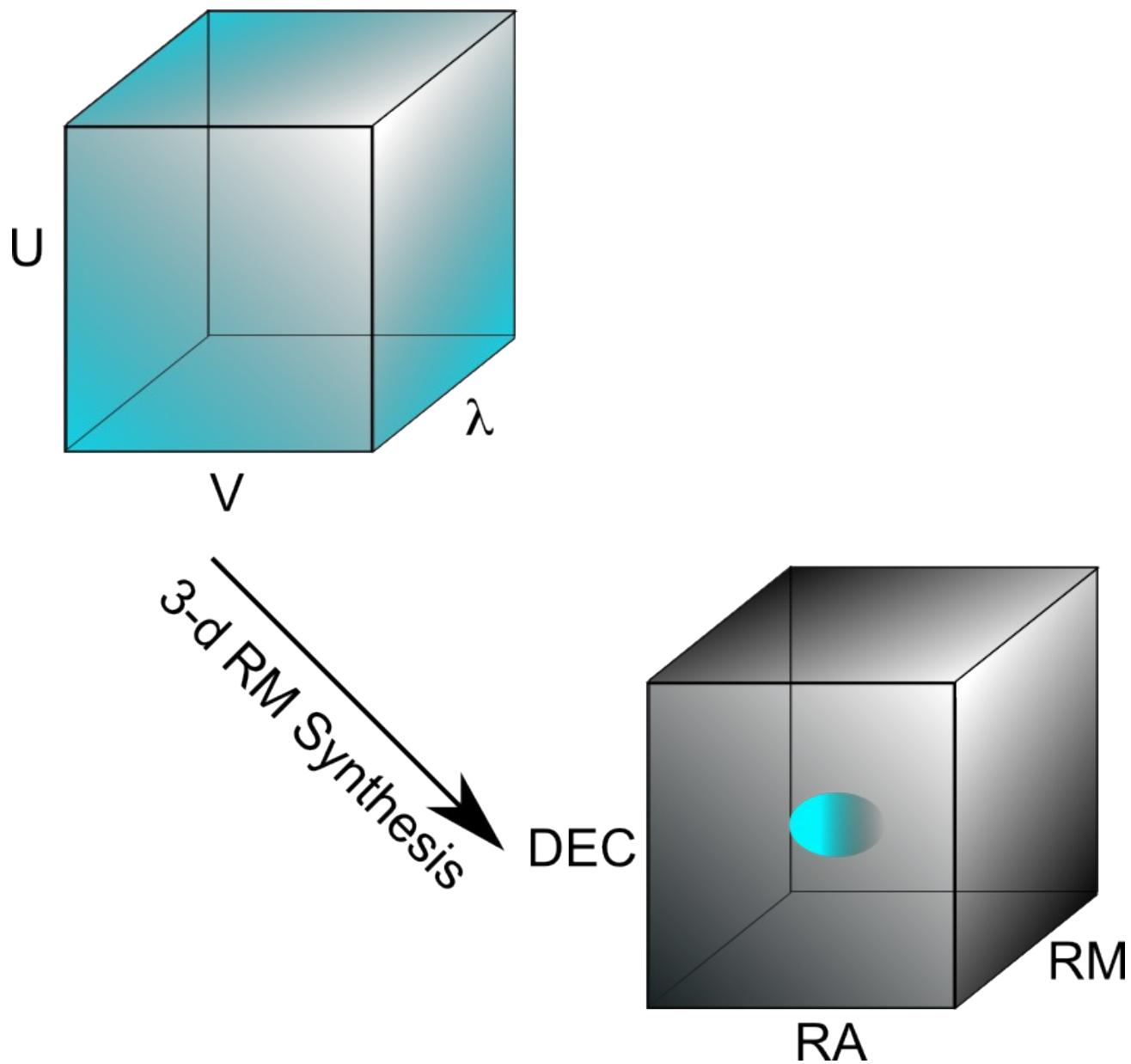


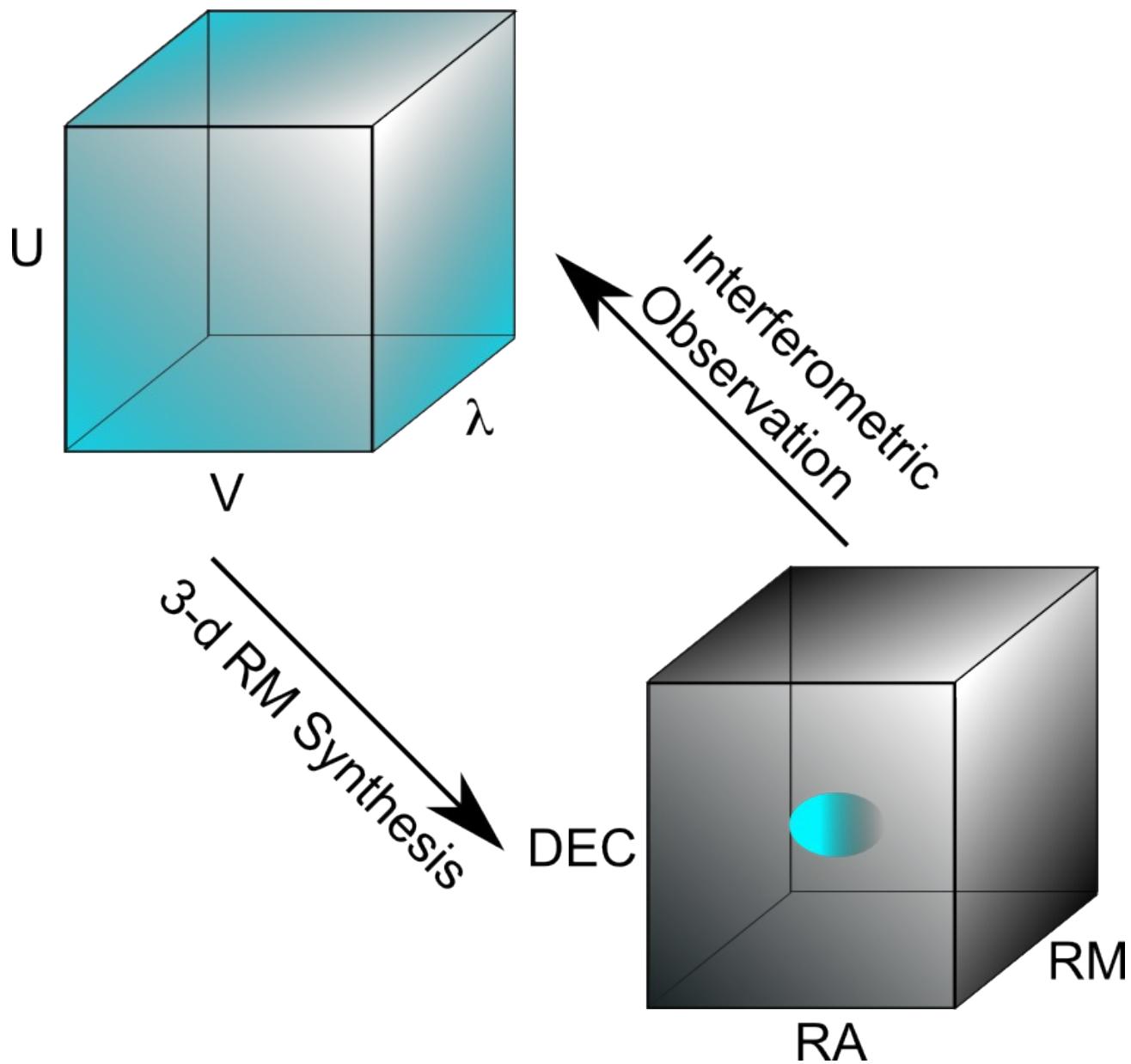
1-d RM Synthesis



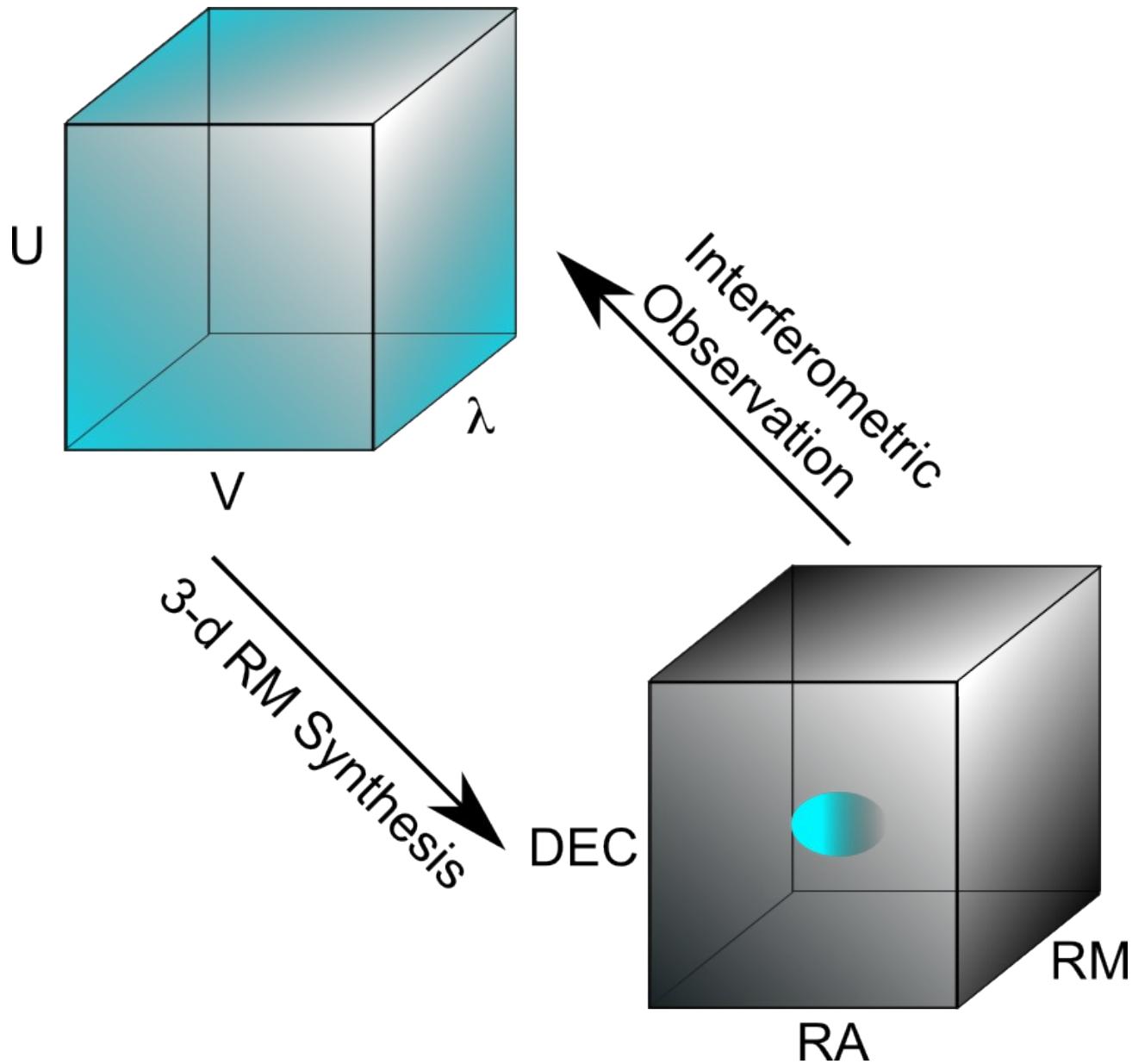
2-d Image Synthesis







information

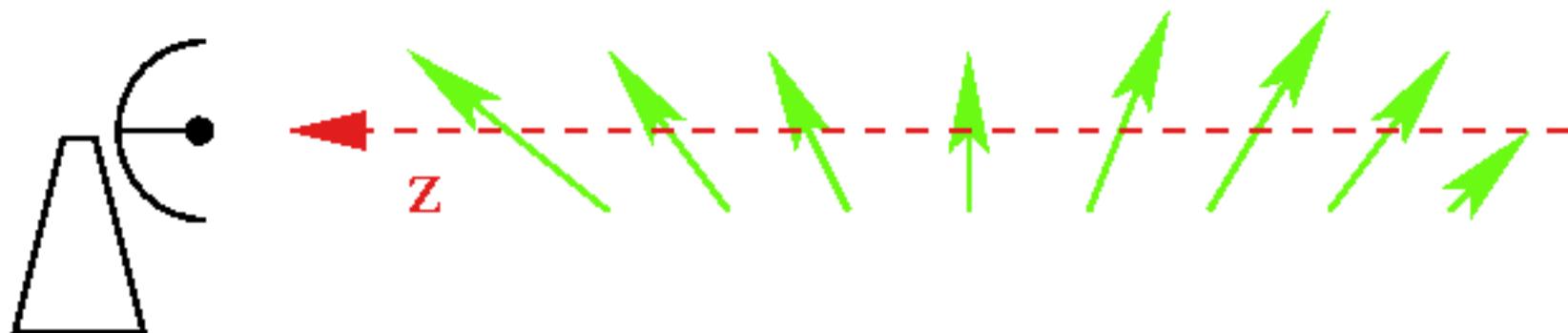


How to obtain information ?

physical signal s

1) **prior knowledge:** $P(s)$

2) **measurement:** $d = R(s) + n, \quad P(d|s)$



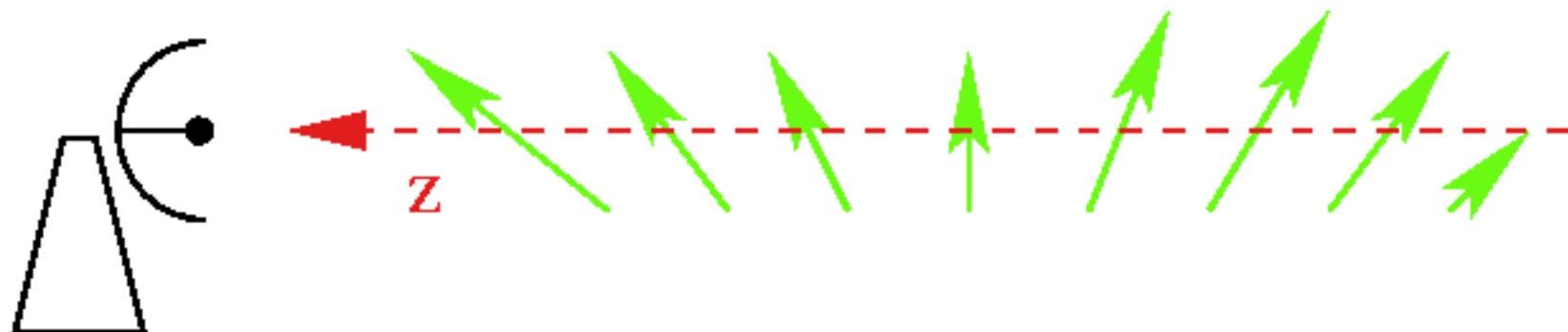
3) **inference:** $P(s|d) = P(d|s) P(s) / P(d), \quad d \rightarrow s$

Wiener filter

physical signal s

1) **prior knowledge:** $P(s) = G(s, S)$ Gaussian

2) **linear measurement:** $d = R s + n, \quad P(n) = G(n, N)$



3) **inference:** $P(s|d) = G(s-m, D)$

$$m = (S^{-1} + R^\dagger N^{-1} R)^{-1} R^\dagger N^{-1} d$$

$$D = (S^{-1} + R^\dagger N^{-1} R)^{-1}$$

Prior information ?

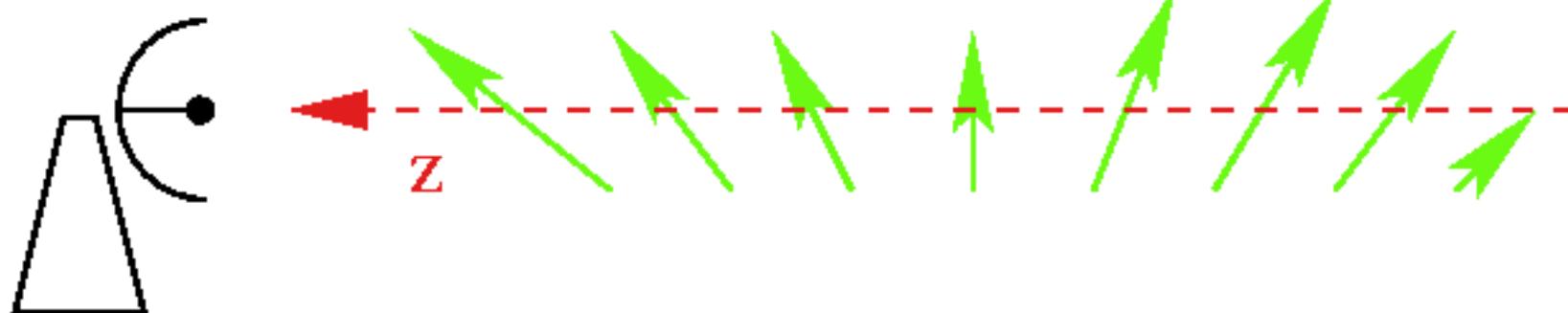
physical signal \mathbf{s}

prior knowledge:

- \mathbf{s} sparse in pixel space
- \mathbf{s} sparse in wavelet space
- \mathbf{s} Gaussian random field
- \mathbf{s} non-Gaussian random field

optimal algorithm:

- CLEAN
- Compressed Sensing
- Wiener Filtering
- Information Field Theory

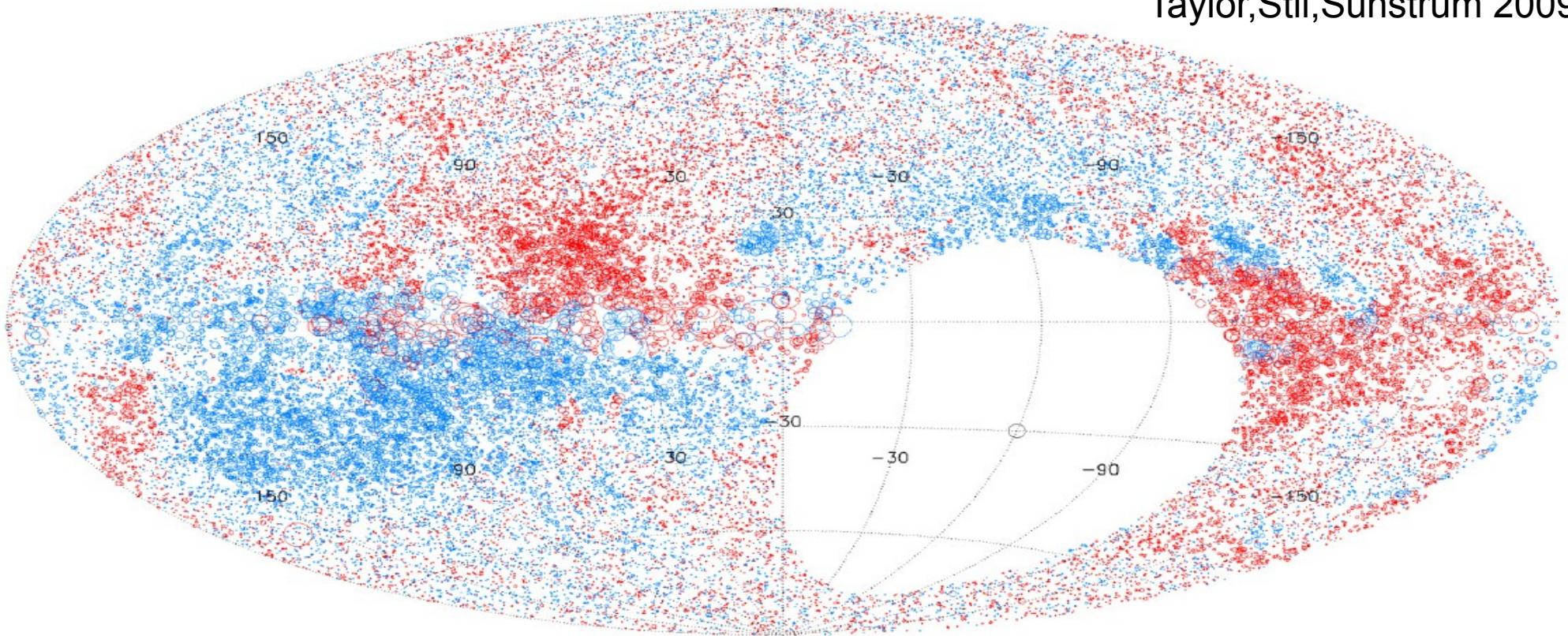


chicken & egg problem:

signal statistics & signal reconstruction
are interdependent

The RM Sky

Taylor,Stil,Sunstrum 2009



37,534 RM extragalactic RM-sources in the northern sky

The RM Sky

Taylor, Stil, Sunstrum 2009

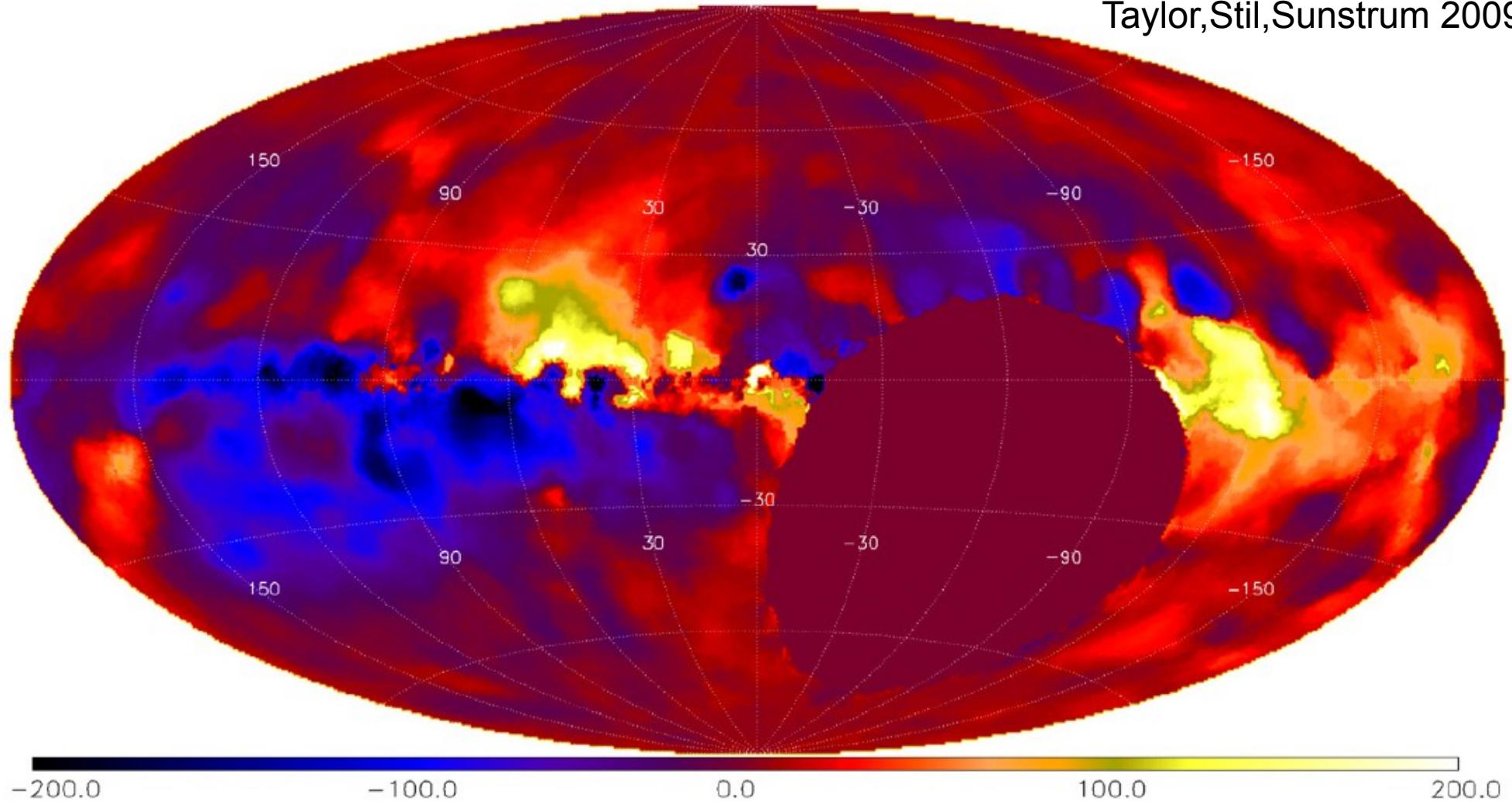
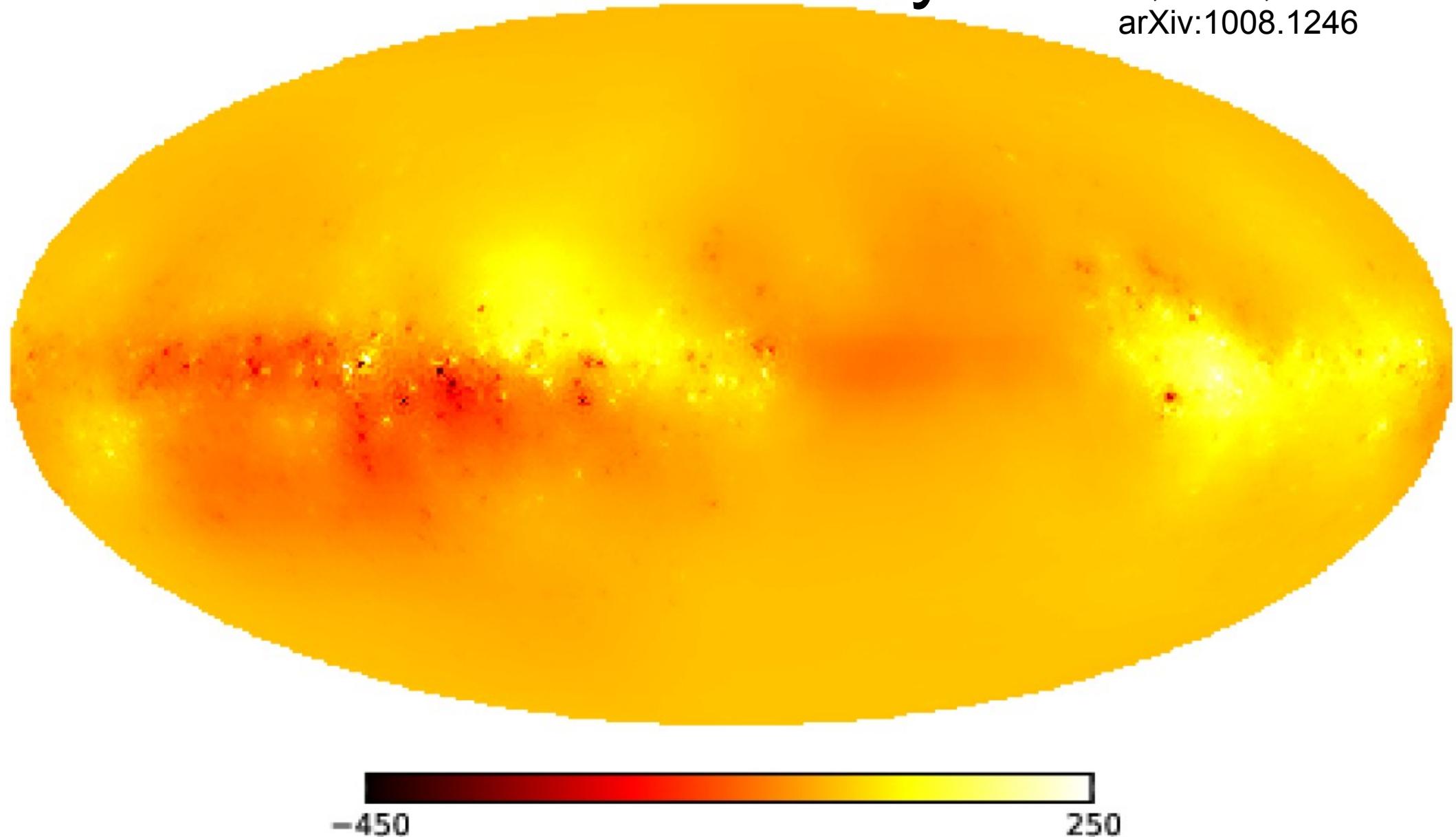


Image of the median value of RM

The RM Sky

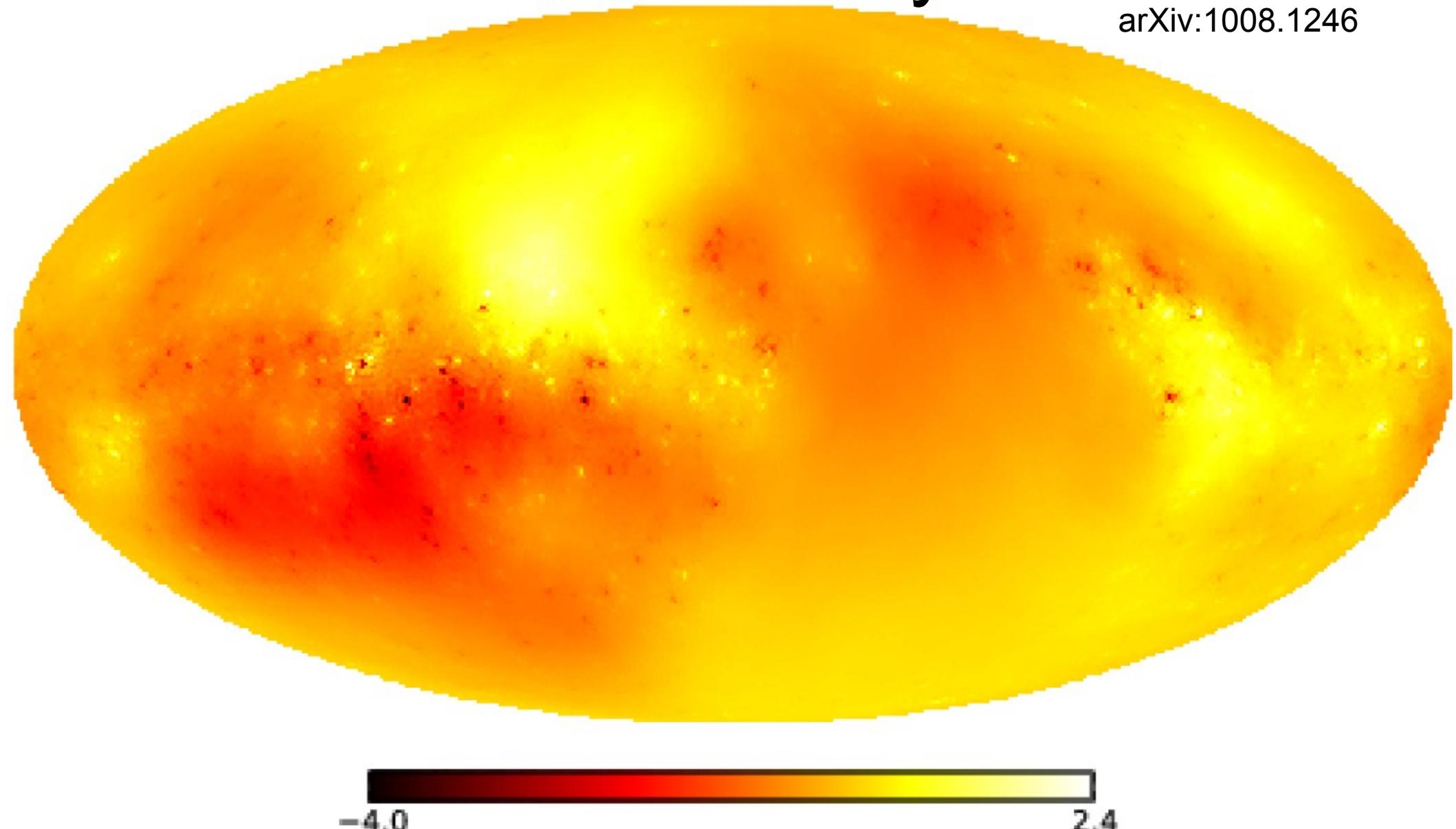
Oppermann, Junklewitz,
Robbers, Enßlin, 2010
arXiv:1008.1246



Wiener filtered RM-map with inferred power spectrum

The RM Sky

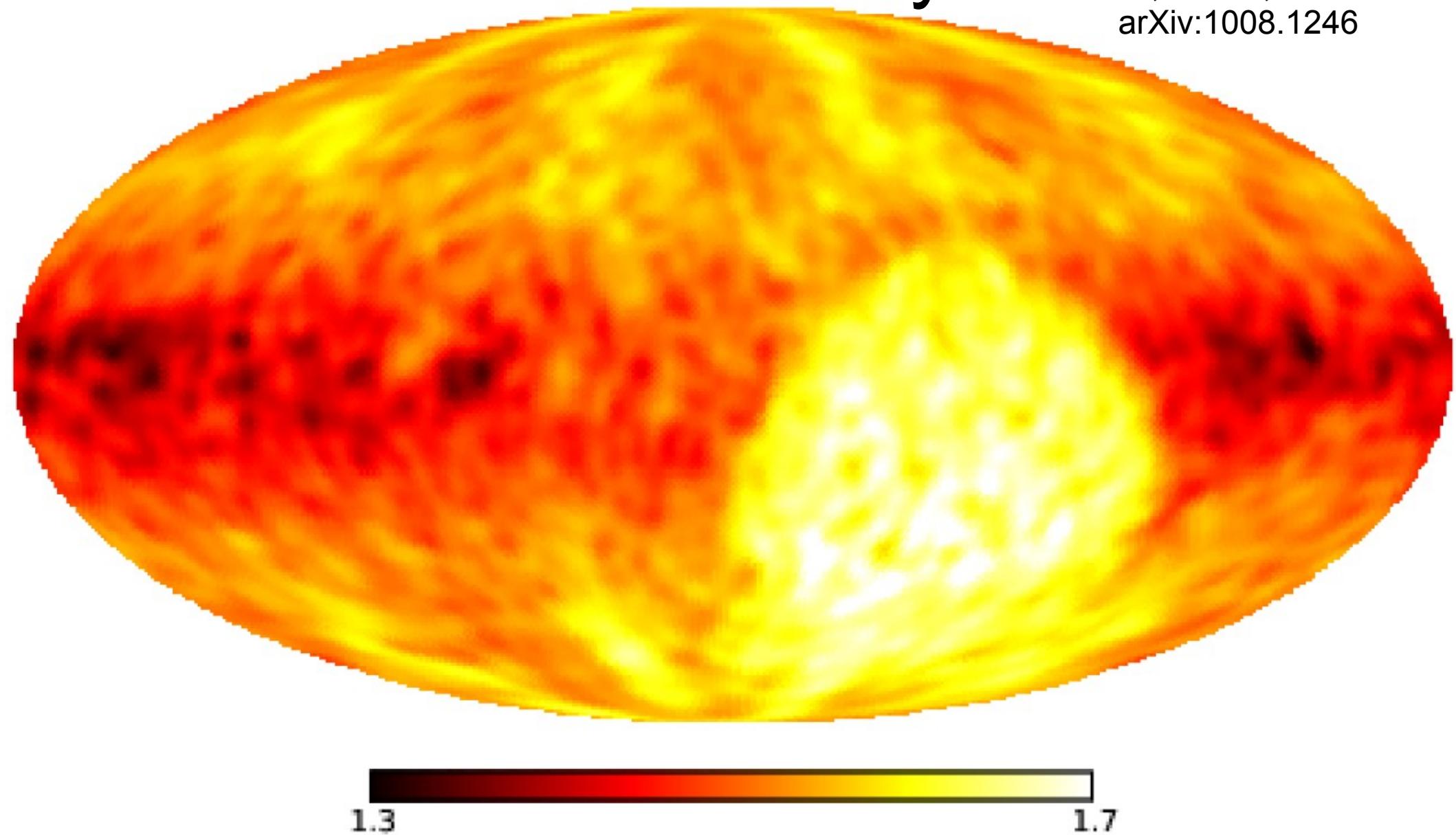
Oppermann, Junklewitz,
Robbers, Enßlin, 2010
arXiv:1008.1246



Wiener filtered RM-map with latitude profile removed

The RM Sky

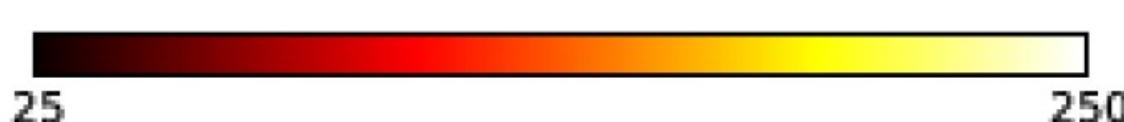
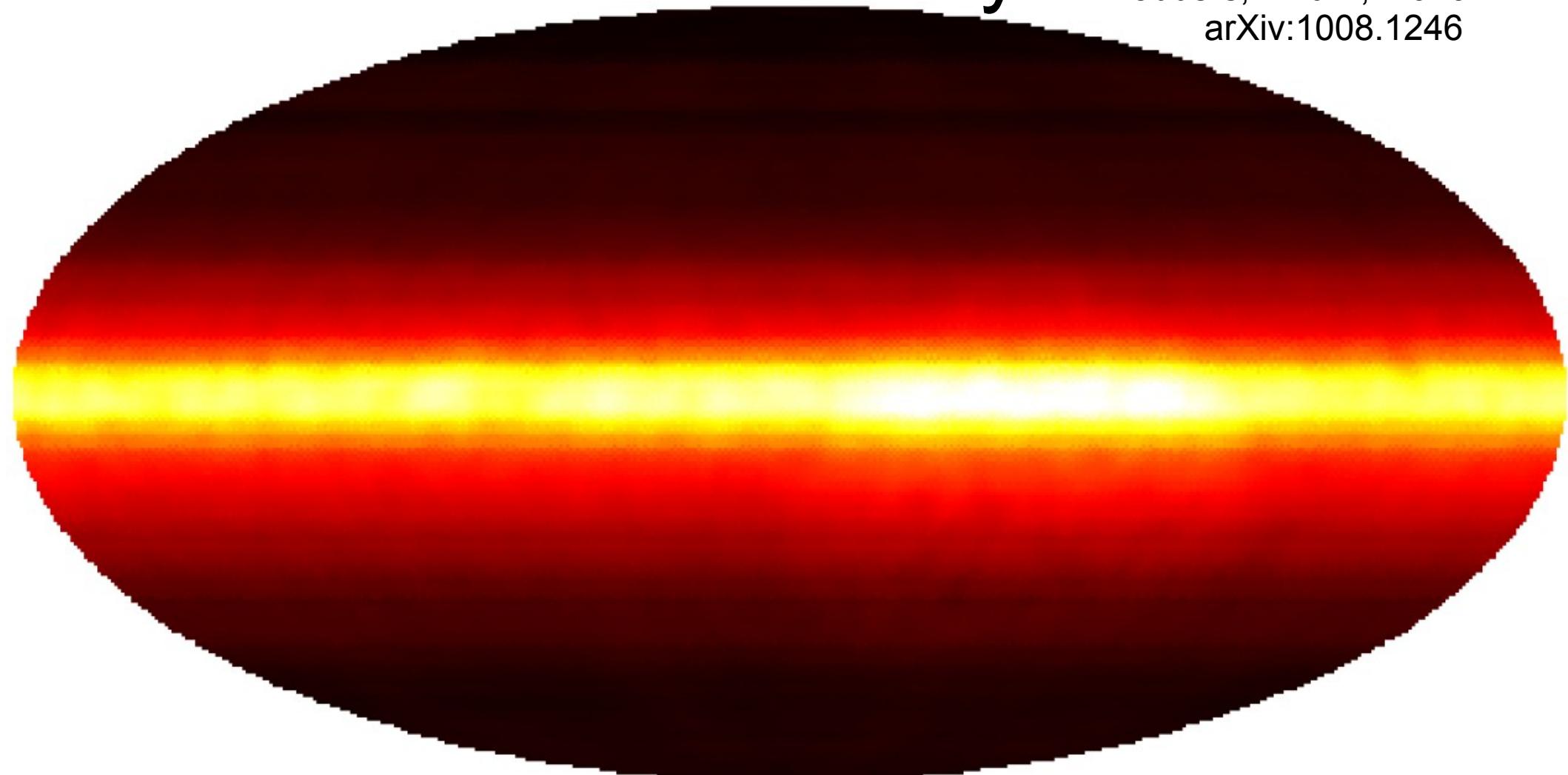
Oppermann, Junklewitz,
Robbers, Enßlin, 2010
arXiv:1008.1246



RM uncertainty map with latitude profile removed

The RM Sky

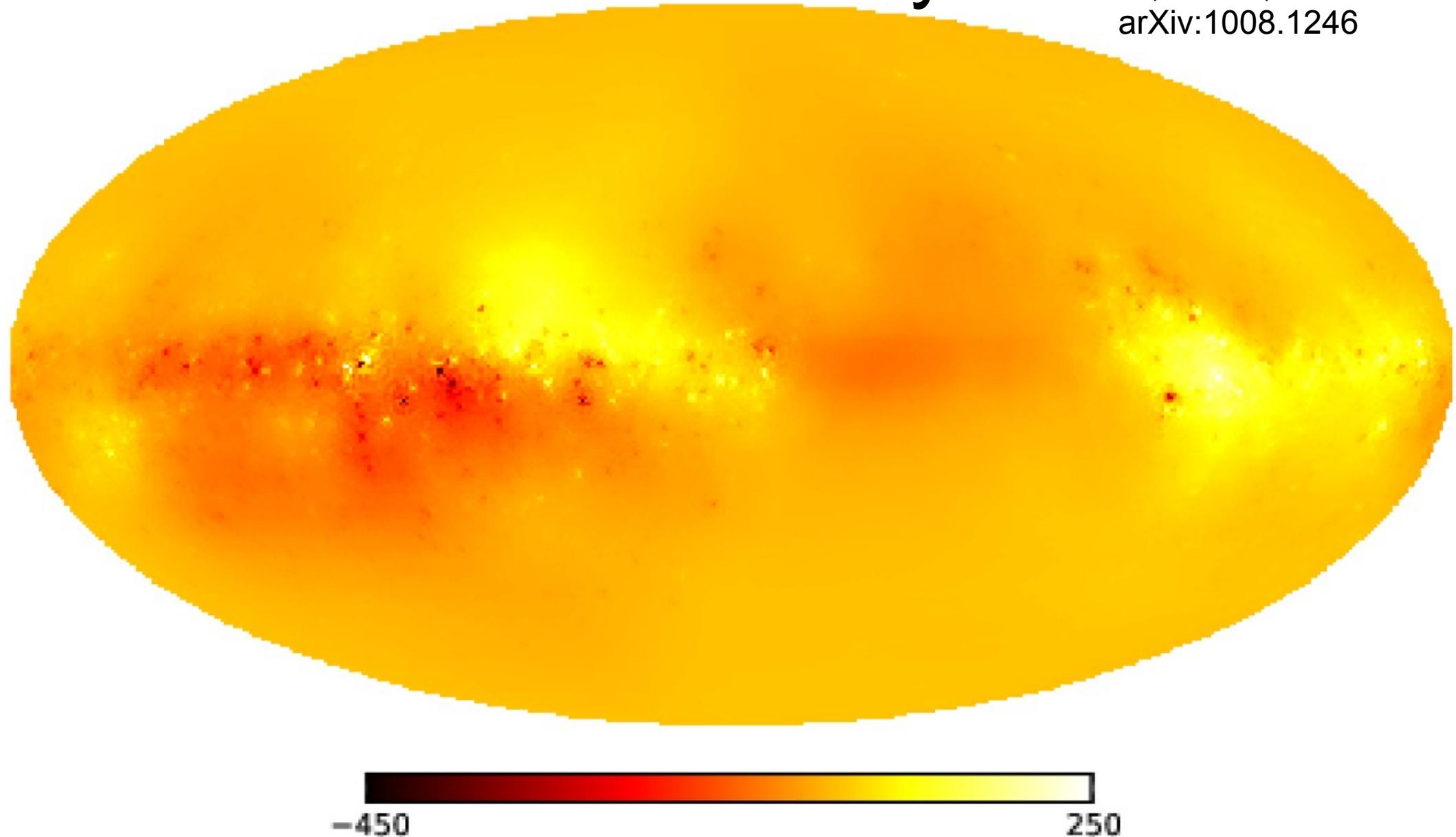
Oppermann, Junklewitz,
Robbers, Enßlin, 2010
arXiv:1008.1246



RM uncertainty map

The RM Sky

Oppermann, Junklewitz,
Robbers, Enßlin, 2010
arXiv:1008.1246



Wiener filtered RM-map with inferred power spectrum

Summary

A8: research tools for the research unit

magnetic field statistics:

Hammurabi – simulation of galactic observables

REALMAF – magnetic power spectrum inference

Litmus Test – magnetic helicity detection

Faraday tomography:

1d & 3d RM Synthesis – in development

[implementing CLEAN, Wiener Filter, ...]







Why Information Field Theory ?

inverse problem => Information Theory
spatially distributed quantity => Field Theory

Information Field Theory

Translation:

inference problem → statistical field theory

$$P(s|d) = \frac{P(d|s) P(s)}{P(d)} \equiv \frac{1}{Z} e^{-H[s]}$$

Dictionary:

log-Posterior	=	negative Hamiltonian
Evidence	=	partition function Z
Wiener variance	=	information propagator
noise weighted data	→	information source
inference algorithms	←	Feynman diagrams
maximum a Posteriori	=	classical solution
uncertainty correct.	=	loop corrections
Shannon information	=	negative entropy

