Luminous buried AGNs in the local universe ULIRGs

origin of galaxy down-sizing ?

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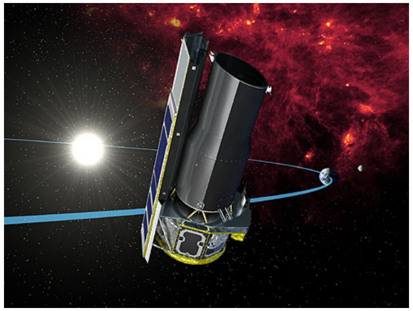
Subaru

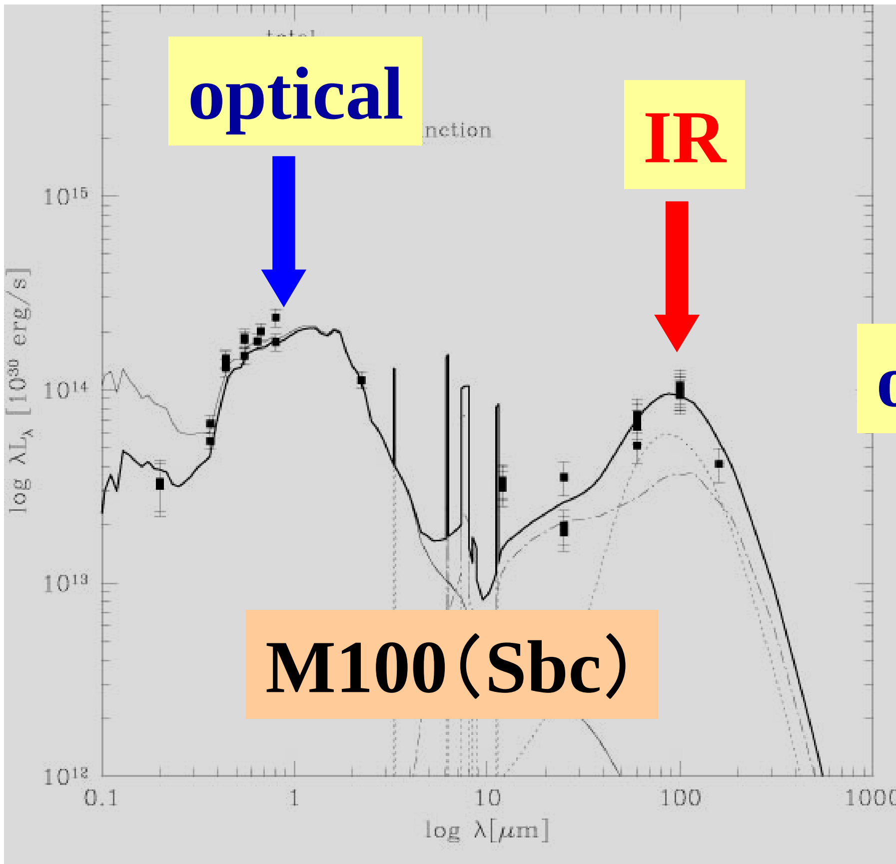


Spitzer

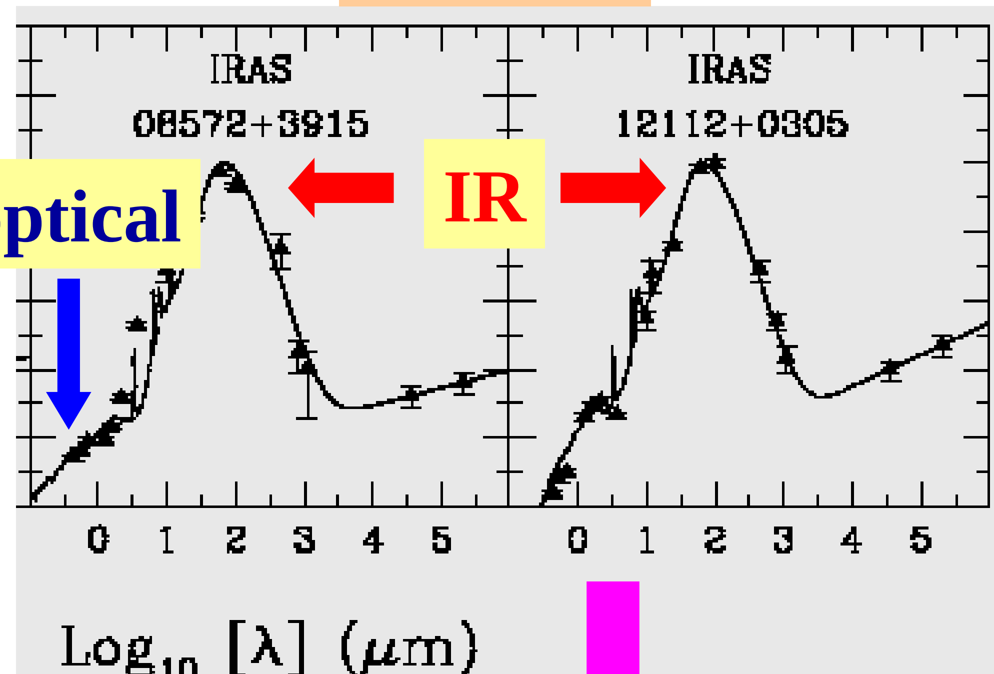
AKARI

Ultraluminous infrared galaxies（ULIRGs）

L(IR) > 10^12 Lsun (Normal spiral ~ 10^10 Lsun)



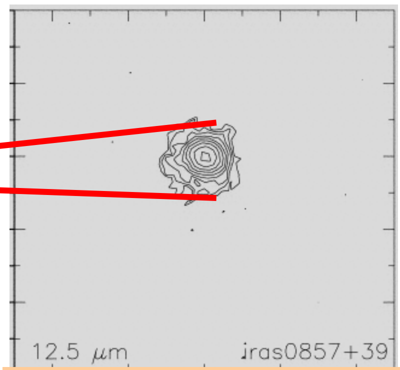
ULIRGs



Luminous energy source is hidden behind dust

ULIRGs

optical



IR(12um)

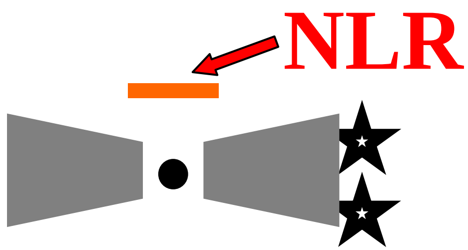
Soifer et al. 2000

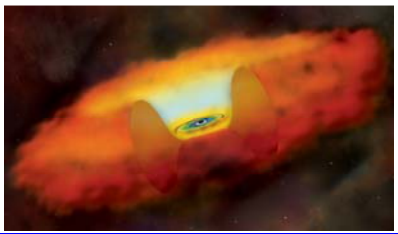
Compact cores (<500pc) are energetically dominant

Very compact starburst

or AGN ?

AGNs in ULIRGs are buried

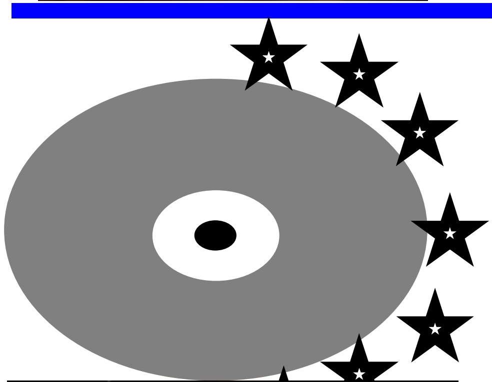


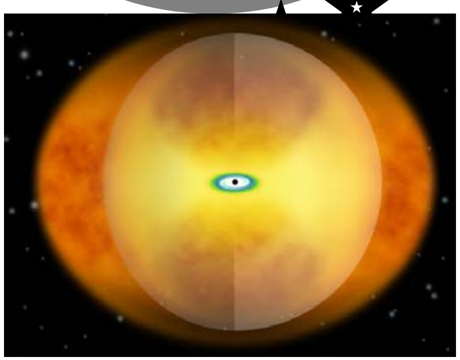


AGNs obscured by torus-shaped dust

Sy2

Detectable via optical spectroscopy





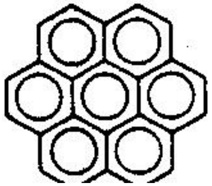
ULIRGs have a large amount of nuclear gas and dust

Buried AGNs are elusive

70% ULIRGs = non-Sy

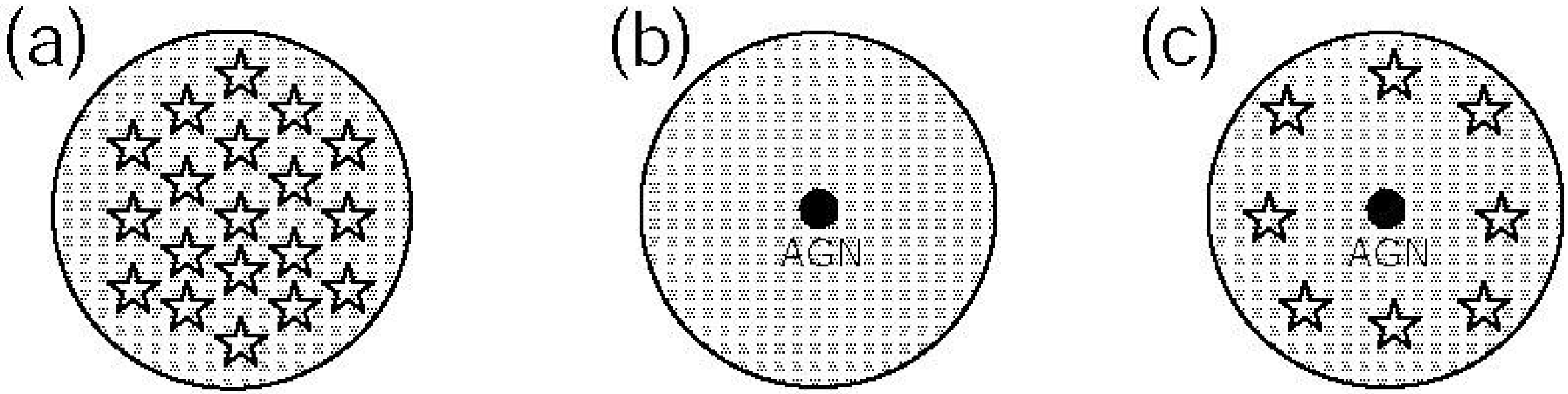
1. Infrared spectral shape

PAH



PAHs are excited in starburst PDRs but destroyed near an AGN

Starburst(SB) | **Buried AGN** | **composite**

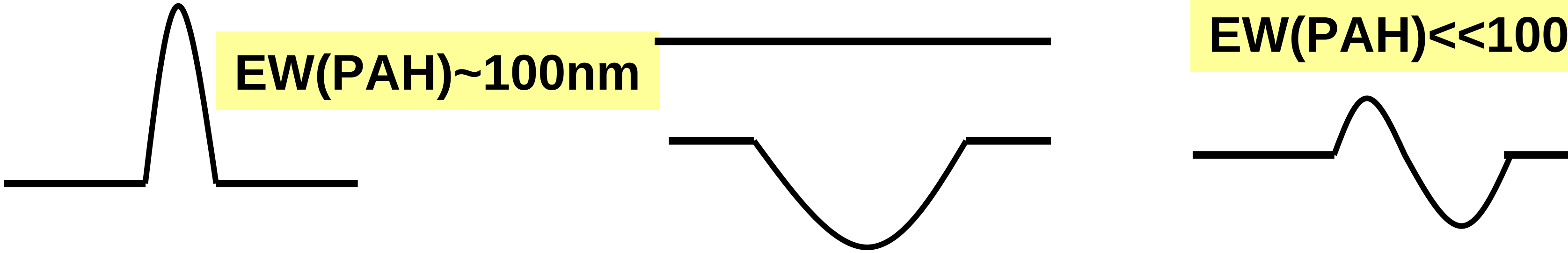


3.3um PAH

EW(PAH)~100nm

featureless

EW(PAH)<<100nm

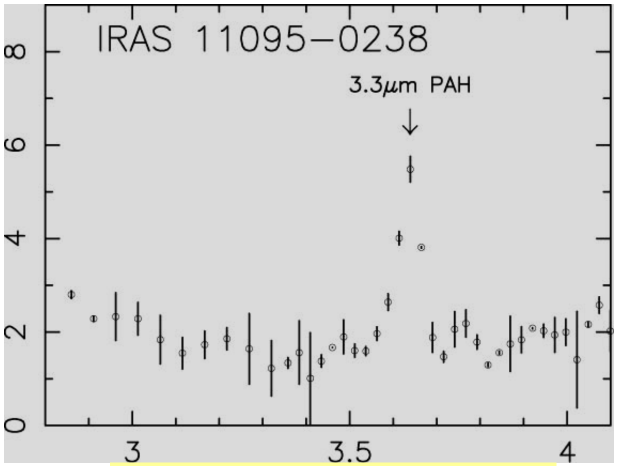


3.4um/3.1um

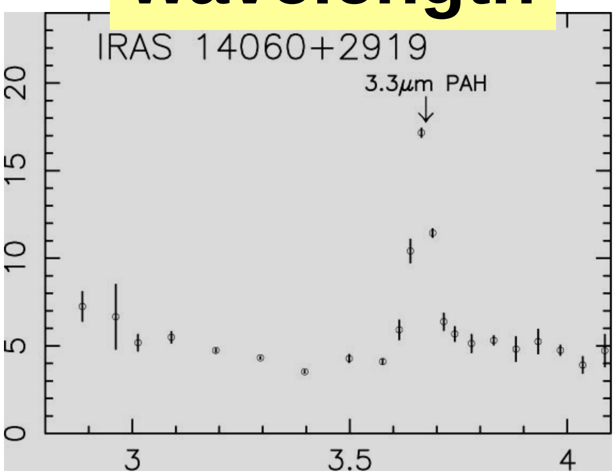
3-4 um

Subaru

Starburst(SB)

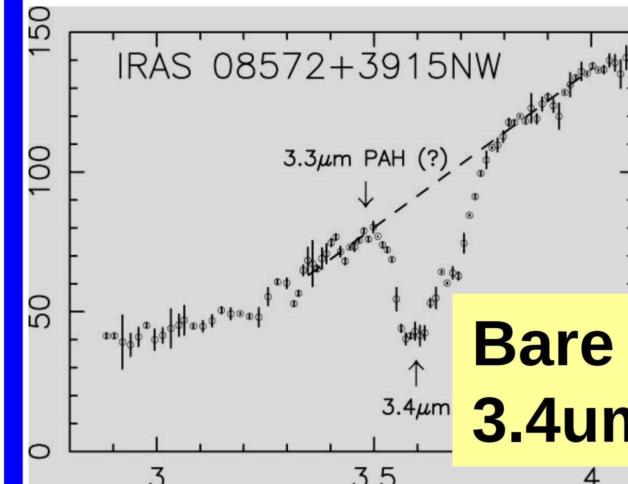


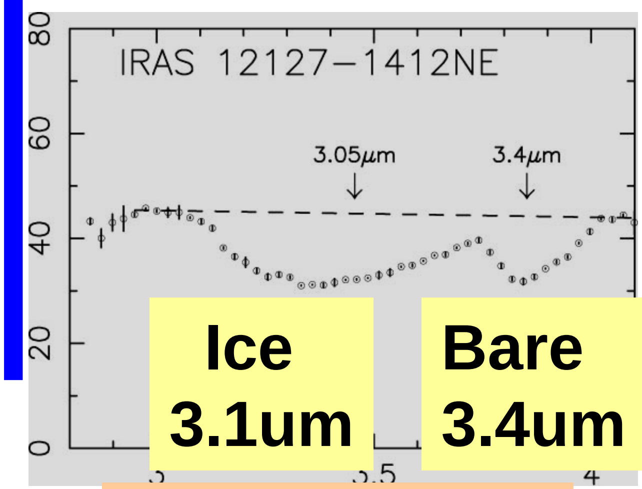
wavelength



Strong PAH

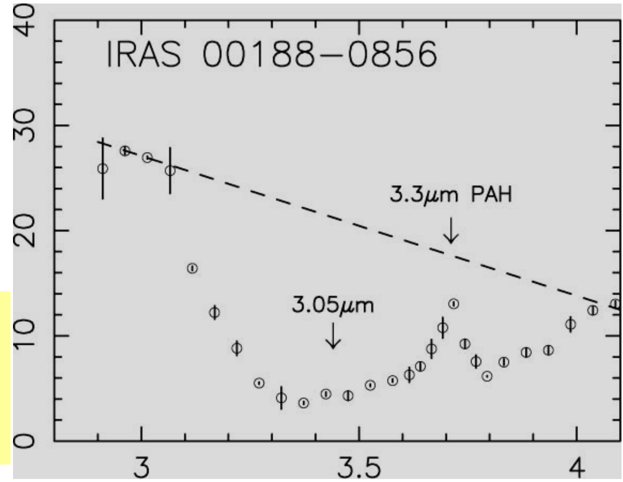
Buried AGN

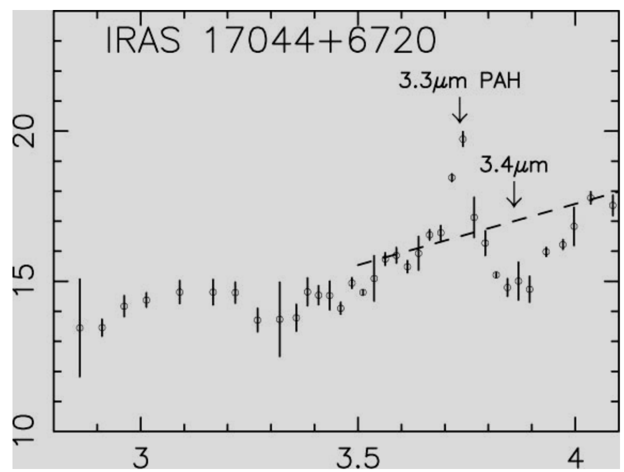




Abs. feature

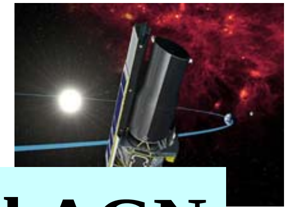
AGN/SB composite



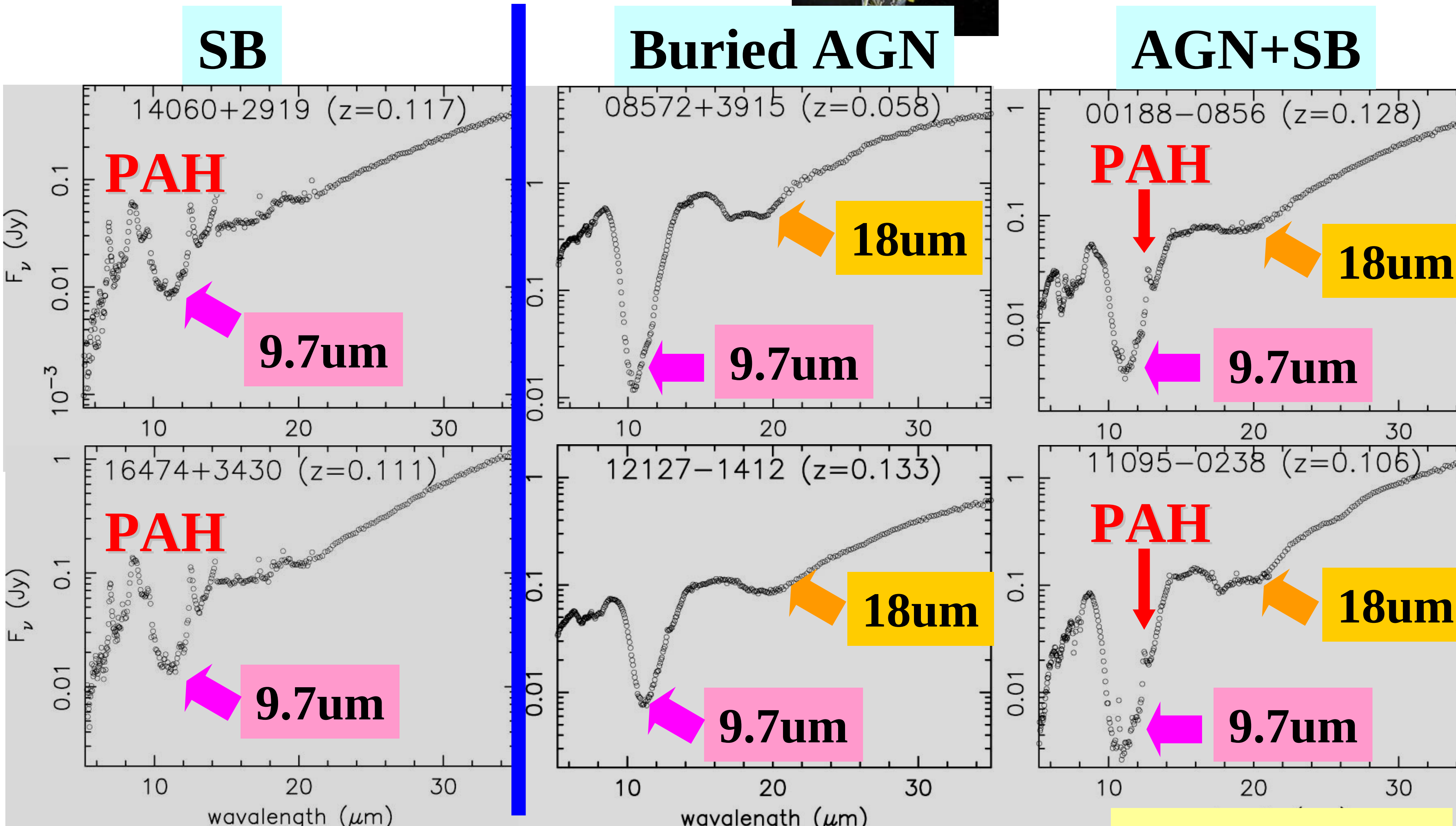


Low EW(PAH)

5-35 um



Spitzer GO1

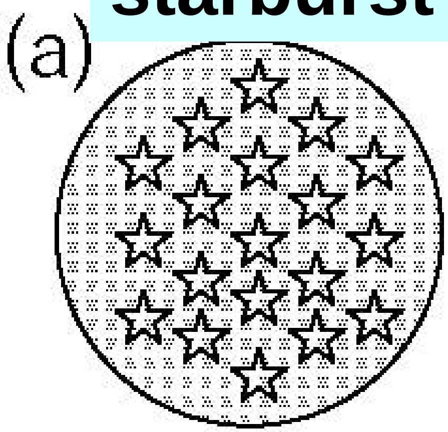


PAH strong

PAH weak
Silicate Abs. strong

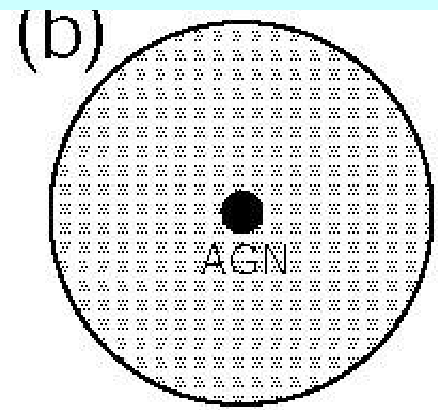
2. Dust absorption feature strength

starburst



ULIRG core
<500pc

Buried AGN



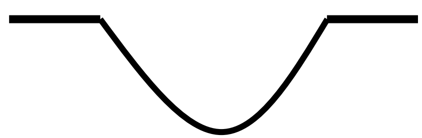
Mixed dust model

$$\frac{1-\exp(-\mathrm{Tau})}{\mathrm{Tau}}$$

Foreground screen dust model

$$\exp(-\mathrm{Tau})$$

Dust absorption feature: weak



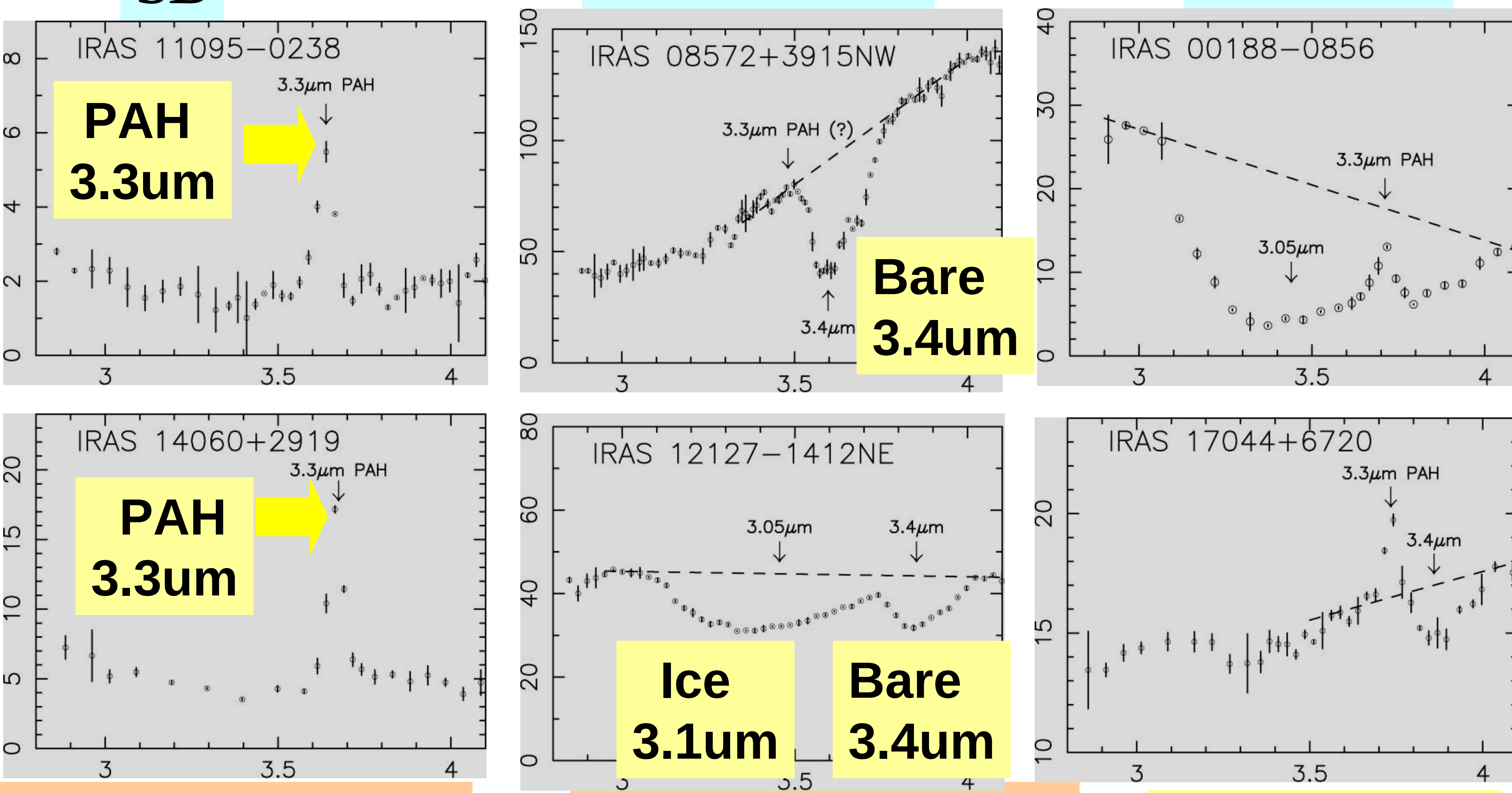
strong

Tau(3.1) < 0.3 Tau(9.7) < 1.7

Tau(3.4) < 0.2 (Imanishi & Maloney 2003 ApJ 588 165

3-4 um

SB | **Buried AGN** | **AGN+SB**

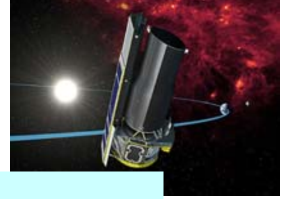


PAH strong (SB):
Dust abs. weak

PAH weak (AGN):
Dust abs. strong

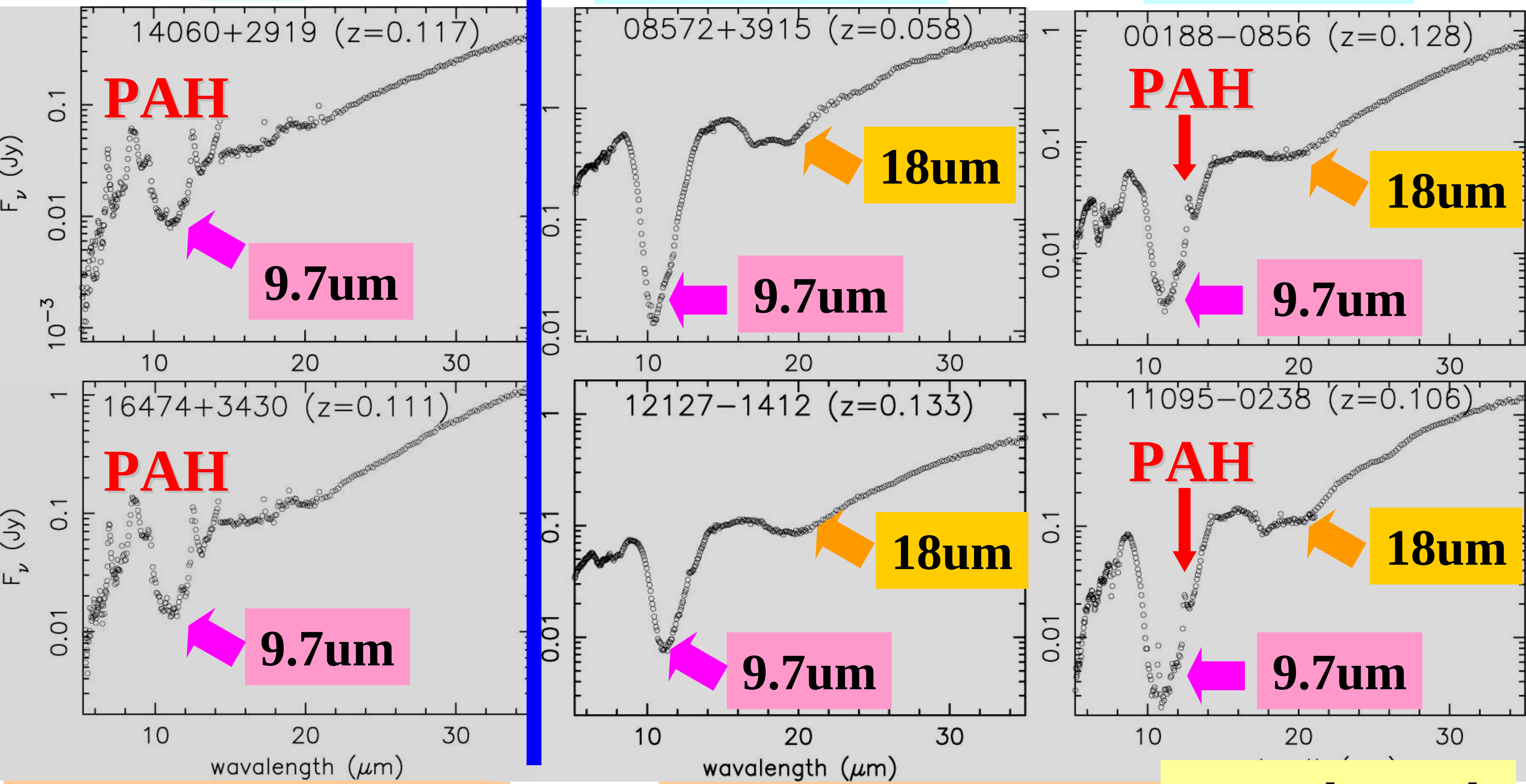
wavelength

5-35 um



Spitzer GO1

SB | **Buried AGN** | **AGN+SB**

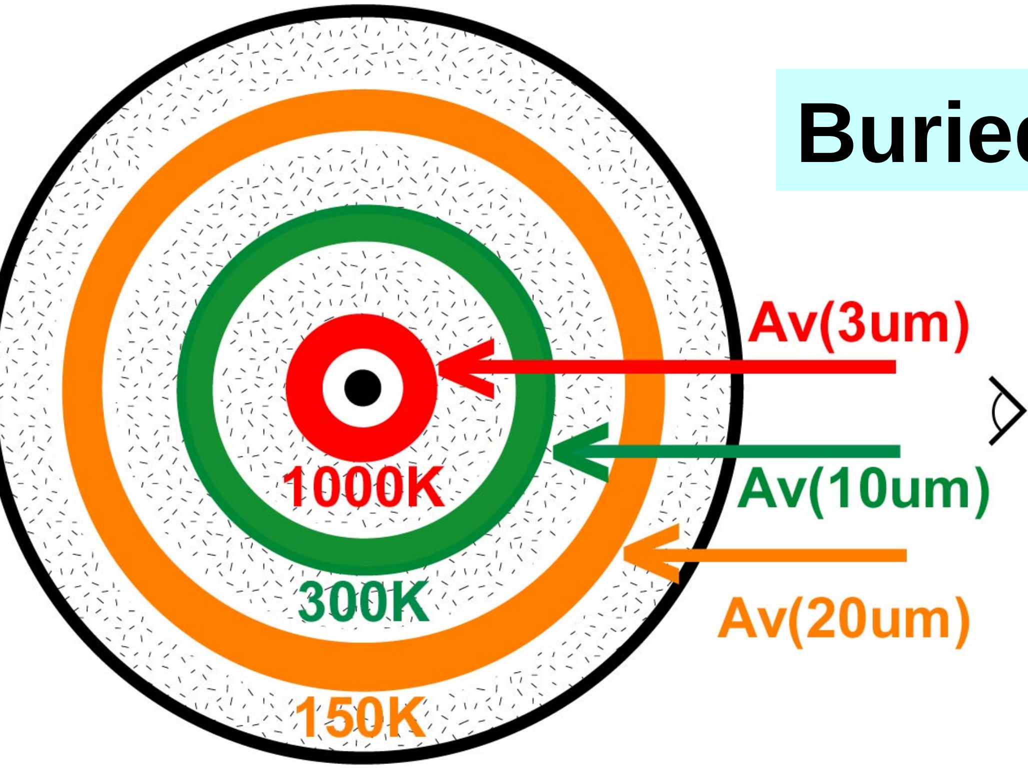


wavelength

PAH strong :
Silicate Abs. weak

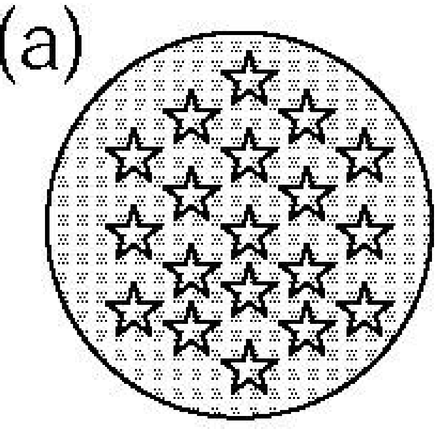
PAH weak:
Silicate Abs. strong

3. Dust temperature gradient



Buried AGN

Av(3um) > Av(10um) > Av(20um)

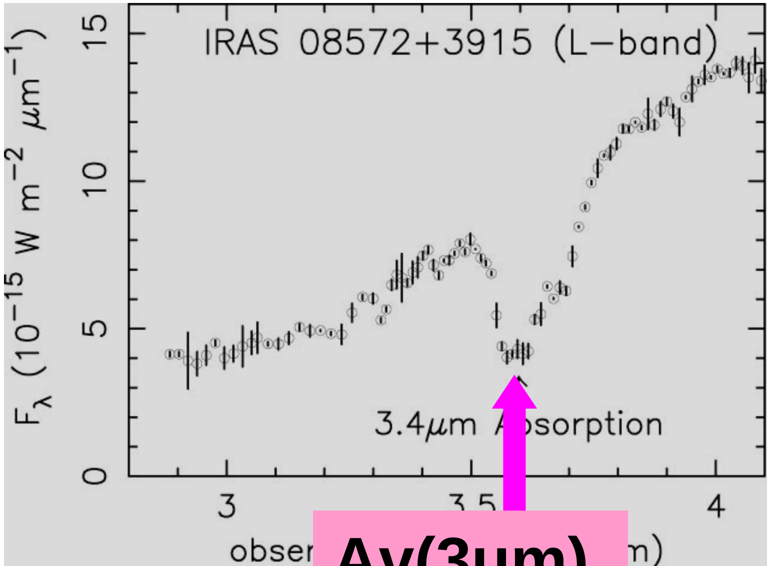


dust in edge-on host

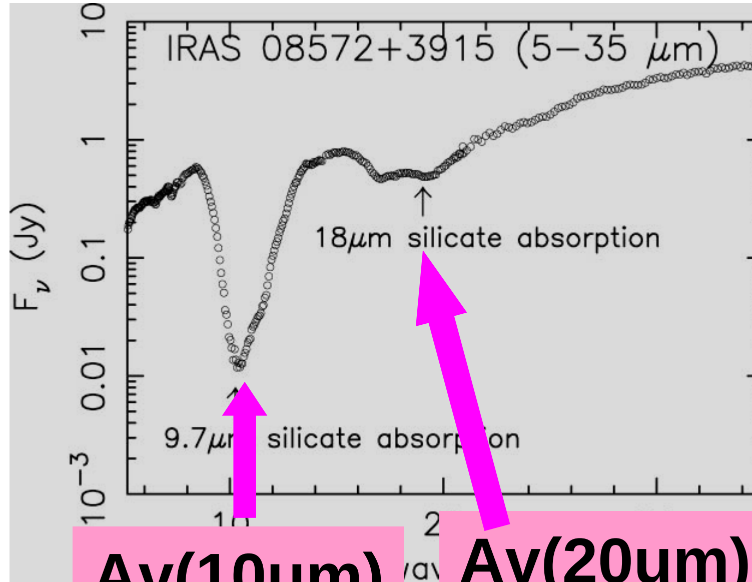
Starburst

Av(3um) =< Av(10um) =< Av(20um)

How to detect T-gradient ?



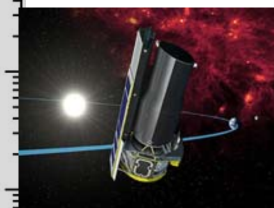
Av(3um)
~110mag



Av(10um)
~40mag

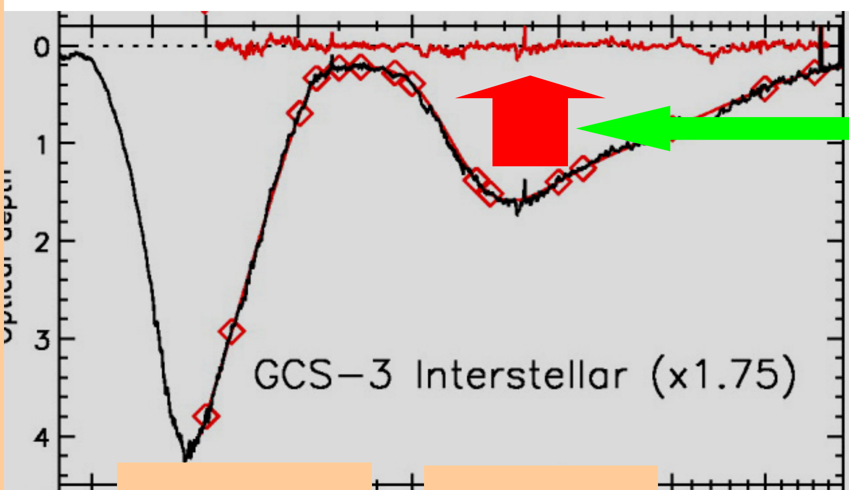
Av(20um)
~20mag

Subaru



Spitzer

Optical depth

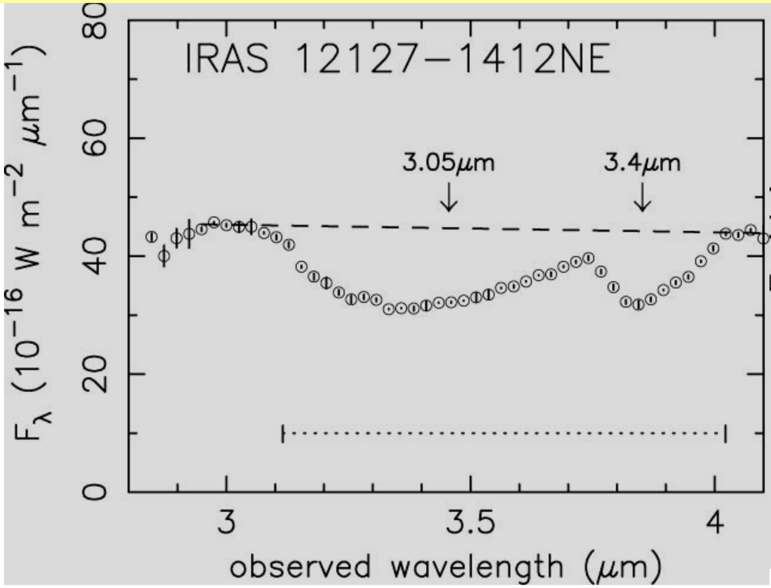


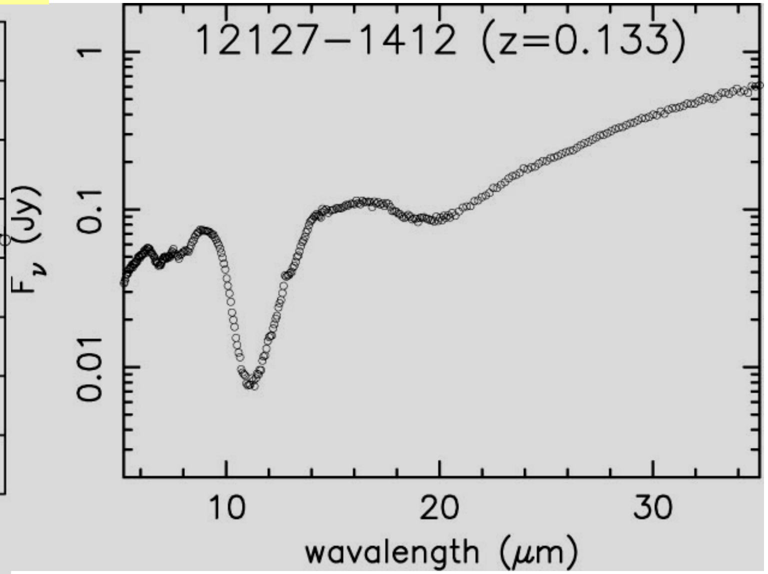
9.7um

18um

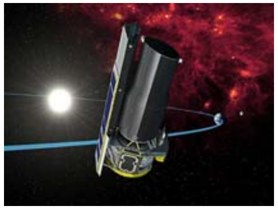
Dust temperature gradient

Strong T-gradient (II)

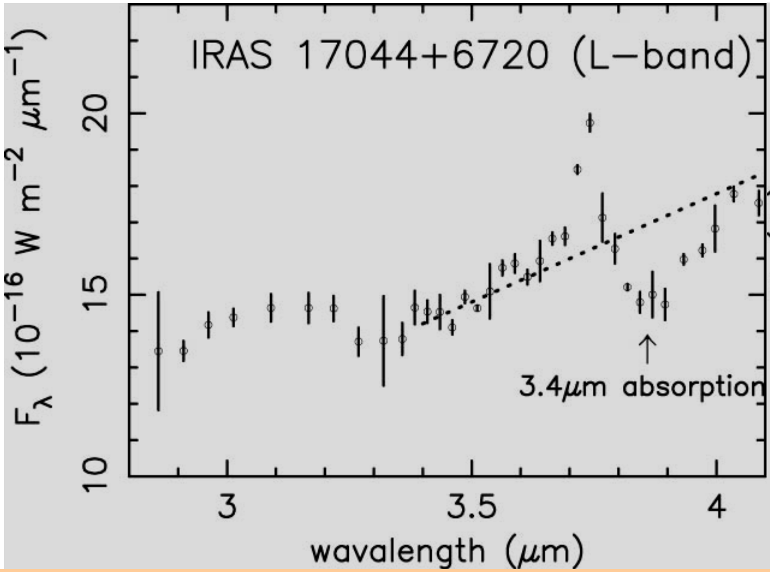


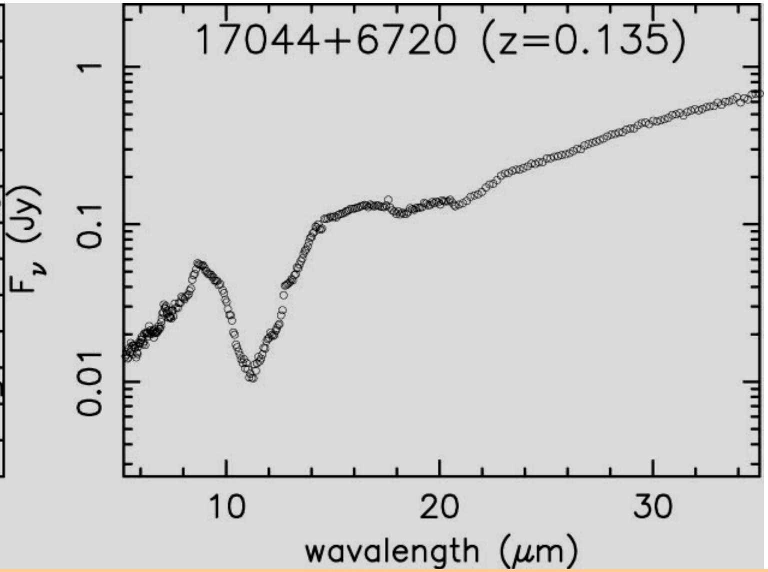


Subaru



Spitzer

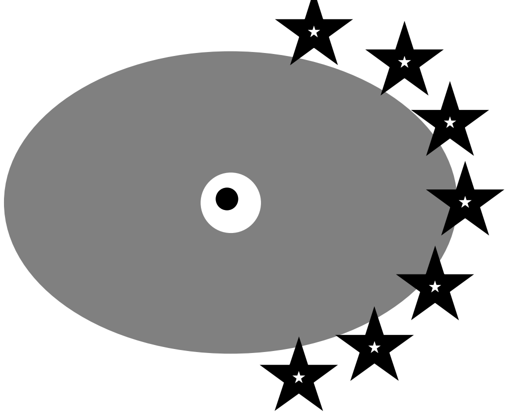




Strong abs ULIRGs -> often show T-gradient

Results

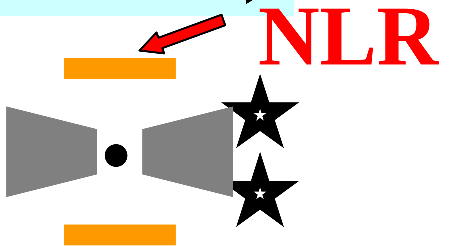
nearby(z<0.15)



Opical non-Seyfert ULIRGs

Luminous buried AGNs = 30-50%

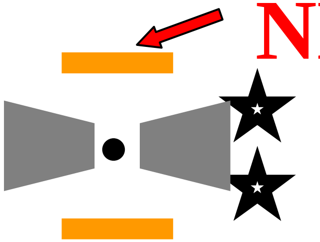
30% ULIRGs = optical Sy (AGN + torus)

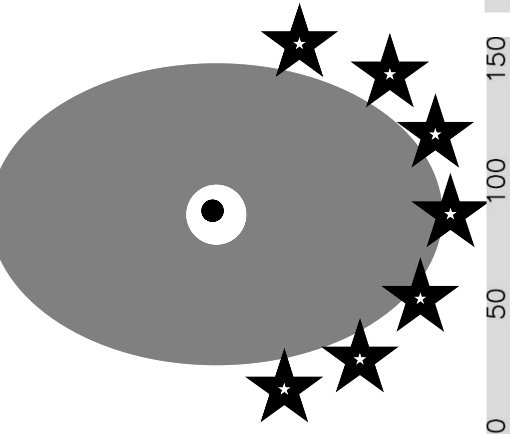


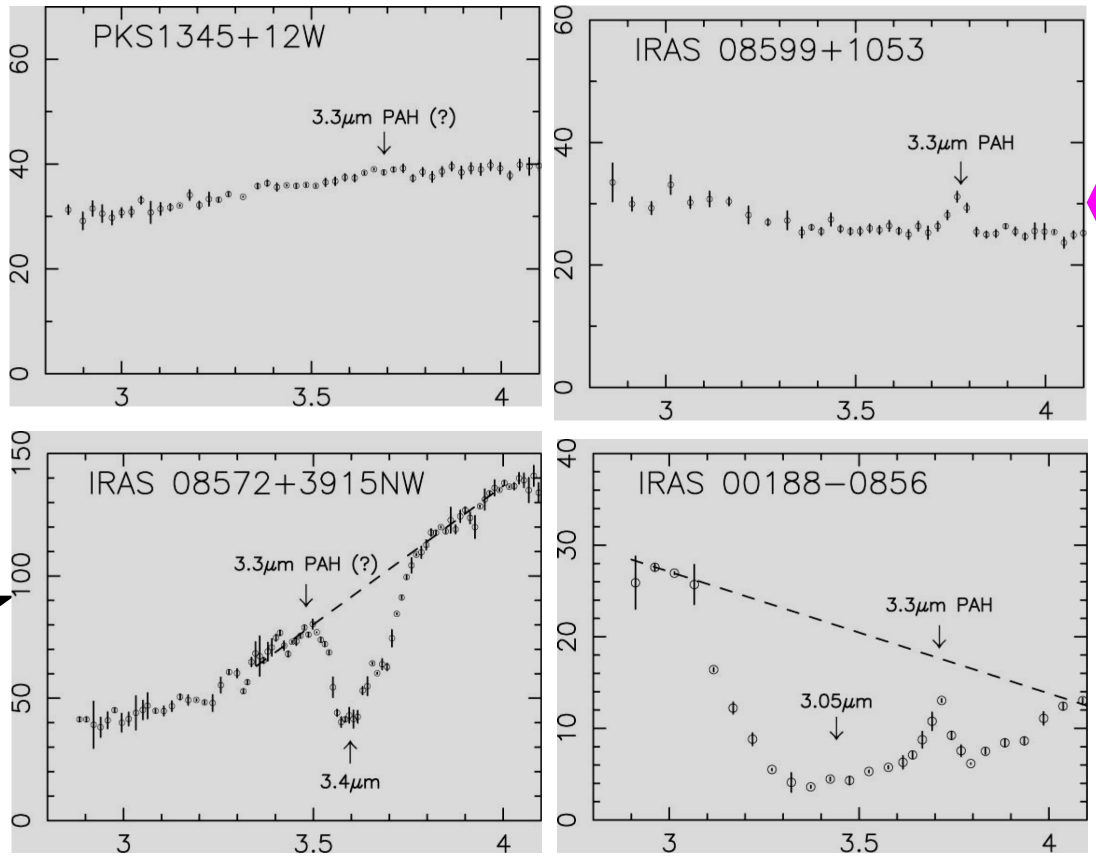
>50% ULIRGs = luminous AGN

Our line-of-sight obscuration: Non-Sy >> Sy2

NLR







Sy2: Abs weak

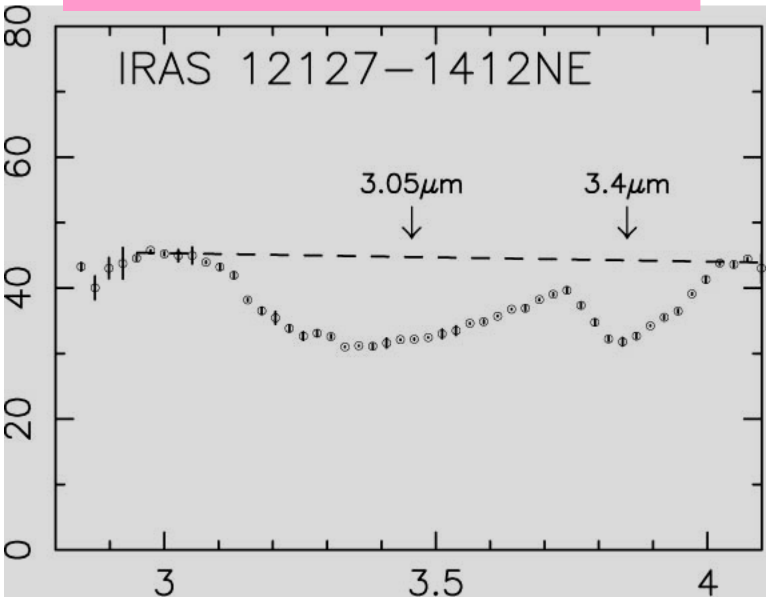
Non-Sy: strong

L(intrinsic AGN) ~L(IR)

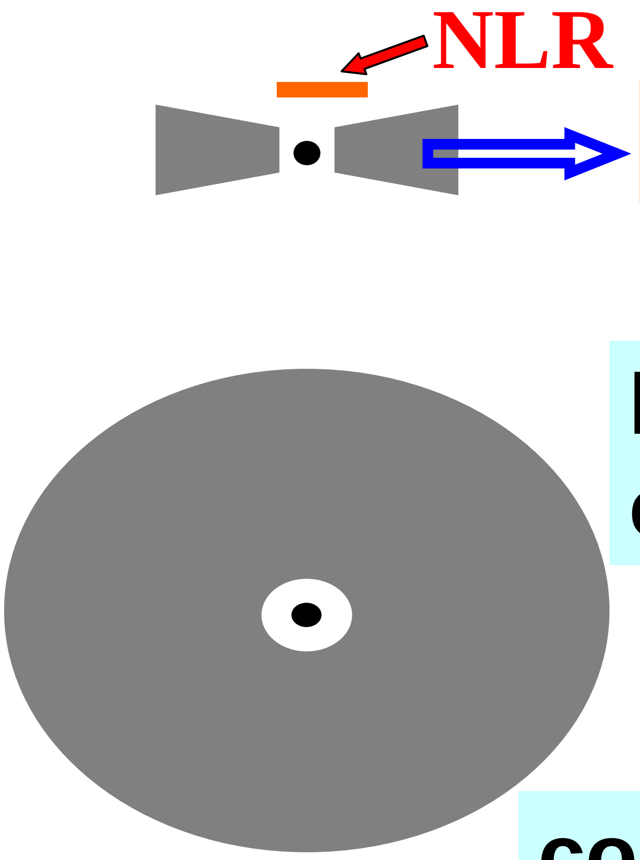
Amount of nuclear dust: Non-Sy >> Sy2

Buried AGNs: both warm/cool FIR colors

pure buried AGN



F25/F60=0.16(cool)



warm

larger dust column

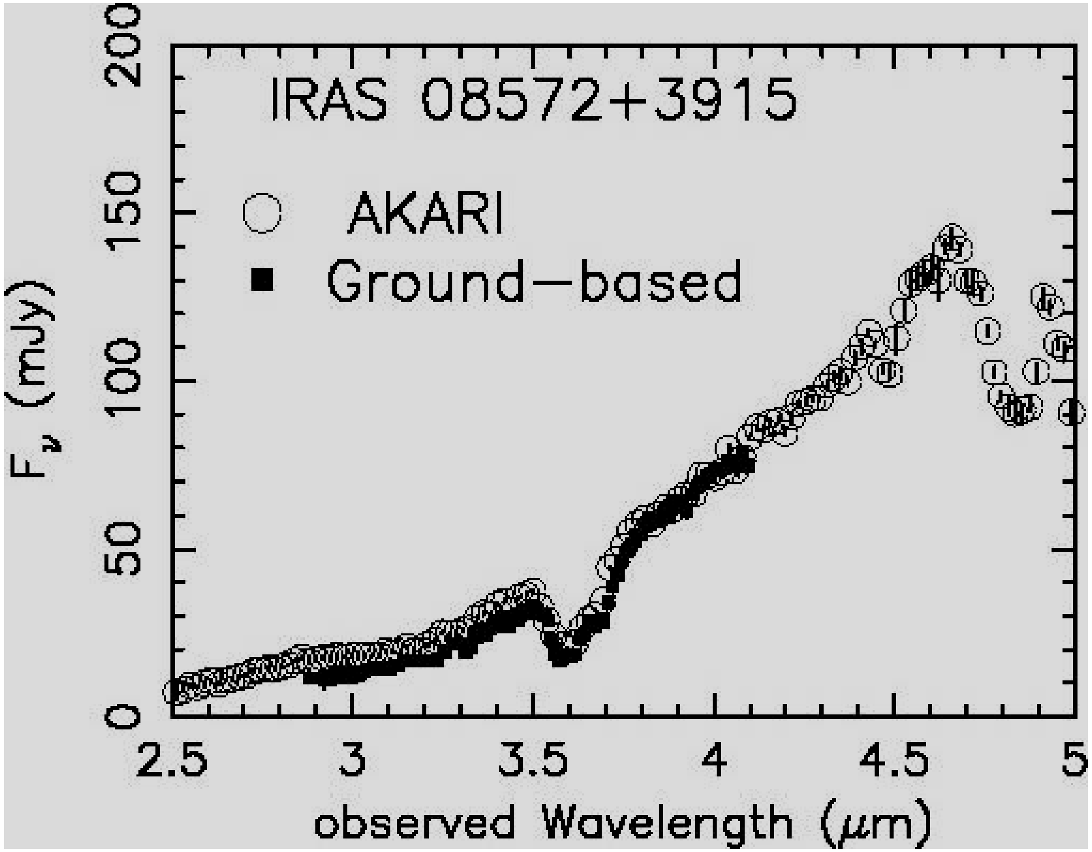
cool FIR color

cool ≠ starburst

AKARI

2.5-5 um spectroscopy → **z > 0.15 ULIRG**

Unaffected by Earth's atmosphere



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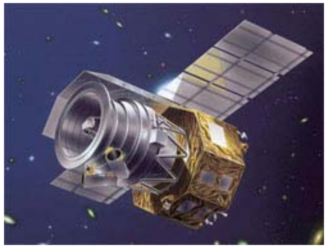
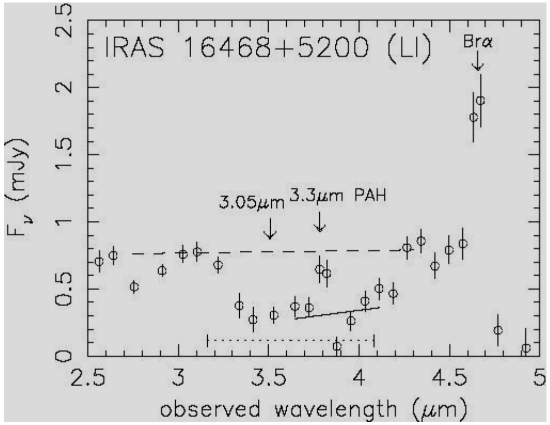
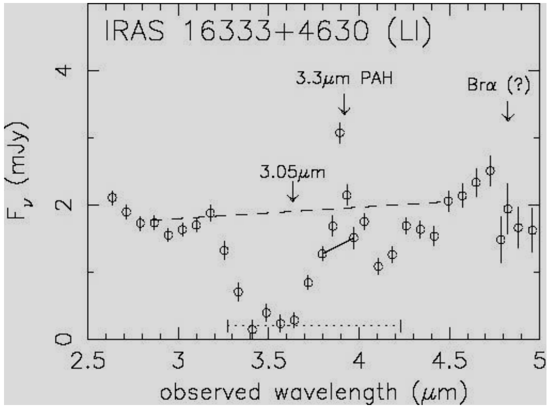


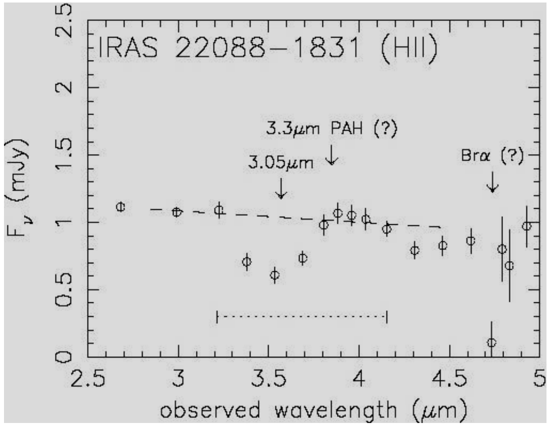
z>0.15 → Higher L_{IR} ULIRGs

↓

Buried AGNs increase with L_{IR}







↓

AGN-feedback for galaxy down-sizing ?

Summary

1. **Buried AGNs : 30-50% non-Sy ULIRGs**

warm & cool

2. **Nuclear dust amount:**
   non-Sy ULIRGs > Sy2 ULIRGs

→ **Optical Sy (non-)detectability depends on the amount of nuclear dust**

Imanishi et al. 2006 ApJ 637 114 (Subaru)
Imanishi et al. 2007 ApJS 171 72 (Spitzer)
Imanishi et al. 2008 PASJ submitted (AKARI)